Binary black hole remnants of first stars for the gravitational wave source

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GW150914

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Observation of Gravitational Waves from a Binary Black Hole Merger

B. P. Abbott *et al.**
(LIGO Scientific Collaboration and Virgo Collaboration)
(Received 21 January 2016; published 11 February 2016)

On September 14, 2015 at 09:50:45 UTC the two detectors of the Laser Interferometer Gravitational-Wave Observatory simultaneously observed a transient gravitational-wave signal. The signal sweeps upwards in frequency from 35 to 250 Hz with a peak gravitational-wave strain of 1.0×10^{-21}. It matches the waveform predicted by general relativity for the inspiral and merger of a pair of black holes and the ringdown of the resulting single black hole. The signal was observed with a matched-filter signal-to-noise ratio of 24 and a false alarm rate estimated to be less than 1 event per 203 000 years, equivalent to a significance greater than 5.1σ. The source lies at a luminosity distance of 410^{+160}_{-180} Mpc corresponding to a redshift $z = 0.09^{+0.03}_{-0.04}$. In the source frame, the initial black hole masses are $36^{+5}_{-4} M_\odot$ and $29^{+4}_{-4} M_\odot$, and the final black hole mass is $62^{+4}_{-4} M_\odot$, with $3.0^{+0.5}_{-0.5} M_\odot c^2$ radiated in gravitational waves. All uncertainties define 90% credible intervals. These observations demonstrate the existence of binary stellar-mass black hole systems. This is the first direct detection of gravitational waves and the first observation of a binary black hole merger.

- $36M_\odot + 29M_\odot$, circular orbit

GW150914

- More than factor 2-3 larger mass of BH compared with that in X-ray binary
- Many theories exist such as
- 1)Pop II BBH
- 2)Pop III BBH

Low metal field binaries

- 3)Primordial Binary BH (PBBH)
- 4)Three body origin from Globular Cluster
- 5)Fragmentation of very massive stars
- ……………………..

Why field binaries?

- There are many massive close binaries

  Example

  Milky way young open clusters

  71 O stars   fbinary=69+/-9% (P<3200days)

  Sana et al. 2012

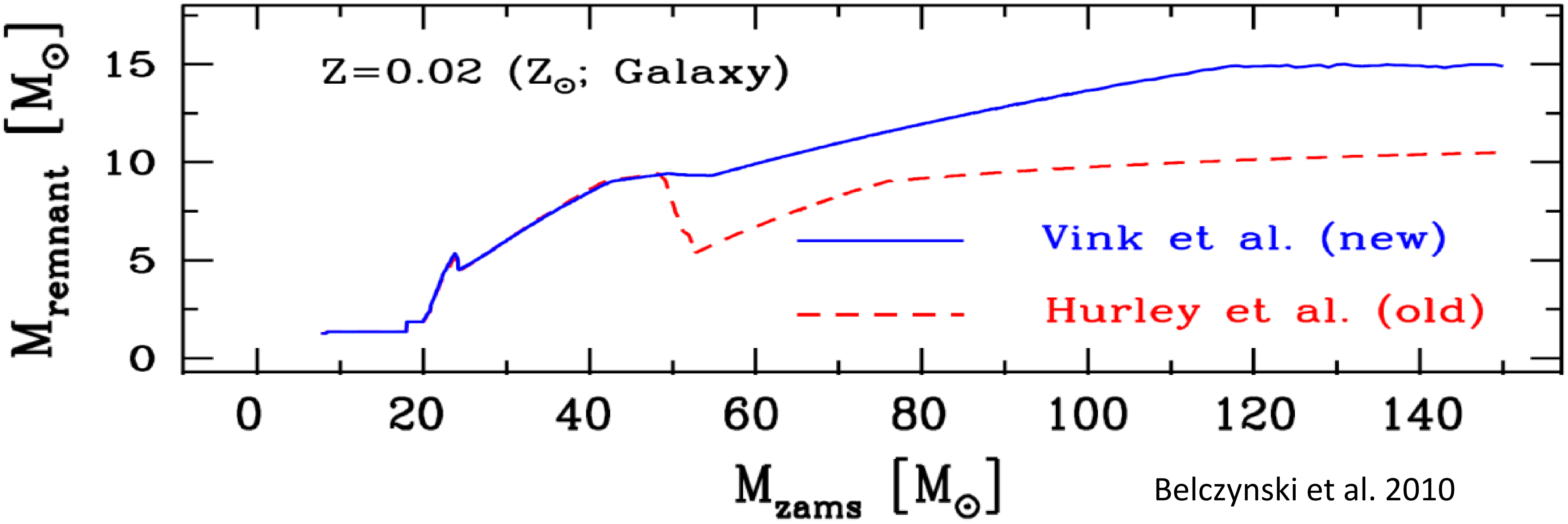
  30 Doradus (Tarantula Nebula)

  362 O stars fbinary=51+/-4%(P<3200days)

  Sana et al. 2013

Why low metal?

- If the progenitor of BH is Pop I (=Solar metal stars)



Belczynski et al. 2010

- The orbit become wide due to wind mass loss

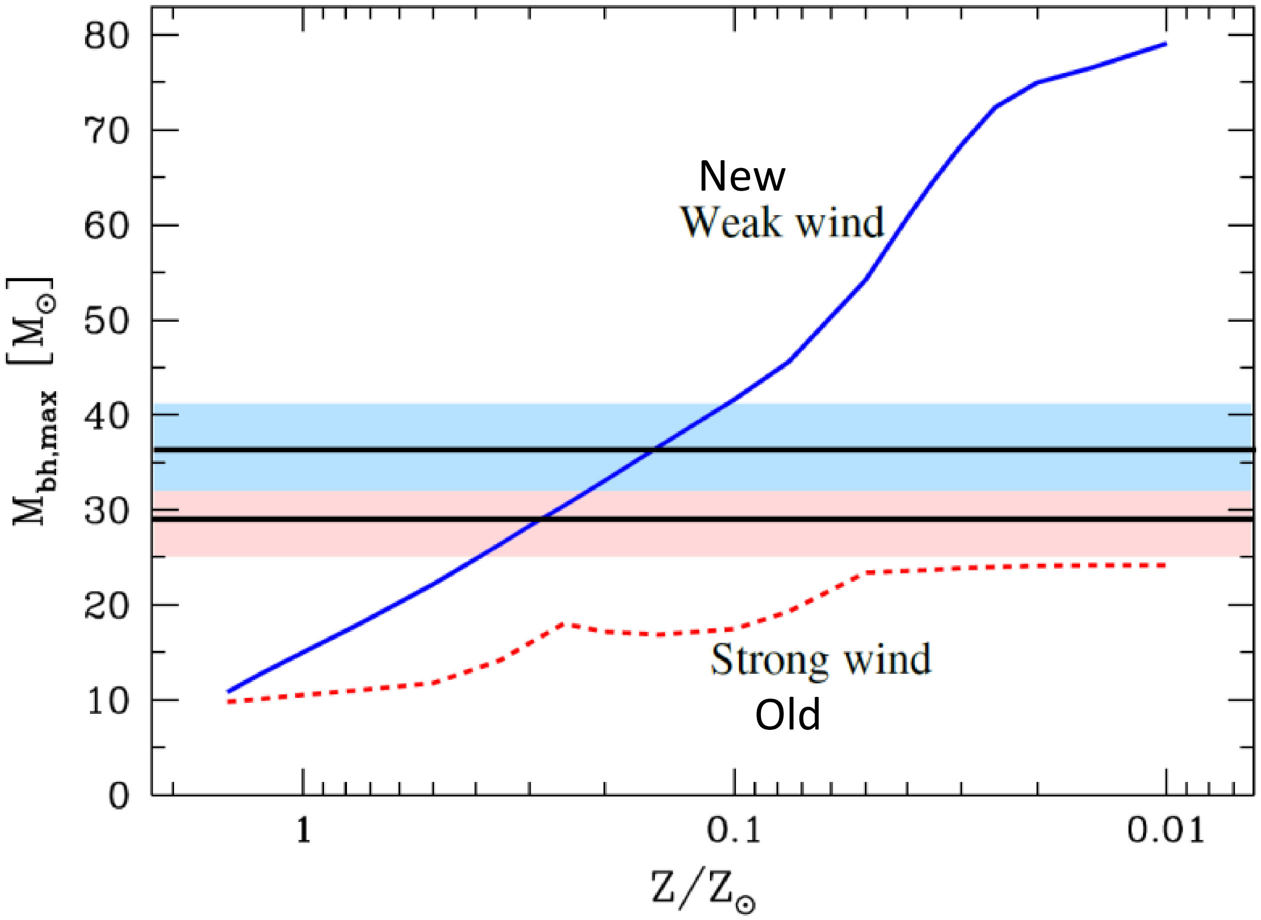
Why low metal?

- If the progenitor is low metal,

- Pop II (Z<0.1Zsun)

  Typical mass is same as Pop I

  But, week wind mass loss



Minitial: 8Msun<M<150Msun

Single stellar evolution

with 2 stellar wind models.

(Belczynski et al.2010,

Abbot et al.2016)

- Pop III (No metal)

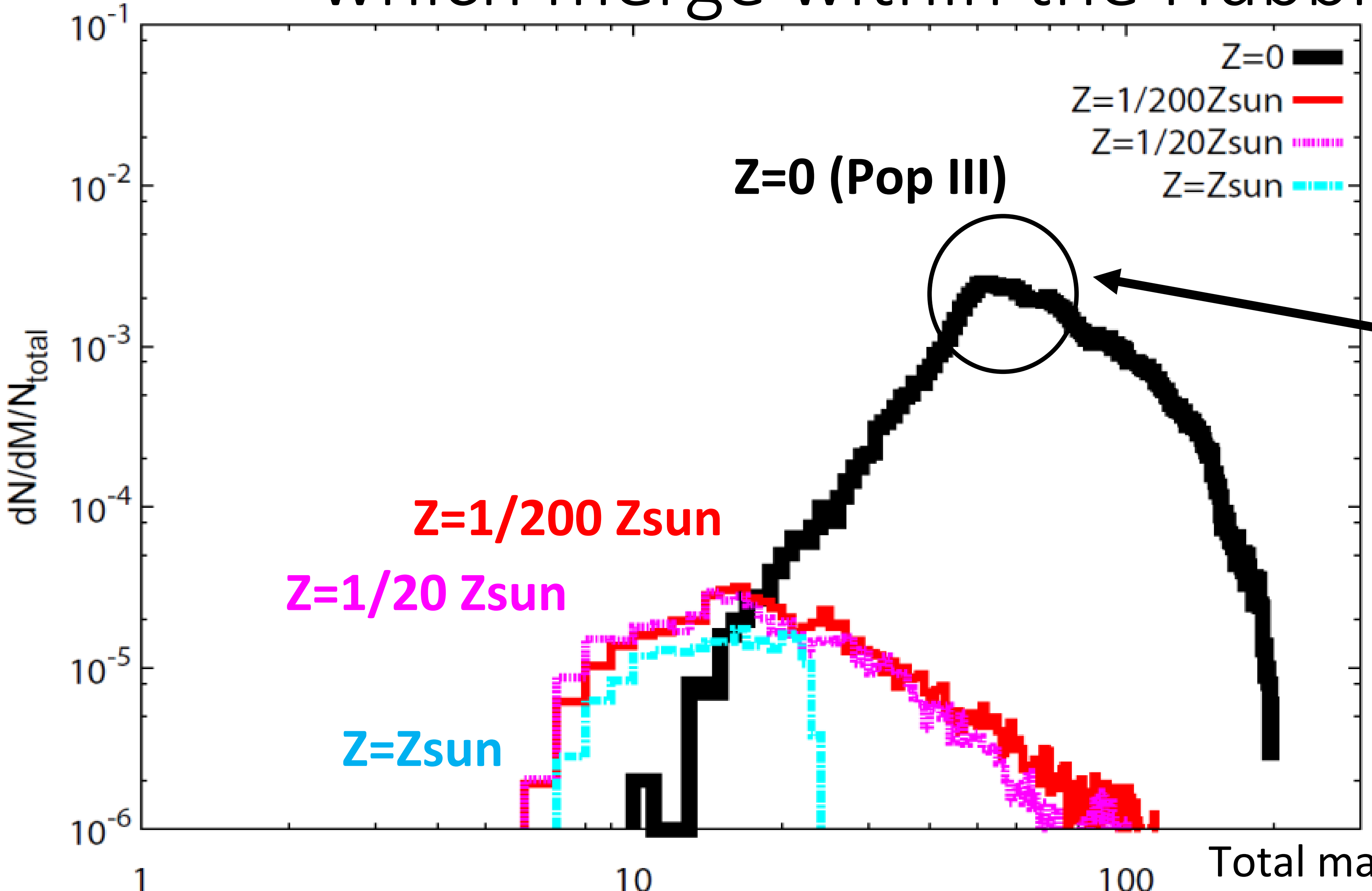
  Pop III stars are ***the first stars*** after the Big Bang.

  Typical mass is more massive than Pop I, II

  MpopIII~10-100Msun

  No wind mass loss due to no metal.

Total mass distribution of BBH which merge within the Hubble time



Typical total mass

M～60 $M_\odot$

(30 $M_\odot$ +30 $M_\odot$)

TK et al. 2014,2016

IMF:Flat

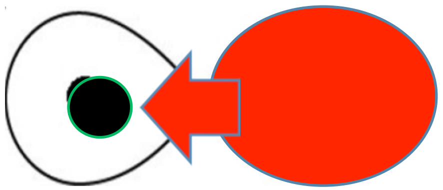
(10M<M<140M)

e.g. Pop I, Pop II
(Z=0.02,0.001,0.0001)
IMF:Salpeter
(1Msun<M<140Msun)
Typical mass ～10 $M_\odot$

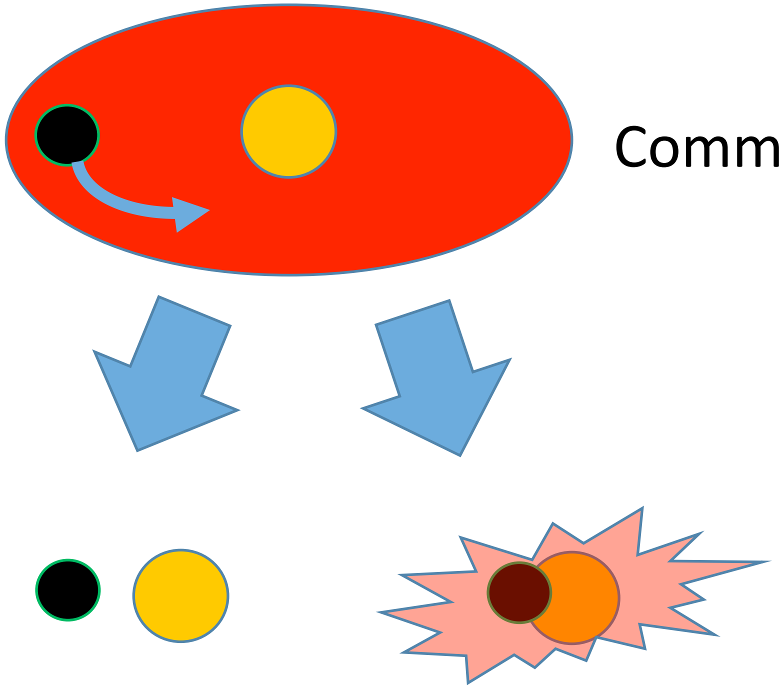
What do determine the BH-BH mass?

- Steller wind mass loss
- Binary interactions

(Mass transfer, Common envelope)



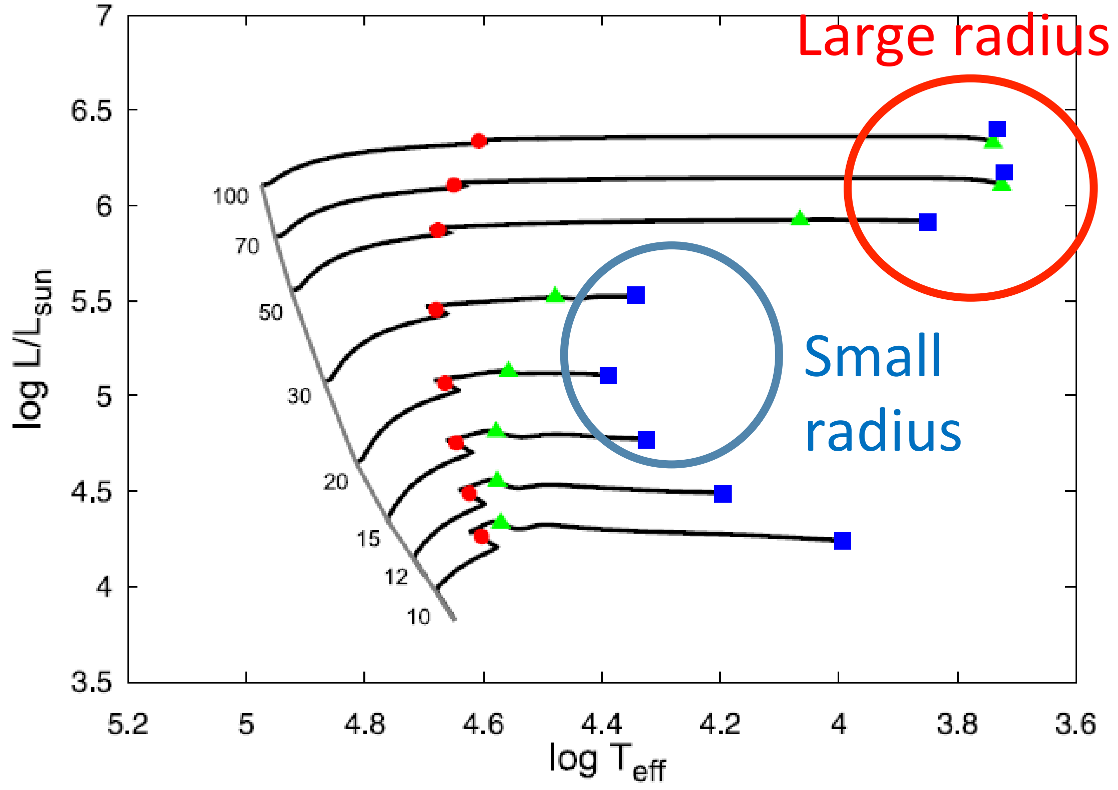
Mass transfer



Common envelope

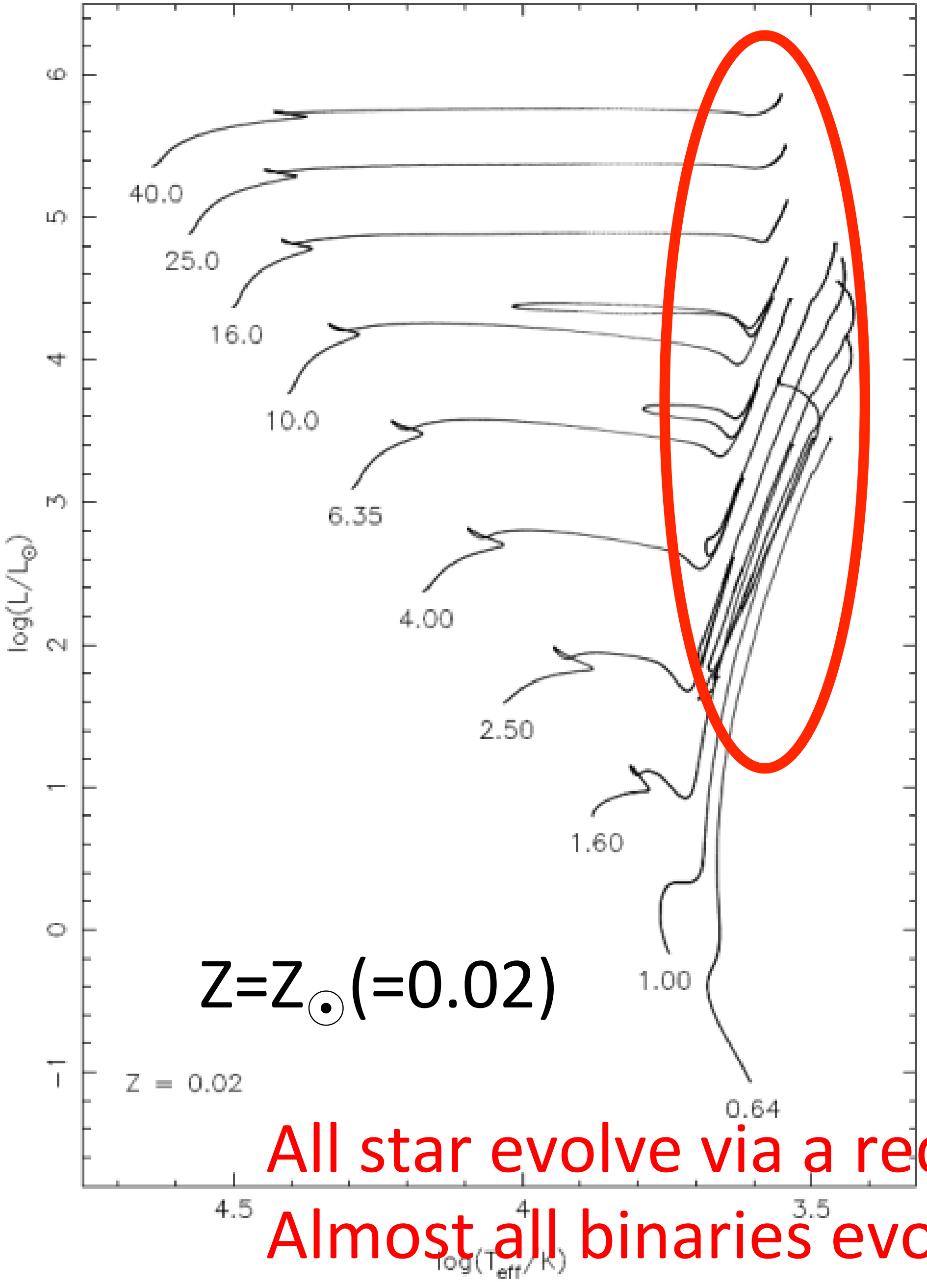
Close binary or merge

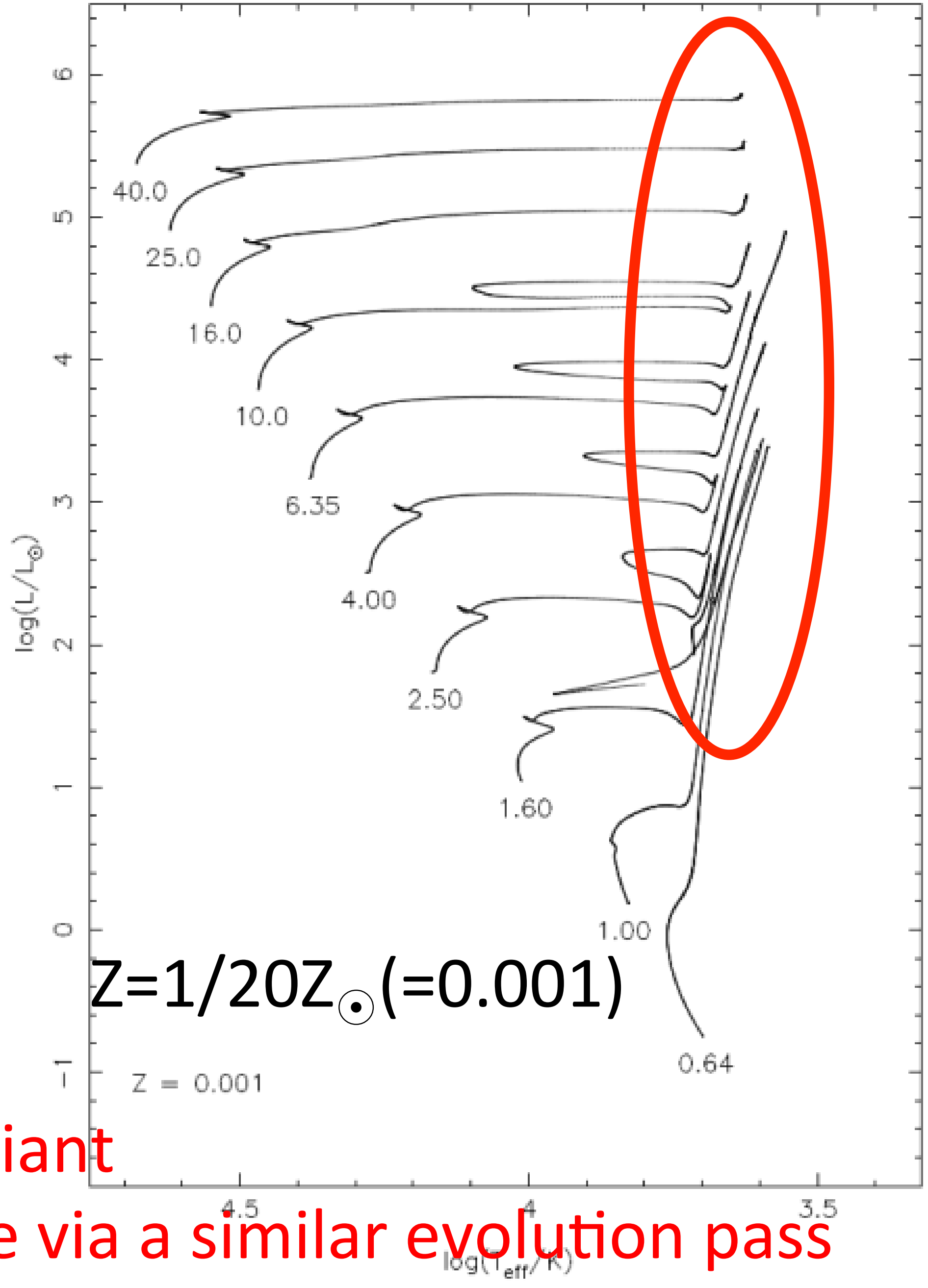
Why Pop III binaries become 30Msun BH-BH



- M>50Msun red giant
  →Mass transfer is unstable
  →common envelope
  →1/3~1/2 of initial mass (~25-30Msun)

- M<50Msun blue giant
  →Mass transfer is stable
  →mass loss is not so effective
  →2/3~1 of initial mass (25-30Msun)





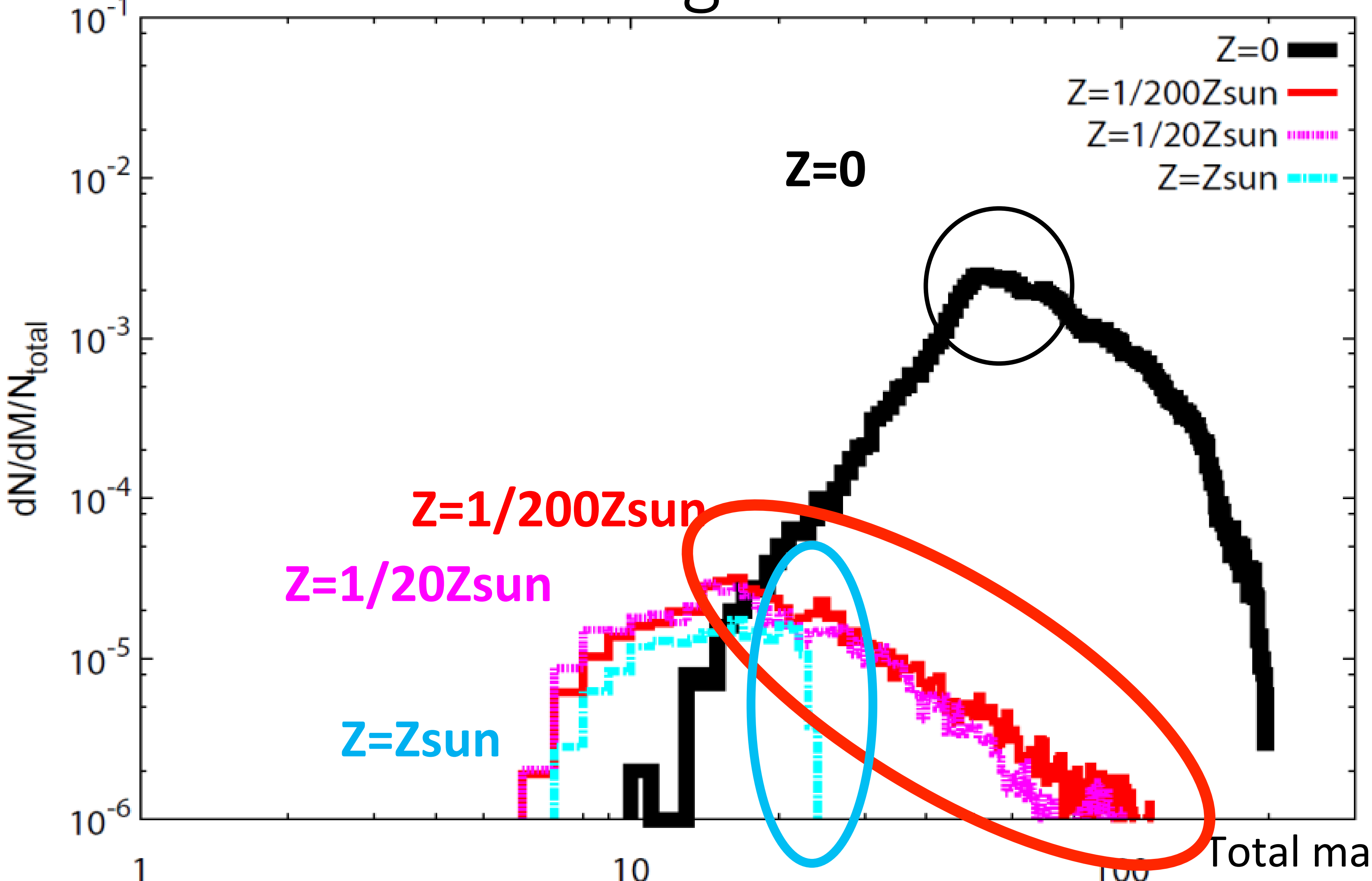
All star evolve via a red giant

Almost all binaries evolve via a similar evolution pass

Figure 1. Selected OVS evolution tracks for $Z = 0.02$, for masses 0.64, 1.0, 1.6, 2.5, 4.0, 6.35, 10, 16, 25 and 40 $M_\odot$.

Figure 2. Same as Fig. 1 for $Z = 0.001$. The 1.0 $M_\odot$ post He flash track has been omitted for clarity.

Total mass distribution of BBH which merge within the Hubble time



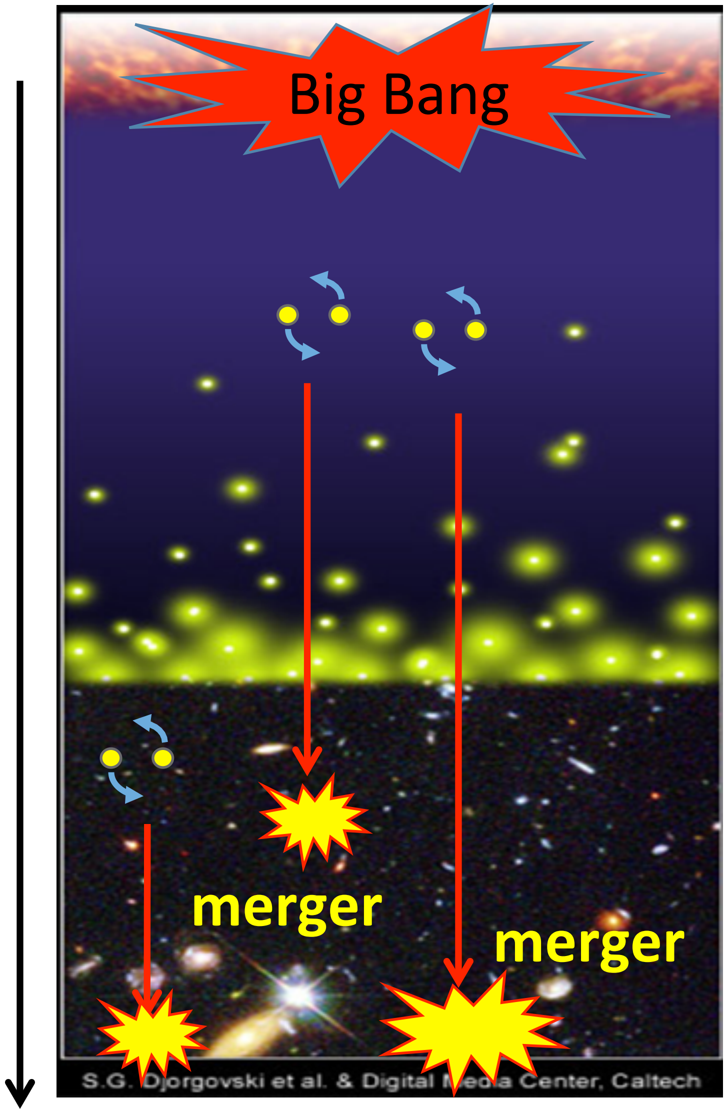
This shape reflects the influence of Pop III stellar evolution

These shapes have the influence of IMF

and the influence of stellar wind mass loss

Pop III BBH remnants for gravitational wave

- Pop III stars were born and died at z~10
- The typical merger time of compact binaries ~$10^{8\text{-}10}$yr
- We might see Pop III BBH at the present day.



time

Djorgovski et al.&Degital Media Center

Pop III BBH?

ASTROPHYSICAL IMPLICATIONS OF THE BINARY BLACK-HOLE MERGER GW150914

ApJL Abbot. et al 2016

2014; Dominik et al. 2015).

On the extreme low-metallicity end, it has been proposed that BBH formation is also possible in the case of stellar binaries at zero metallicity (Population III [PopIII] stars; see Belczynski et al. 2004; Kinugawa et al. 2014). The predictions from these studies are even more uncertain, since we have no observational constraints on the properties of first-generation stellar binaries (e.g., mass function, mass ratios, orbital separations). However, if one assumes that the properties of PopIII massive binaries are not very different from binary populations in the local universe (admittedly a considerable extrapolation), then recently predicted BBH total masses agree astonishingly well with GW150914 and can have sufficiently long merger times to occur in the nearby universe (Kinugawa et al. 2014). This is in contrast to the predicted mass properties

The star formation rate of Pop III

In order to calculate merger rate,

we need to know

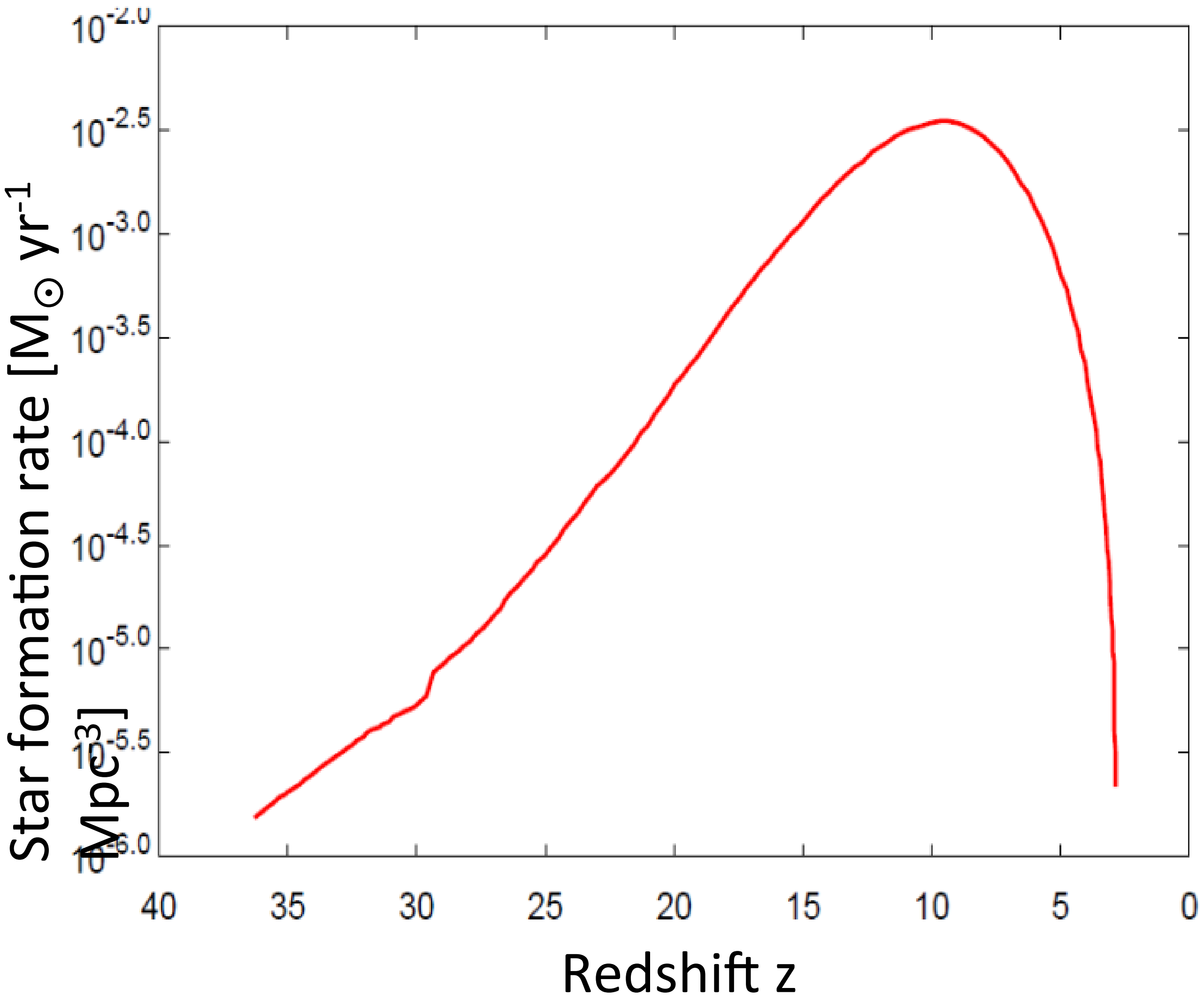
- When were Pop III stars born?
- How many Pop III stars were born?

⇒Star formation rate

We adopt the Pop III SFR

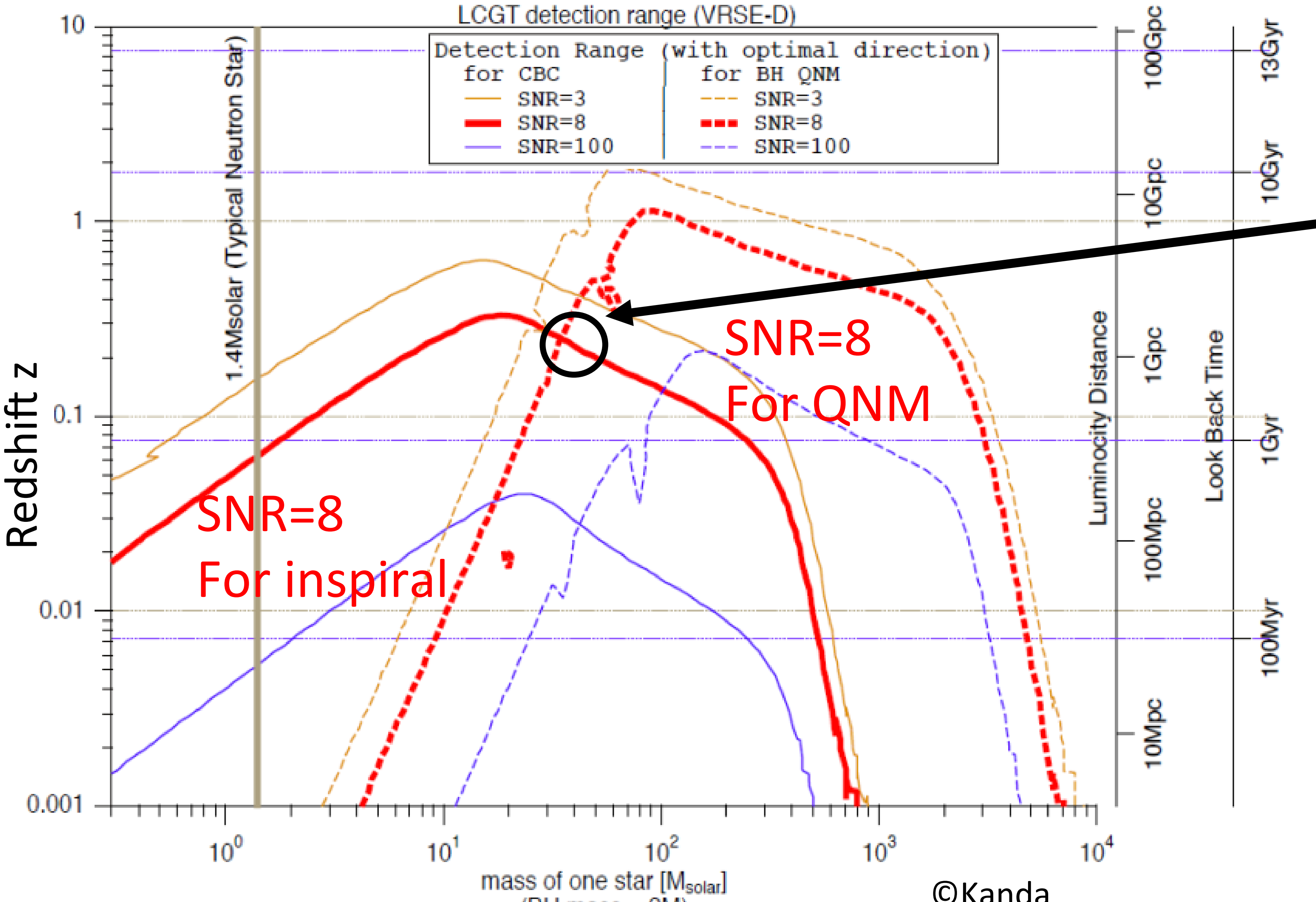
by de Souza et al. 2011

$SFR_{peak} \sim 10^{-2.5}$ [$M_\odot$ yr^{-1} Mpc^{-3}]



(de Souza et al. 2011)

Detection range of KAGRA and Adv. LIGO



$M_{BH} \sim 30 M_{\odot}$
SNR=8

Luminocity distance
$\sim$1.5 Gpc

Redshift z$\sim$0.28

©Kanda

Detection rate of Pop III BH-BH

- **Detection rate of Pop III BBH (GW150914 like BBH) in our standard model**

  R～180 $\left(SFR_{peak}/10^{-2.5}\right)\left(f_b/(1+f_b)/0.33\right)$ **[yr^{-1}]**(S/N>8)

- Typical mass

  M～30 M$_{\odot}$ →We can see the QNM of merged BBH

  We might detect the Pop III BBH by GW

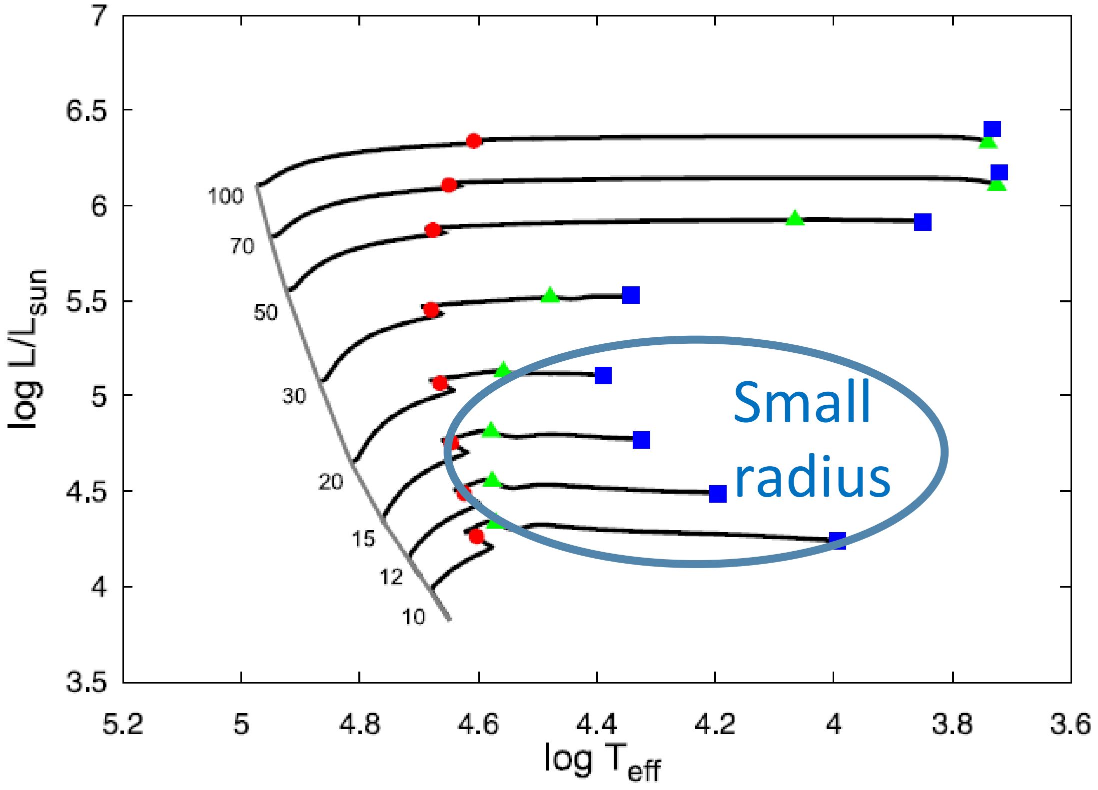
Other Pop III compact binaries cases

- Pop III NSNS

  Almost all binary NS disrupt

- Pop III NSBH

Pop III NS progenitor evolution



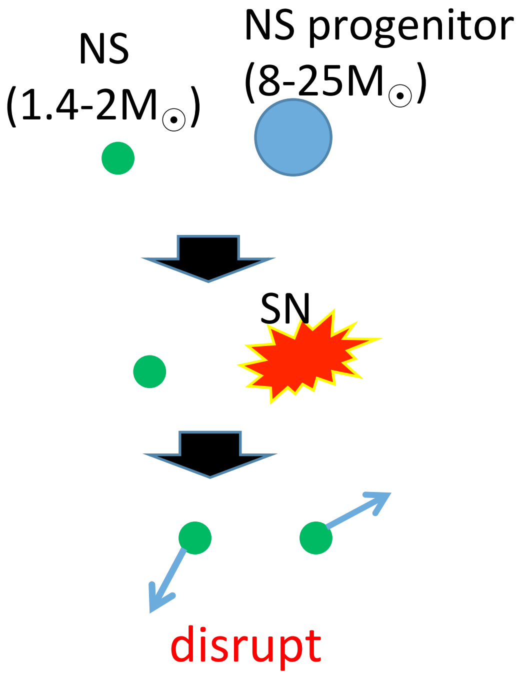
- blue giant

→Mass transfer is stable

→mass loss is not so effective before supernova

Pop III NS-NS disrupt

For example, we consider NS and NS progenitor binary.



In the case of Pop III NS progenitor, wind mass loss and the mass loss due to binary interaction is not effective.
When NS progenitor becomes supernova, NS progenitor suddenly loses mass and becomes NS.
Then, due to instant mass loss the binding energy of binary decreases and binary NS disrupts.

⇨ Binary NS cannot survive!

Other Pop III compact binaries cases

- Pop III NSNS

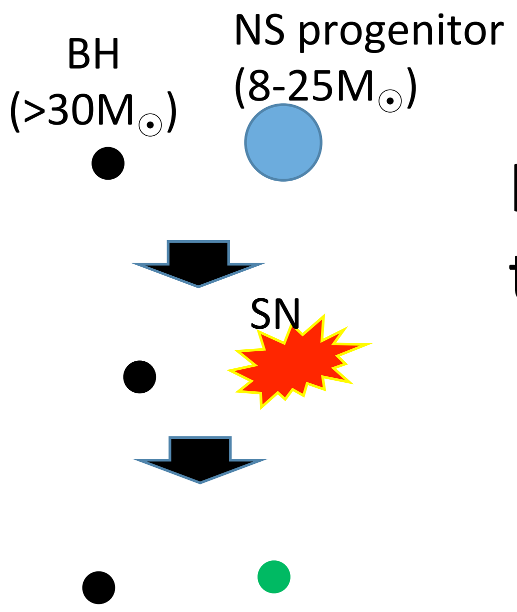
  Almost all binary NS disrupt

- Pop III NSBH

  NSBH do not disrupt

Pop III NS-BH do not disrupt

For example, we consider BH and NS progenitor binary.



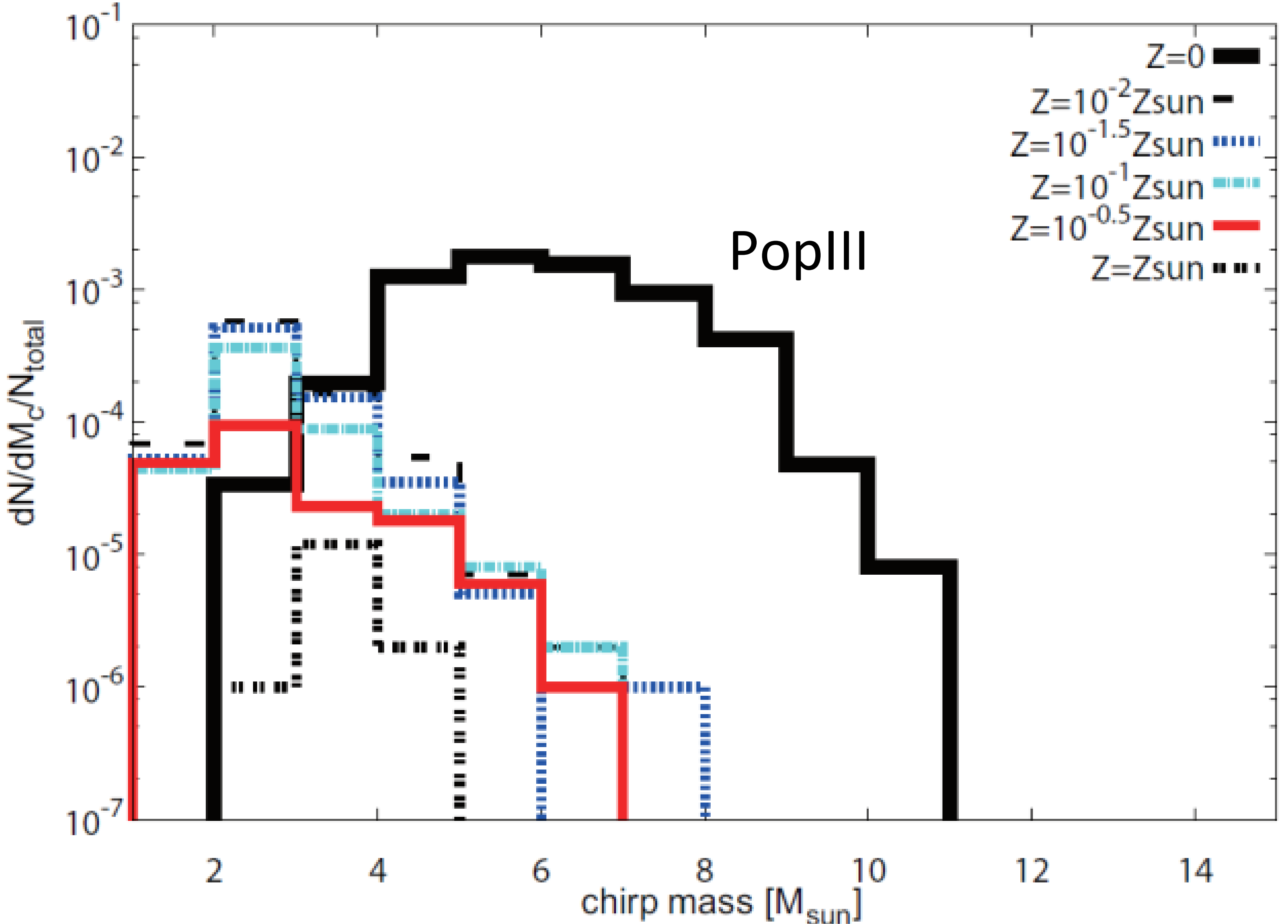
In the case of Pop III NS progenitor, wind mass loss and the mass loss due to binary interaction is not effective.

When NS progenitor becomes supernova, NS progenitor suddenly loses mass and becomes NS.

But, due to massive BH, NS do not disrupts.

⇨ NS BH can survive!

Merging NSBH chirp mass distribution



- **Typical chirp mass**

  M～6 $M_\odot$

  (1.4Msun NS+50Msun BH)

NSBH detection rate

	Merger rate [/yr/Gpc^3]	aLIGO O2 detection rate [/yr]	aLIGO (design sensitivity) detection rate [/yr]
Pop I+II	28.8 (Belczynski et al. 2016)	1.41 (Belczynski et al. 2016)	~10
Pop III	1.25	0.658 (*)	5.24(*)

*For simplicity, as the assumption of the chirp mass of Pop III NSBH, we fixed Mc = $6M_\odot$ (Kinugawa et al.2016)

Summery

- **Detection rate of Pop III BBH (GW150914 like BBH)**

  R～180 $\left(SFR_{peak}/10^{-2.5} \right) \left(f_b/(1+f_b)/0.33 \right)$ **[yr^{-1}]**(S/N>8)

- Typical chirp mass

  M～30 M$_\odot$

  We might detect (detected?) the Pop III BBH by GW

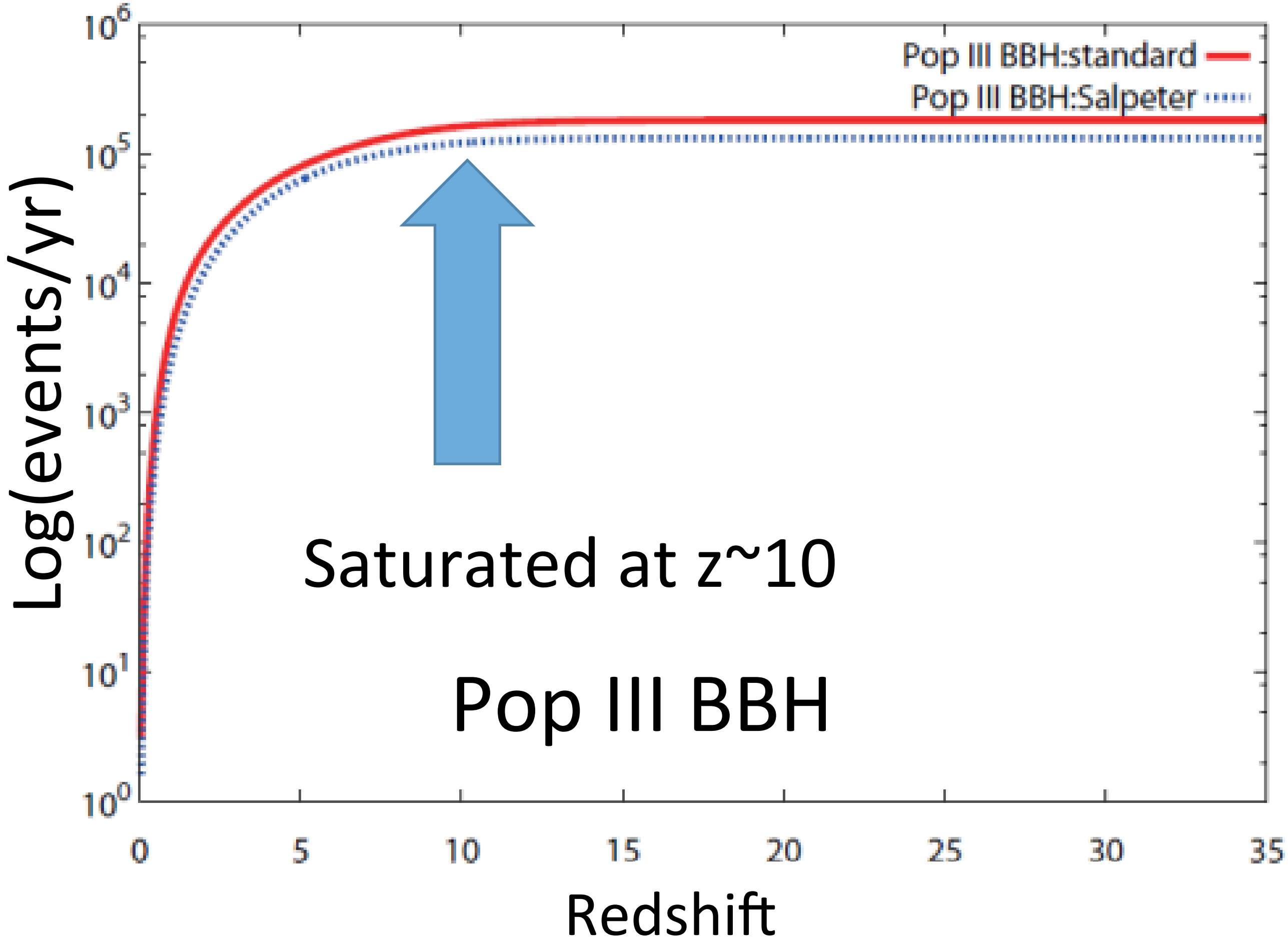
- **Detection rate of Pop III NSBH**

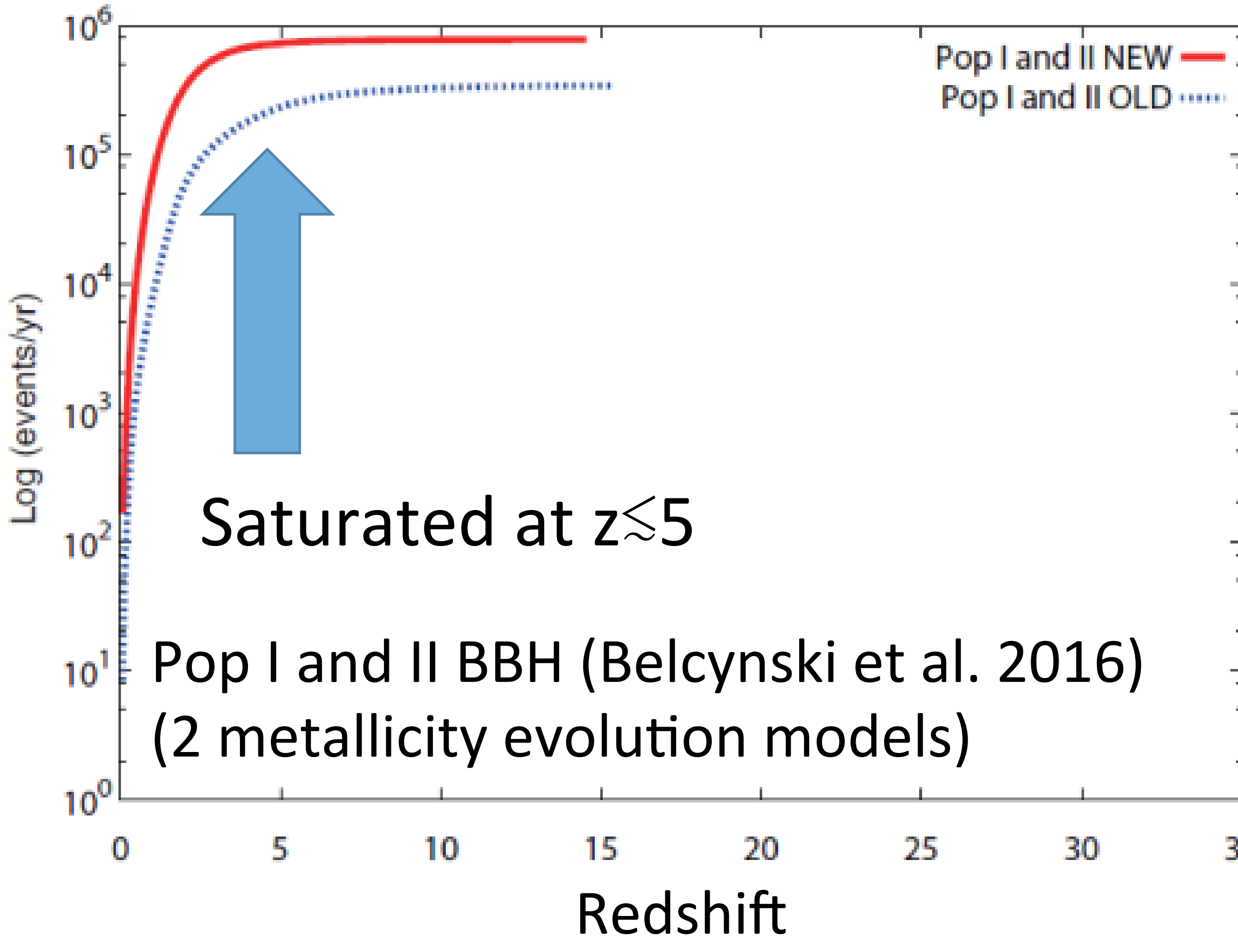
  R～5 $\left(SFR_{peak}/10^{-2.5} \right) \left(f_b/(1+f_b)/0.33 \right)$ **[yr^{-1}]**(S/N>8)

- **Typical chirp mass**

  M～6 M$_\odot$ (1.4Msun NS+50Msun BH)

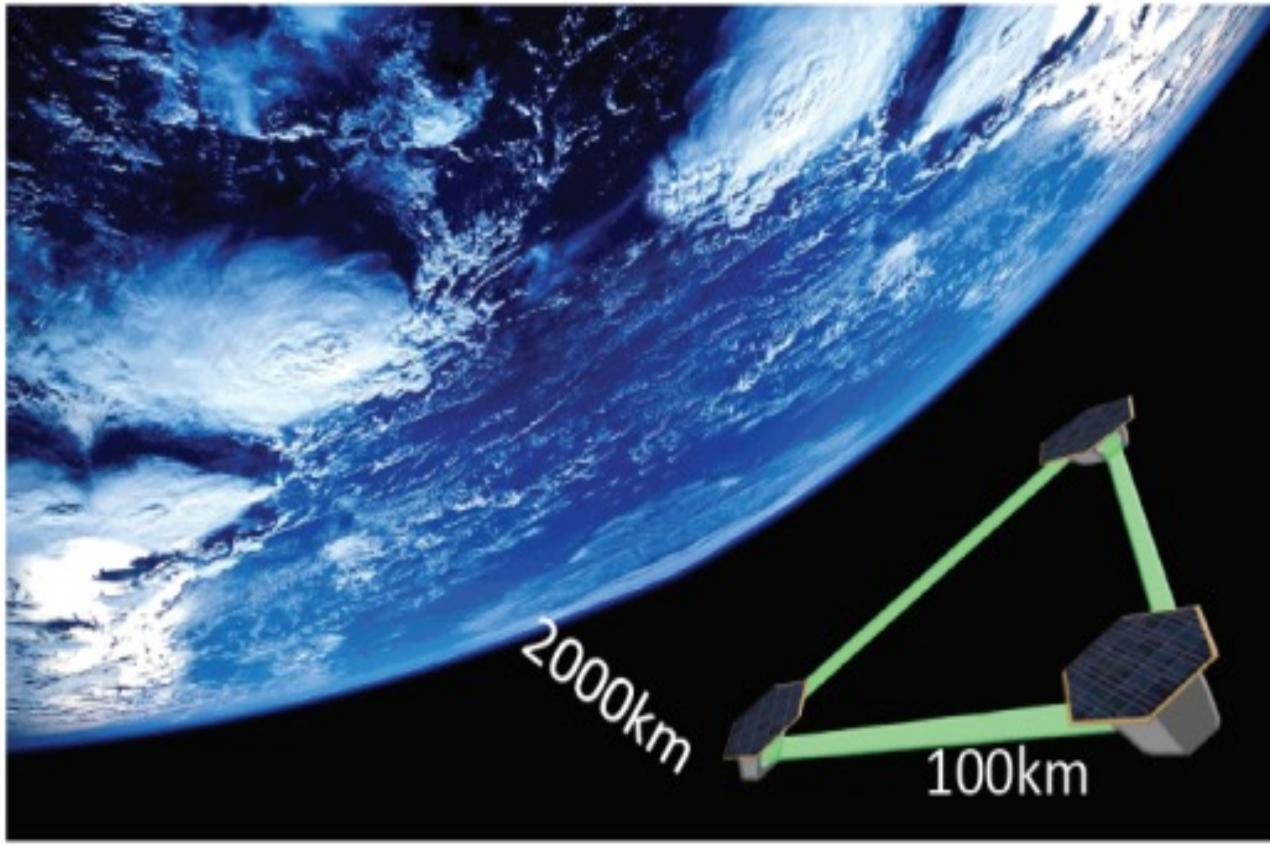
Cumulative BBH merger rate





Future plan of GW observer : pre-DECIGO and DECIGO

- DECIGO: Japanese space gravitational wave observatory project
- Pre-DECIGO: test version of DECIGO



- Pre-DECIGO : z~10 (30 Msun BH-BH)
  ~10^5 events/yr
- DECIGO can see Pop III BH-BHs
  when Pop III stars were born!
  (Nakamura, Ando, Kinugawa et al. 2016)

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Summary

- Pop III binaries tend to become 30Msun+30Msun BH-BH
- **Pop III BBH detection rate of aLIGO in our standard model**

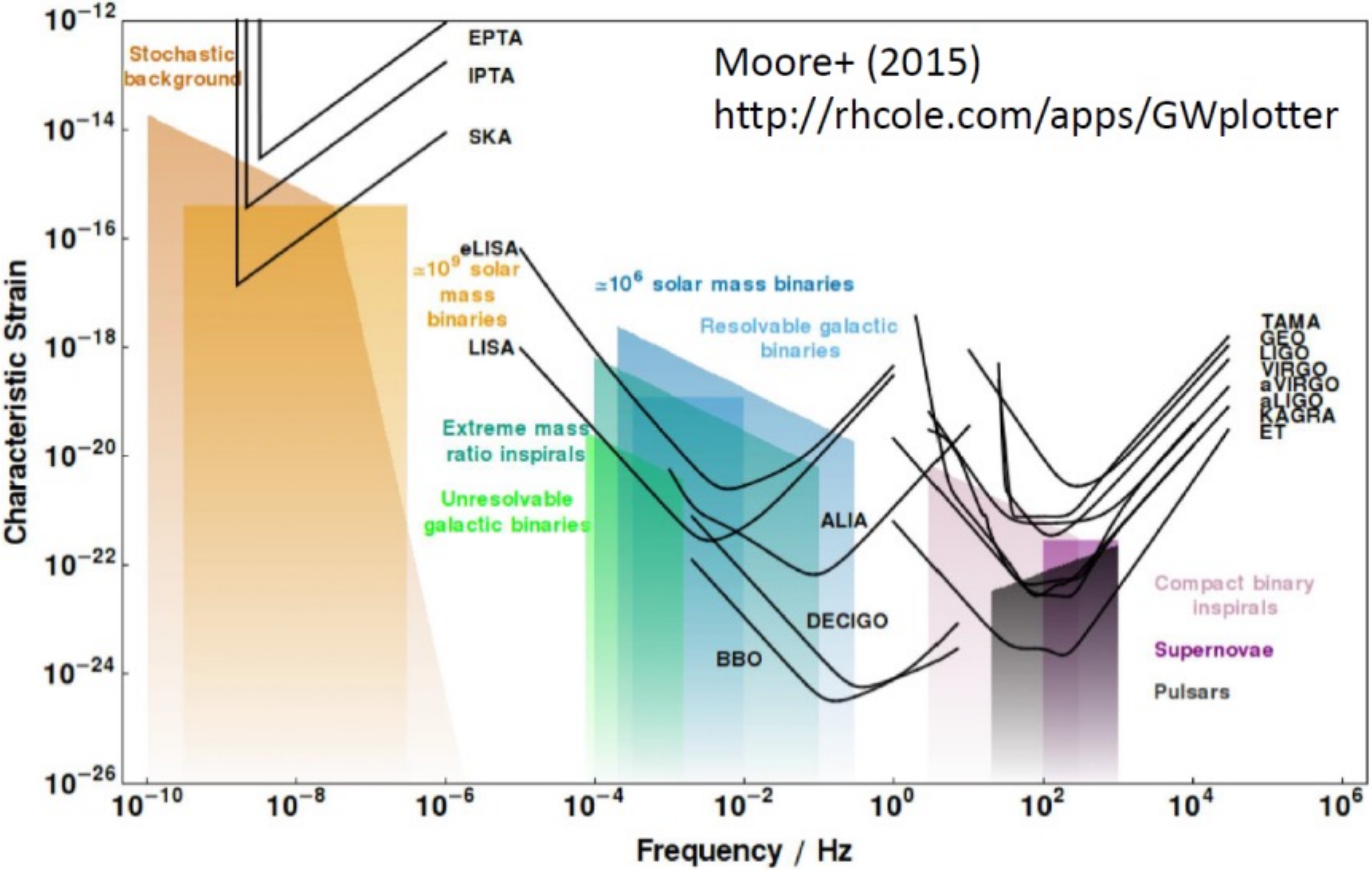
  $$R \sim 180 \left(SFR_{peak}/10^{-2.5} \right) \left(f_b/(1+f_b)/0.33 \right) [\mathrm{yr}^{-1}] (S/N>8)$$

- The mass distribution or the redshift dependence might distinguish Pop III from Pop I,II.
- DECIGO can see Pop III BH-BH merger when they were born

Appendix

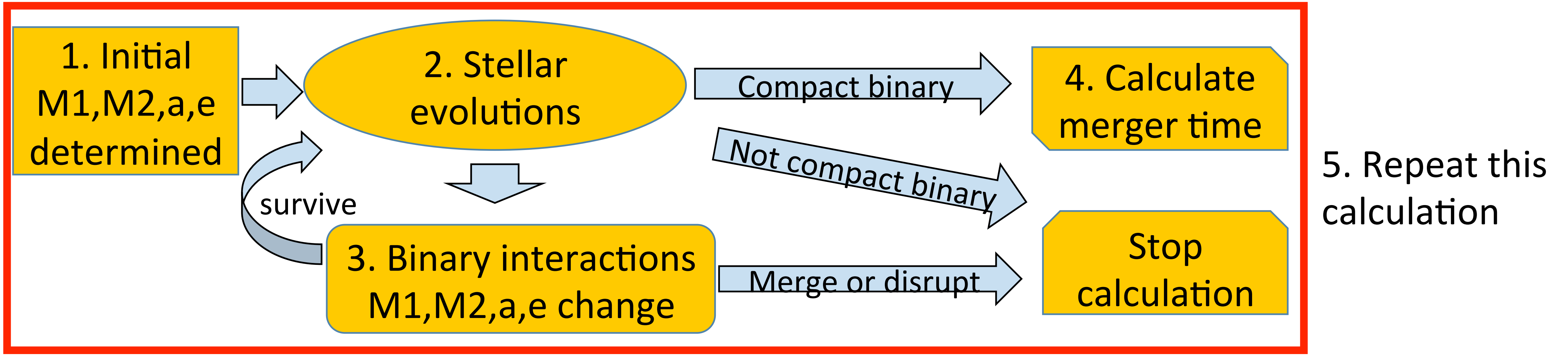
Pop I and Pop II case (Dominik et al. 2015)

- From 1/200 Zsun to 1.5 Zsun
- BH-BH detection rate (Their standard model) ~300/yr
- 25% of above rate is >20 Msun BHBH
- Thus, Detection rate of high mass BHBH ~80/yr



Moore+ (2015)
http://rhcole.com/apps/GWplotter
Characteristic Strain
Frequency / Hz
Stochastic background
EPTA
IPTA
SKA
≥10^9 solar mass binaries
eLISA
LISA
≥10^6 solar mass binaries
Resolvable galactic binaries
Extreme mass ratio inspirals
Unresolvable galactic binaries
ALIA
DECIGO
BBO
TAMA
GEO
LIGO
VIRGO
aVIRGO
aLIGO
KAGRA
ET
Compact binary inspirals
Supernovae
Pulsars

How to calculate Pop III binaries?



1. Initial stellar parameters are decided by Monte Carlo method with initial distribution functions

   (primary mass: M1, secondary mass: M2, separation: a, orbital eccentricity: e)

2. We calculate evolution of stars

3. If star fulfills the condition of binary interactions (BIs), we calculate BIs and change M1, M2, a, e .

   ･If binary merges or disrupts due to BIs before binary becomes compact binary, we stop calculation.

   ･If binary survives from BIs, we calculate stellar evolutions again.

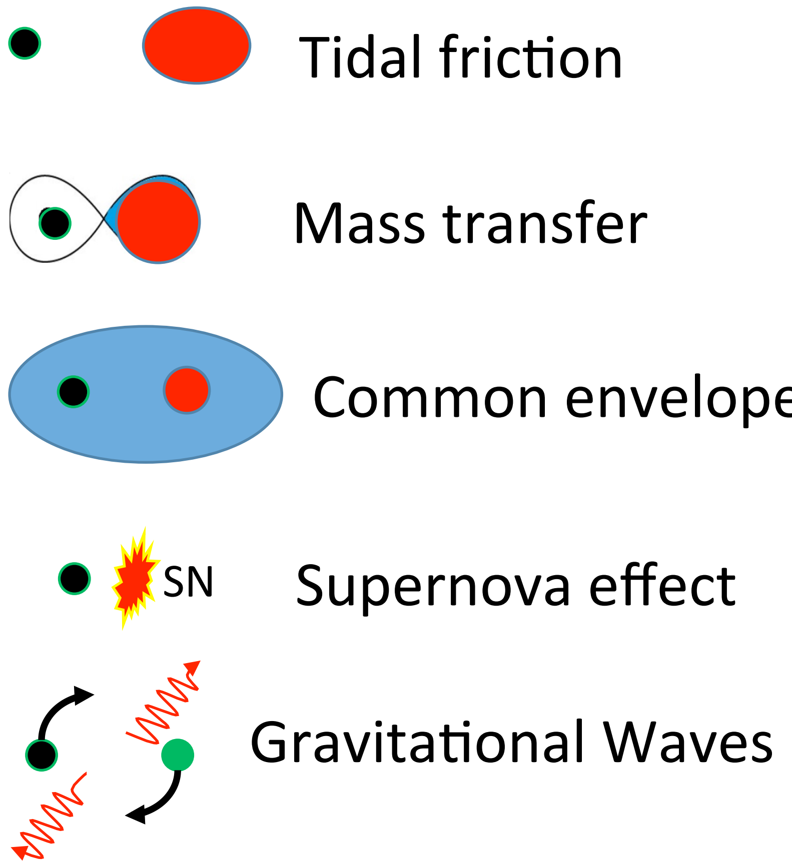
4.If binary becomes compact binary (NS-NS, NS-BH, BH-BH), we calculate when binary merge due to GW.

5.We repeat these calculations and take the statistics of compact binary mergers.

Binary Interactions

- Tidal friction
- Mass transfer
- Common envelope
- Supernova effect
- Gravitational radiation

Change
M1,M2,a, e



We need to specify some parameters to calculate these effects.

We use the parameters adopted for Pop I population synthesis in Our standard model.

Pop III binary population synthesis

We simulate 10^6 Pop III-binary evolutions and estimate how many binaries become compact binary which merges within Hubble time.
×84 models (Kinugawa et al.2016)

Initial stellar parameters are decided by Monte Carlo method with initial distribution functions

- Initial parameter (M1,M2,a,e) distribution in our standard model

M1 : Flat ($10\ M_{\odot} < M < 100\ M_{\odot}$)

q=M2/M1 : P(q)=const. ($0<q<1$)

a : $P(a) \propto 1/a$ ($a_{min} < a < 10^6 R_{\odot}$)

e : $P(e) \propto e$ ($0<e<1$)

} The same distribution functions adopted for Pop I population synthesis

Results

The numbers of the compact binaries which merge within Hubble time for 10^6 binaries

NSNS	NSBH	BHBH
0	50	115056

Our standard model

- A lot of Pop III BH-BH binaries form and merge within Hubble time
- Close NS binaries do not form

The star formation rate of Pop III

In order to calculate merger rate,

we need to know

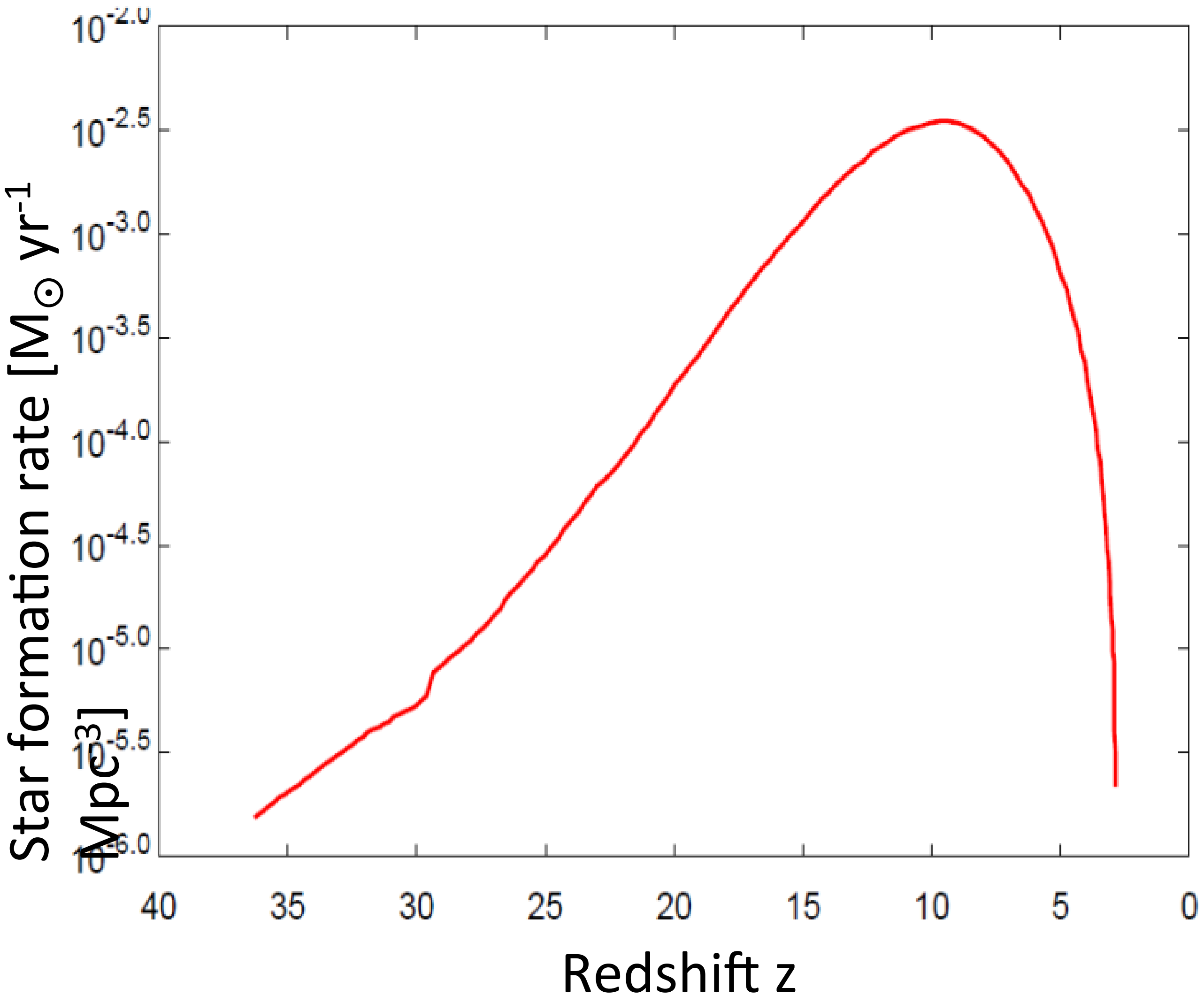
- When were Pop III stars born?
- How many Pop III stars were born?

⇒Star formation rate

We adopt the Pop III SFR

by de Souza et al. 2011

$SFR_{peak} \sim 10^{-2.5}$ [$M_{\odot}$ yr^{-1} Mpc^{-3}]



(de Souza et al. 2011)

Consistency with LIGOS6 and Adv.LIGO

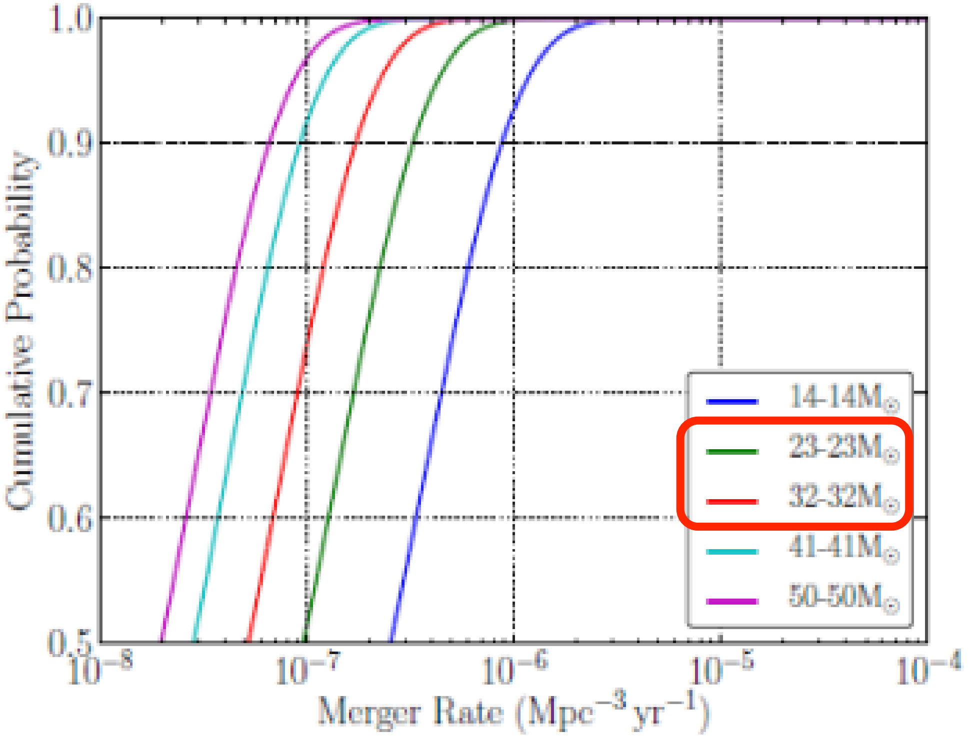


FIG. 6: Cumulative posterior probabilities over astrophysical merger rate, for the bins shown in Figure 5 with central values $m_1 = m_2 = 50, 41, 32, 23, 14\, M_{\odot}$ (left to right). We show the probability level corresponding to the 90% confidence rate limit (dashed horizontal line). These posteriors were evaluated for signals described by the EOBNRv2 waveform family in S6 data using S5 search results as prior information.

Aasi, Abadie, Abbott et al. (2013)

- **LIGOS6 upper limit of BH-BH merger rate**

  left figure

  $\sim 10^{-7}\ yr^{-1}Mpc^{-3}$

- **Merger rate estimated by GW150914 (z<0.5)**

  $\sim 0.02\text{-}4\times10^{-7}\ yr^{-1}Mpc^{-3}$

- **Pop III BH-BH Merger rate at z~0**

  $\mathbf{R \sim 2.5\times10^{-8}\ (SFR_{peak}/10^{-2.5})\ Err_{sys}\ [yr^{-1}\ Mpc^{-3}]}$

<u>Our result is consistent with LIGO</u>

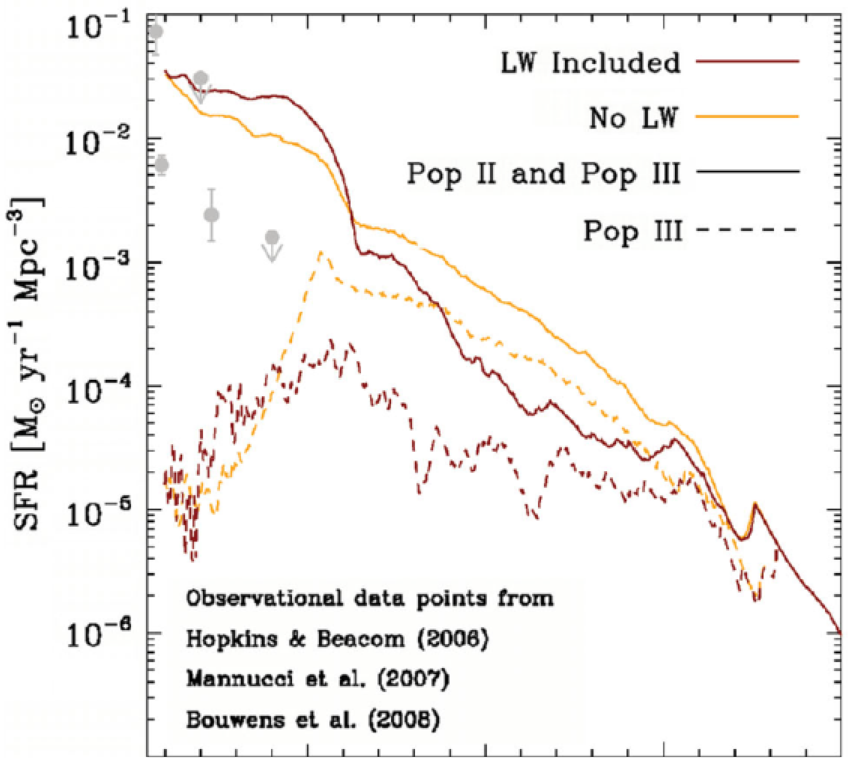
Errsys (Example)

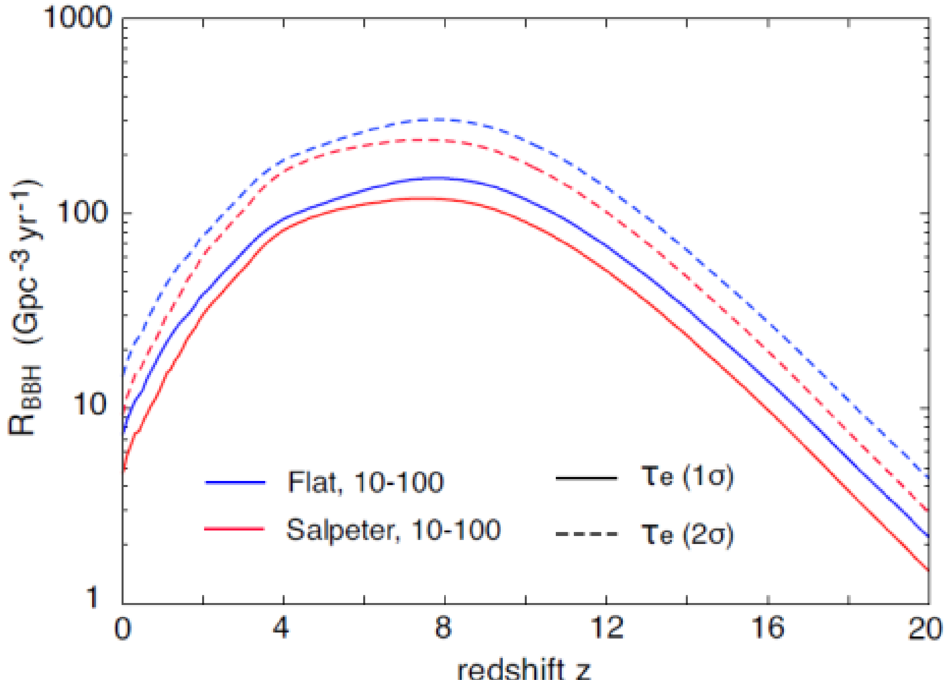
	Errsys
Standard	1 (180 /yr)
Mass range: (10 $M_\odot$<M< 100 $M_\odot$ or *140 $M_\odot$*)	1~3.4
IMF:Flat, M^{-1}, Salpeter	0.42~1
IEF:$f(e)\propto e$,const.,$e^{-0.5}$	0.94~1
BH natal kick: V=0,100,300 km/s	0.2~1
CE:αλ=0.01,0.1,1,10	0.21~1
Mass transfer (mass loss fraction): β=0, 0.5, 1	0.67~1.3
Worst	0.046

- On the other hand, the typical mass is not changed (~30 Msun).

Other Pop III SFRs

- SPH simulation
  (Johnson et al. 2013)
  SFRp~ 10^{-3}-10^{-4} Msun/yr/Mpc^3
- Constraints by Planck
  (e.g.Hartwig et al.2016, Inayoshi et al.2016)
  optical depth of Thomson scattering
  total Pop III density≲$10^{4\text{-}5}$ Msun/Mpc^3
  by Visbal et al.2015





What is the expected Mass of Pop III stars ?

- Without UV feedback

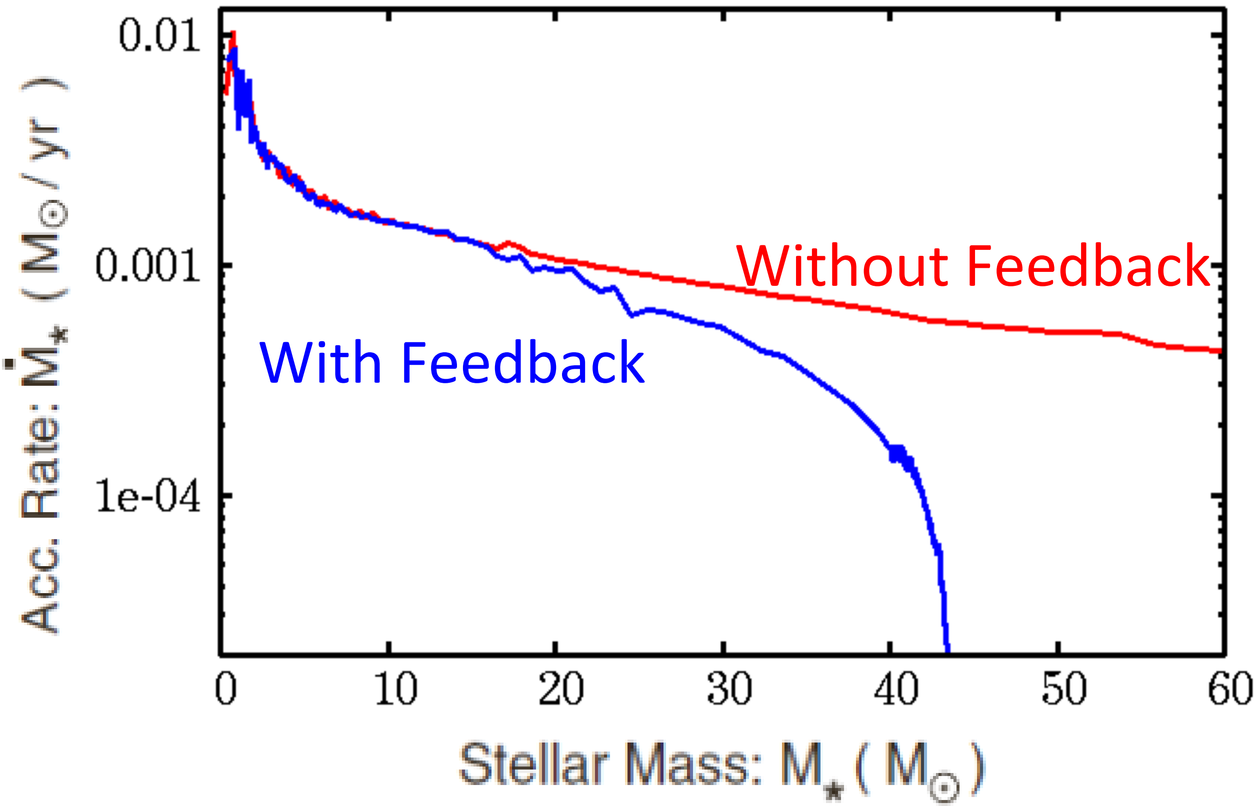
  The typical mass about $10^3\ M_\odot$

  (Omukai & Palla 2003,etc.)

- With UV feedback

  The typical mass 10-100 $M_\odot$

  (Hosokawa et al. 2011, 2012)



Hosokawa et al. 2011

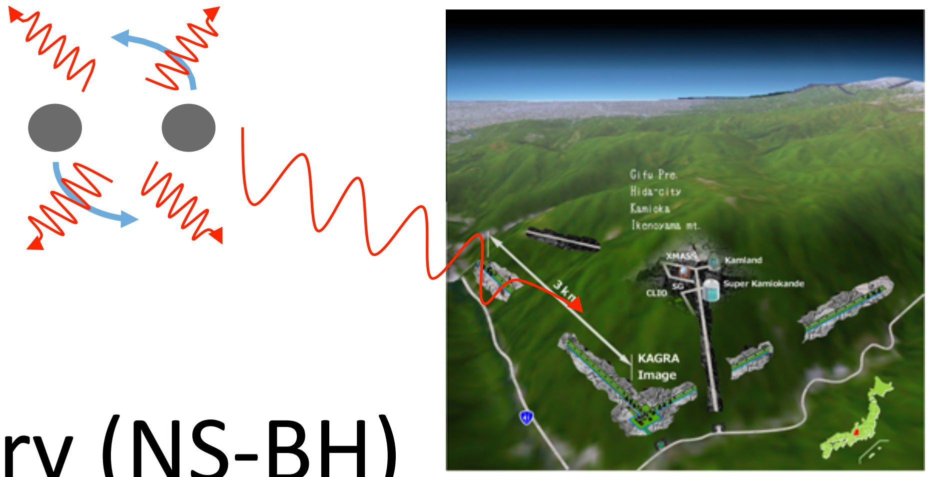
→ Pop III stars → 10-100 $M_\odot$

The differences between Pop III and Pop I

	Pop I stars (Sun like stars)	Pop III stars
Metallicity	2%	0
Radius	Large	Small
Typical Mass	1 Msun	10-100 Msun
Wind mass loss	effective	Not effective

Pop III binaries are easier to be massive compact binary

The main target of gravitational wave source



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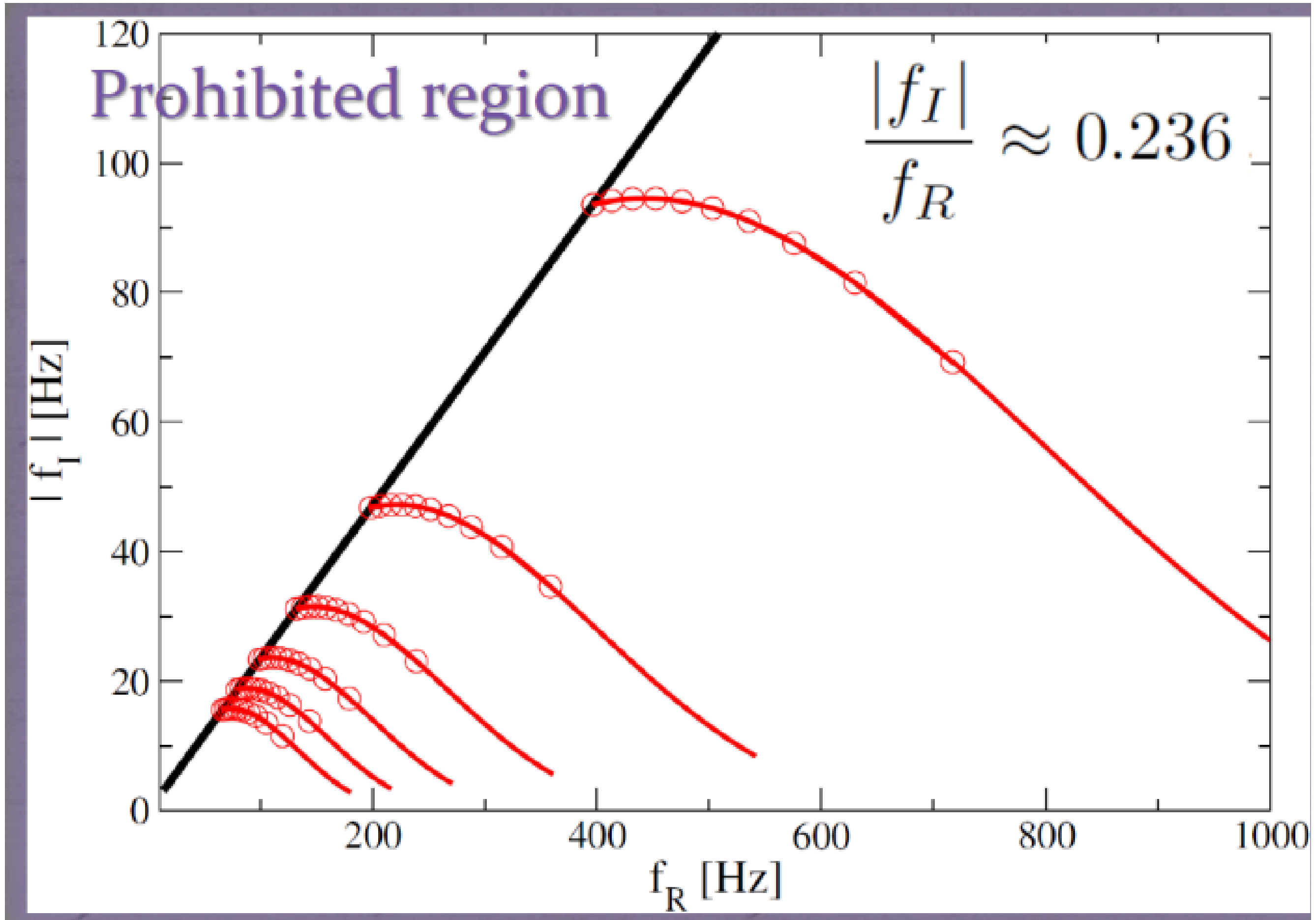
- Compact binary mergers
  - Binary neutron star (NS-NS)
  - Neutron star black hole binary (NS-BH)
  - Binary black hole (BH-BH)

How many times can we detect compact binary mergers？

→Estimated by the binary population synthesis

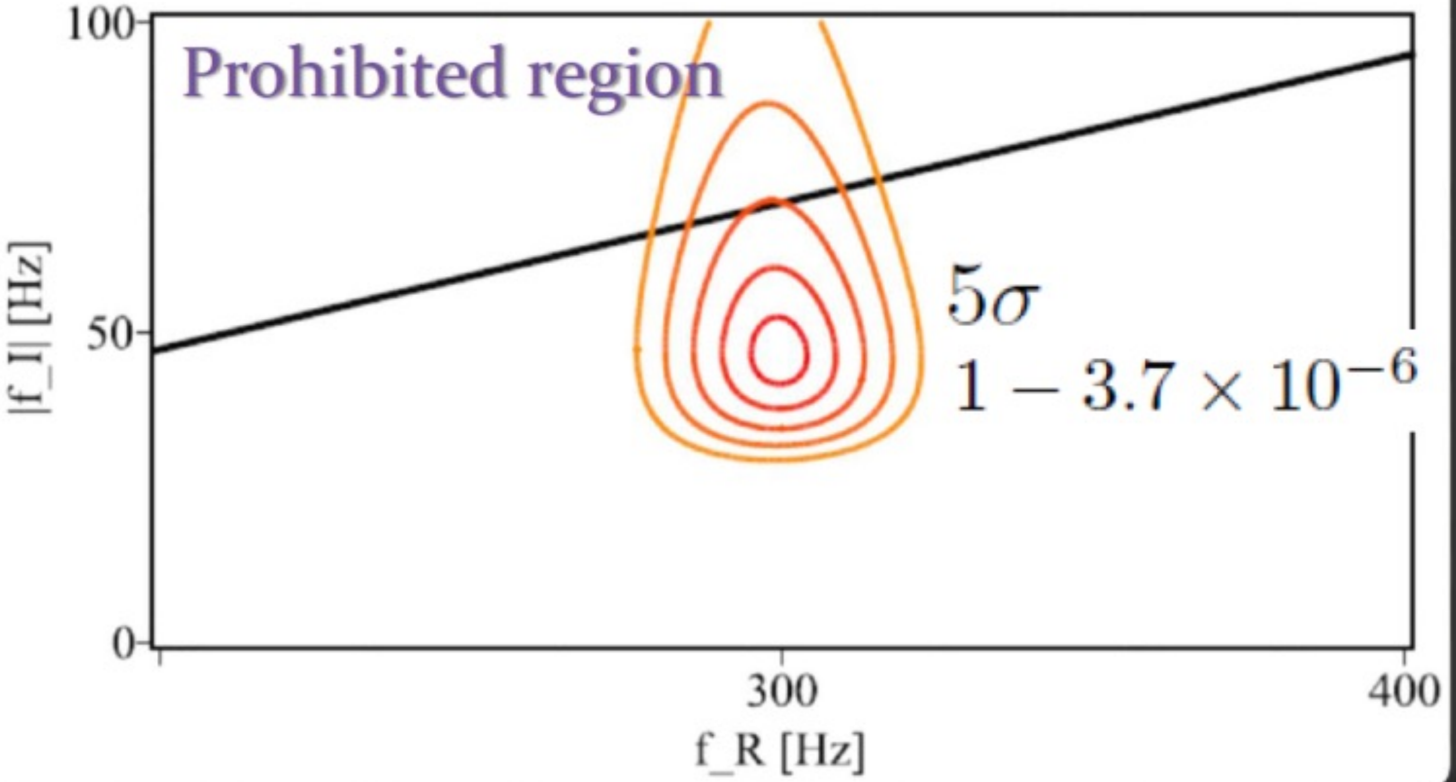
Quasi normal mode



$$f_R = f_c\,, \qquad f_I = -\frac{f_c}{2Q}\,,$$

- fc is frequency of QNM
- Q is the quality factor of QNM which relate to the attenuation of QNM

Only ringdown



- SNR = 20 (500Mpc) for the typical Pop III BBH

$$(M = 60M_\odot\,,\ \eta = 1/4) \rightarrow (M_{\rm rem} = 57.09M_\odot\,,\ \alpha_{\rm rem} = 0.6867)$$
$$f_R = 299.5\text{Hz}\,,\ f_I = -46.34\text{Hz}\,, (f_c = 299.5\text{Hz}\,,\ Q = 3.232)$$

How to calculate the event rate

- NS-NS

  We can get information from binary pulsar observations

  ･The empirical rate from pulsar observations (Kalogera et al. 2004,etc)

  ･Binary population synthesis(Belczynski et al. 2002, 2004, Dominik et al.2012,etc)

- NS-BH,BH-BH

  ･Binary population synthesis

  There were no observation until GW150914.

  Thus, there is no other way except binary population synthesis

Why do Pop III stars have these properties?

- Zero metal stars

-No line cooling and dust cooling at the star formation

-High temperature and high Jeans mass ($M_J \propto T^{3/2}$)

⇒More massive than Pop I stars (Pop I stars are solar like stars)

The typical mass is 10-100$M_{\odot}$

-Missing metal and dust i.e. missing powerful opacity source

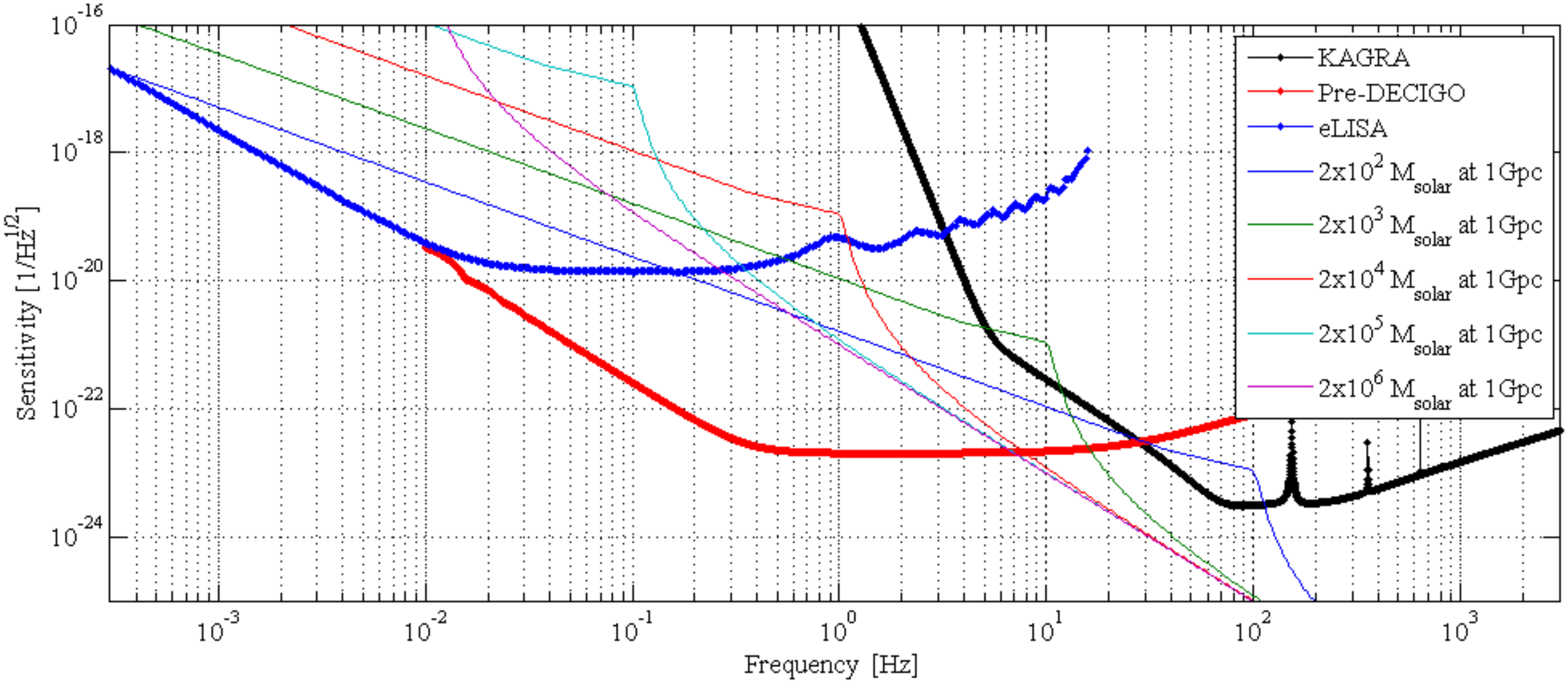
-The stellar photosphere become small

⇒Smaller radius than Pop I stars

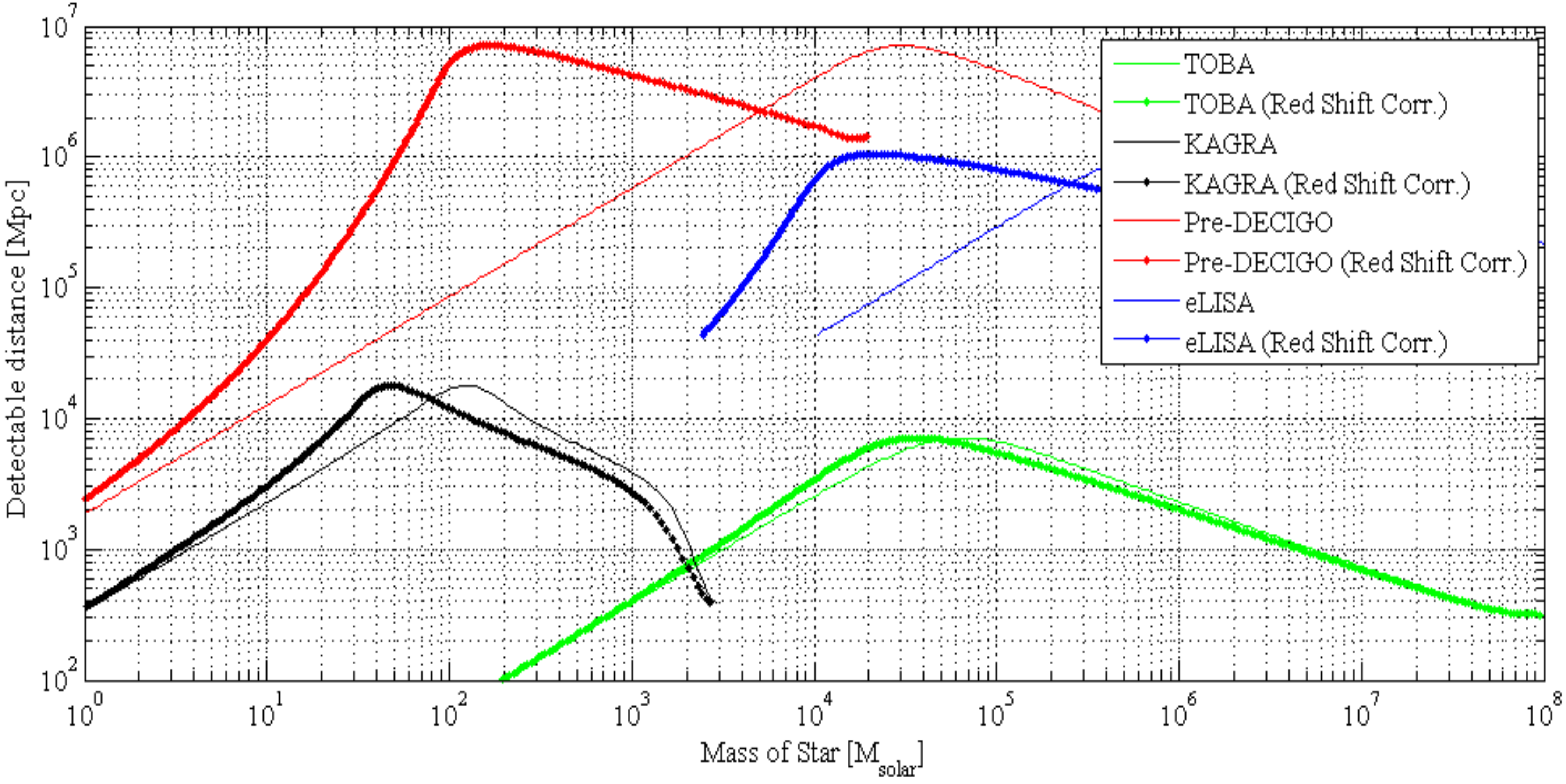
-Stellar wind is driven by radiation pressure on resonance lines of heavier ions or dust grains

-However, Pop III stars do not have heavier ion and dust grain

⇒No wind mass loss

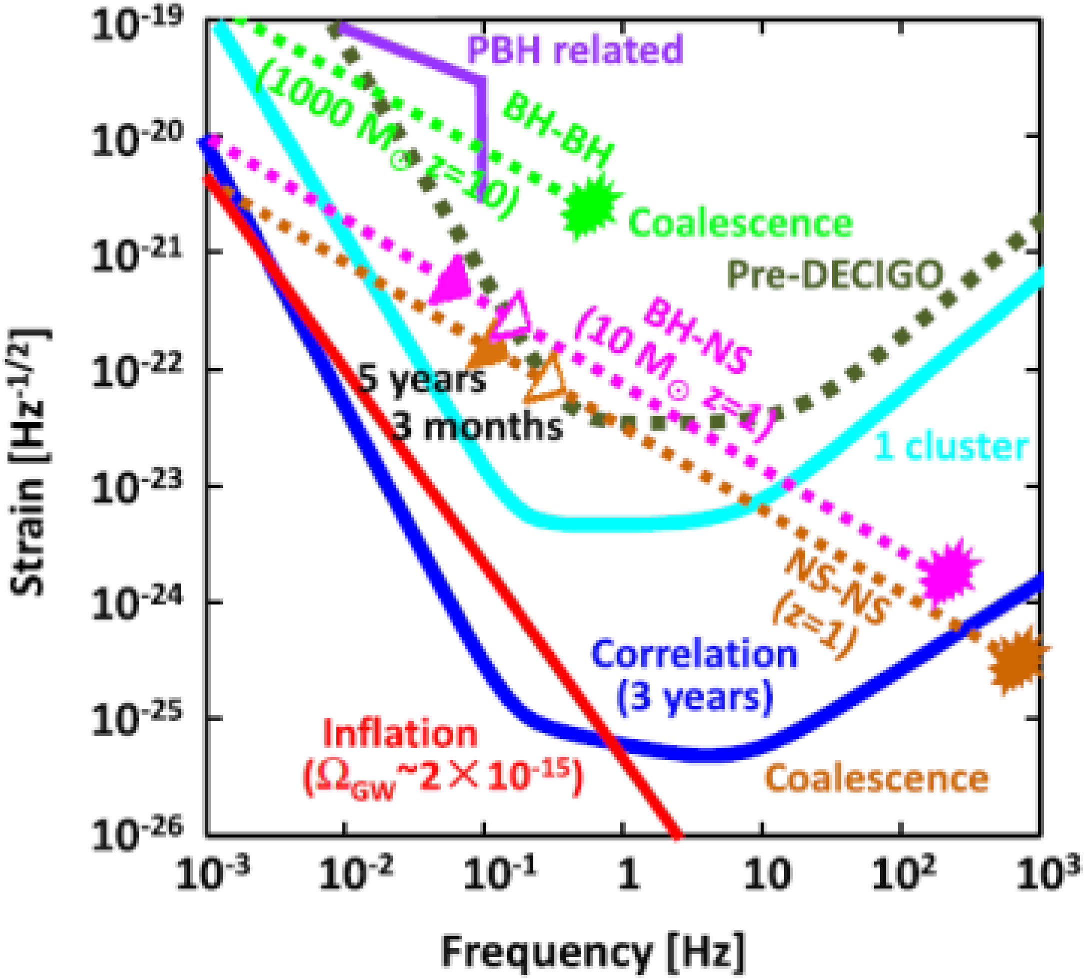


KAGRA
Pre-DECIGO
eLISA
2x10^2 M_solar at 1Gpc
2x10^3 M_solar at 1Gpc
2x10^4 M_solar at 1Gpc
2x10^5 M_solar at 1Gpc
2x10^6 M_solar at 1Gpc
Sensitivity [1/Hz^1/2]
Frequency [Hz]
10^-16
10^-18
10^-20
10^-22
10^-24
10^-3
10^-2
10^-1
10^0
10^1
10^2
10^3



Detectable distance [Mpc]
Mass of Star [M_solar]
TOBA
TOBA (Red Shift Corr.)
KAGRA
KAGRA (Red Shift Corr.)
Pre-DECIGO
Pre-DECIGO (Red Shift Corr.)
eLISA
eLISA (Red Shift Corr.)

DECIGOの感度曲線



Kawamura et al. 2011

- Pop III のSFRのピークはz~9
- Red shift chirp mass=(1+z)Mc
- Pop III BHBH (z~9) ⇒300 Msun (10Hz)

How to calculate the event rate

- NS-NS

  We can get information from binary pulsar observations

  ・The empirical rate from pulsar observations (Kalogera et al. 2004,etc)

  ・Binary population synthesis(Belczynski et al. 2002, 2004, Dominik et al.2012,etc)

- NS-BH,BH-BH

  ・Binary population synthesis

  There is no observation.

  Thus, there is no other way except binary population synthesis

merger rate calculated by population synthesis

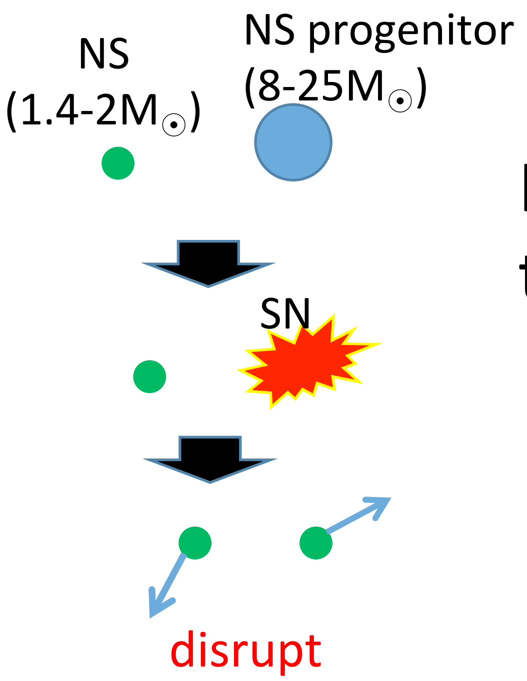
Pop I galactic merger rate [Myr^{-1}] Dominik et al.(2012)

Model	NS-NS	BH-NS	BH-BH
S	23.5 (7.6)	1.6 (0.2)	8.2 (1.9)
V1	0.4 (0.4)	0.002 (0.002)	1.1 (1.1)
V2	11.8 (1.1)	2.4 (0.08)	15.3 (0.4)
V3	48.8 (14.3)	4.6 (0.03)	5.0 (0.03)
V4	20.8 (0.3)	0.9 (0.0)	0.3 (0.0)
		⋮	
V15	39.8 (17.8)	0.01 (0.007)	1.1 (1.0)
Range	0.4–77.4 (0.3–17.8)	0.002–10.6 (0.0–3.9)	0.05–29.7 (0.0–4.2)

These merger rates are calculated by ***Population synthesis (PS)*.**
There are wide differences between models.
I will talk about what is PS and what determine the merger rates.

Why NS-NS disrupt

For example, we consider NS and NS progenitor binary.



In the case of Pop III NS progenitor, wind mass loss and the mass loss due to binary interaction is not effective.
When NS progenitor becomes supernova, NS progenitor suddenly loses mass and becomes NS.
Then, due to instant mass loss the binding energy of binary decreases and binary NS disrupts.

⇒ Binary NS cannot survive!

Binary Interactions

- Supernova effect
- Common envelope

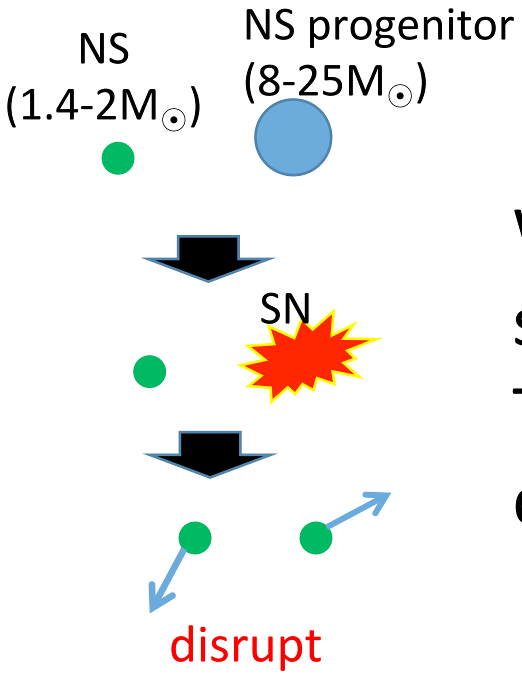
In this talk, I will explain these two binary interactions.

- Stable mass transfer
- Orbital evolution

  (Tidal friction, Gravitational radiation)

Supernova(SN) effect

For example, we consider NS and NS progenitor binary.



When NS progenitor becomes supernova, NS progenitor suddenly loses mass and becomes NS.
Then, due to instant mass loss the binding energy of binary decreases and binary NS disrupts.

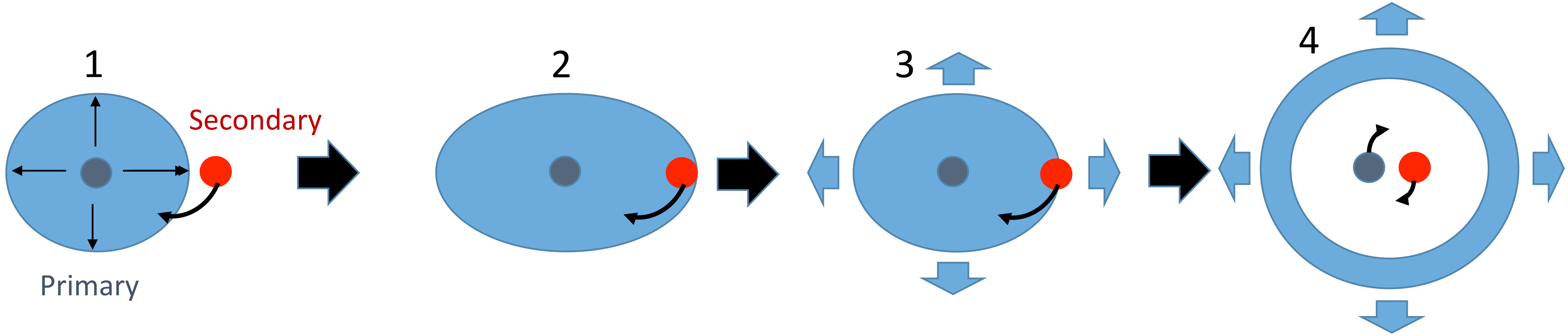
⇨ Binary NS cannot survive!

But in fact binary pulsars have been observed.
Why can binary NS survive?
This reason is common envelope.

Common envelope (CE)

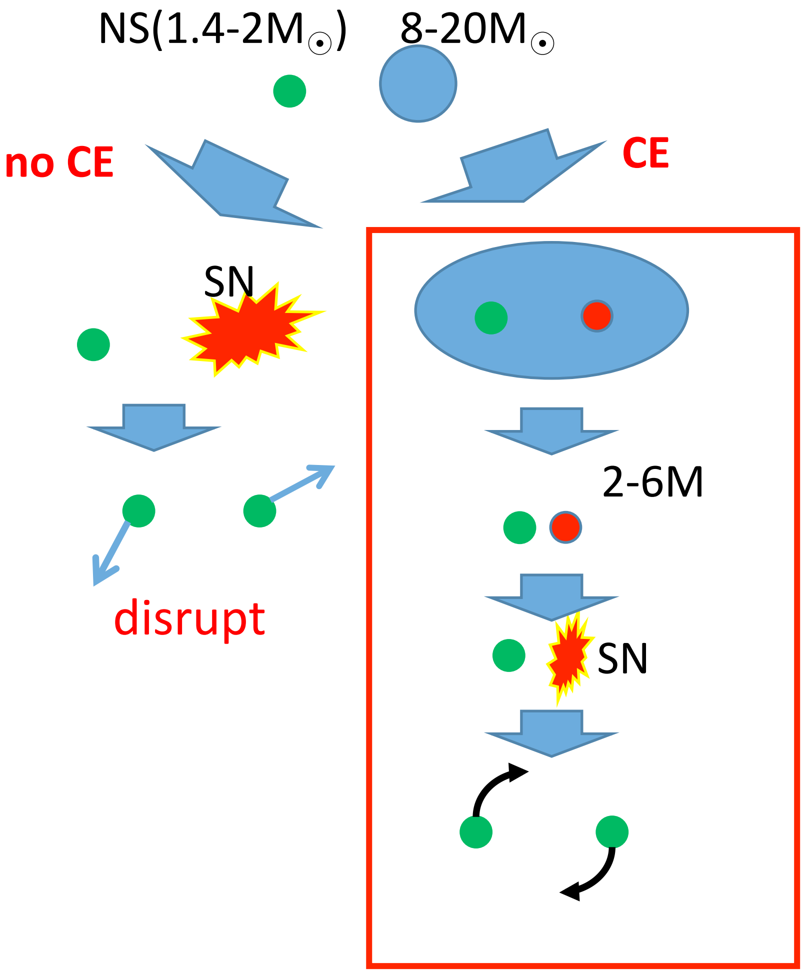
CE is unstable mass transfer phase.

1. Primary star becomes giant and primary radius becomes large.
2. Secondary star plunges in primary envelope.
3. The friction occurs between secondary and primary envelope and transfers angular momentum and energy from orbit to envelope. Due to orbital energy transfer separation decreases and envelope expands and will be expelled.
4. Binary becomes close binary or merges during CE.



Can NS binary survive via CE?

We consider NS and NS progenitor binary again.

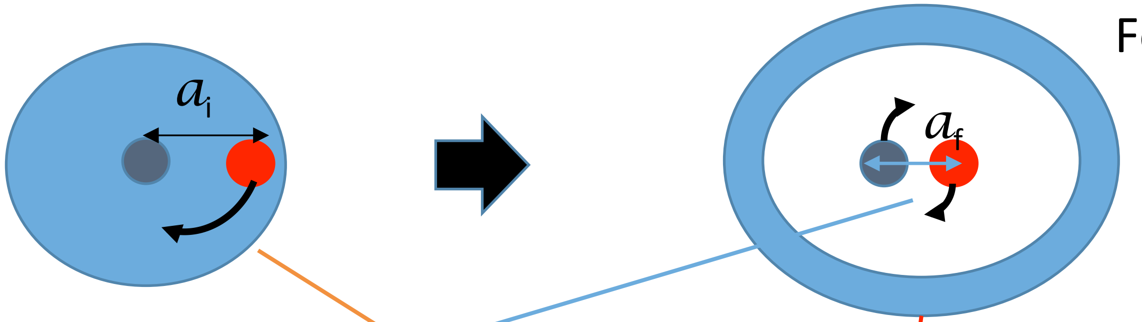


If CE occurs, envelope was already expelled before SN.
Thus, mass ejection at SN becomes smaller than SN mass ejection via no CE.
Due to small mass ejection at SN the loss of binding energy becomes small.
Binary can survive !

Therefore, Common Envelope is important.

The treatment of CE

- We assume the fraction of the orbital energy is used to expel envelope.
- We use simple energy formalism in order to calculate separation after CE a_f



For given Mcore1, Menv1 M2, initial separation a_i

Assuming efficiency of mass ejection

Final separation a_f

$$\alpha\left(\frac{GM_{c1}M_2}{2a_f} - \frac{GM_1M_2}{2a_i}\right) = \frac{GM_1M_{env1}}{\lambda R_1}$$

The loss of orbital energy — the energy required to expel envelope

α: the efficiency of energy transfer from orbit to envelope

λ: the binding energy parameter

These common envelope parameters are uncertain.

- How much the orbital energy can be used to expel envelope?
- How much the internal energy of envelope is used to expel envelope?

The rate dependence on CE parameters

$$\alpha \left(\frac{G M_{c1} M_2}{2 a_f} - \frac{G M_1 M_2}{2 a_i} \right) = \frac{G M_1 M_{env1}}{\lambda R_1}$$

The loss of orbital energy

the energy required to expel envelope

- Separation after CE a_f is dependent on CE parameters.

For simplicity, α=1.

If λ is large i.e, the energy required to expel envelope is small,

the loss of orbital energy during CE becomes small and a_f is large.

- If a_f is large, binary tend not to merge during CE and can survive.
- However, if a_f is too large, binary cannot merge within Hubble time due to GW.

λ ↗, a_f ↗ ⇒

・The number of merger during CE ↘ ⇒ Merger rates ↗

・Merger timescale $t_{GW} \propto a^4$ ↗ ⇒ Merger ↘ rates

The dependence on CE parameters

For example, we consider how Pop I NS-NS merger rate depend on CE parameters.

Pop I NSNS merger rate [Myr^{-1} $galaxy^{-1}$]   Dominik et al.2012

parameter	NS-NS merger rate [events/yr/galaxy]
$\alpha\lambda = 0.01$	0.4
$\alpha\lambda = 0.1$	11.8
$\alpha\lambda = 1$	48.8
$\alpha\lambda = 10$	20.8

αλ
a_f ↗

・The number of coalescence during CE ↘ ⇒ Merger rates ↗

・Merger timescale $t_{GW} \propto a^4$ ↗ ⇒ Merger ↘

Binary population synthesis

- Population synthesis is a method of numerical simulation to research the population of stars with a complex evolutions.
- Population synthesis can predict properties and merger rates of unobserved sources such as NS-BH, BH-BH
- The common envelope of the key process of population synthesis
- However, Common envelope parameters are uncertain.

  This uncertainty change event rate by a factor of several hundreds.

  We should reveal this uncertainty via comparison between result of population synthesis and observations such as GW and other observations and improve binary evolution theory

Example: CE dependence

We calculate $\alpha\lambda$=0.01, 0.1, 1, 10 cases $N_{total}=10^6$

	Standard(1)	0.01	0.1	10
NSNS	0	0	0	1116
NSBH	185335	148290	162814	198408
BHBH	517067	340893	434590	542399
merged NSNS	0	0	0	890
merged NSBH	50	0	45	767
merged BHBH	115056	32283	111696	91787

The number of merged Pop III BH-BH change by a factor of ***several***.
On the other hand, Pop I merger rates changed by a factor of ***several hundreds***.

Why do Pop III stars have these properties?

- Zero metal stars

-No line cooling and dust cooling at the star formation

-High temperature and high Jeans mass ($M_J \propto T^{3/2}$)

⇒More massive than Pop I stars (Pop I stars are solar like stars)

The typical mass is 10-100$M_\odot$

-Missing metal and dust i.e. missing powerful opacity source

-The stellar photosphere become small

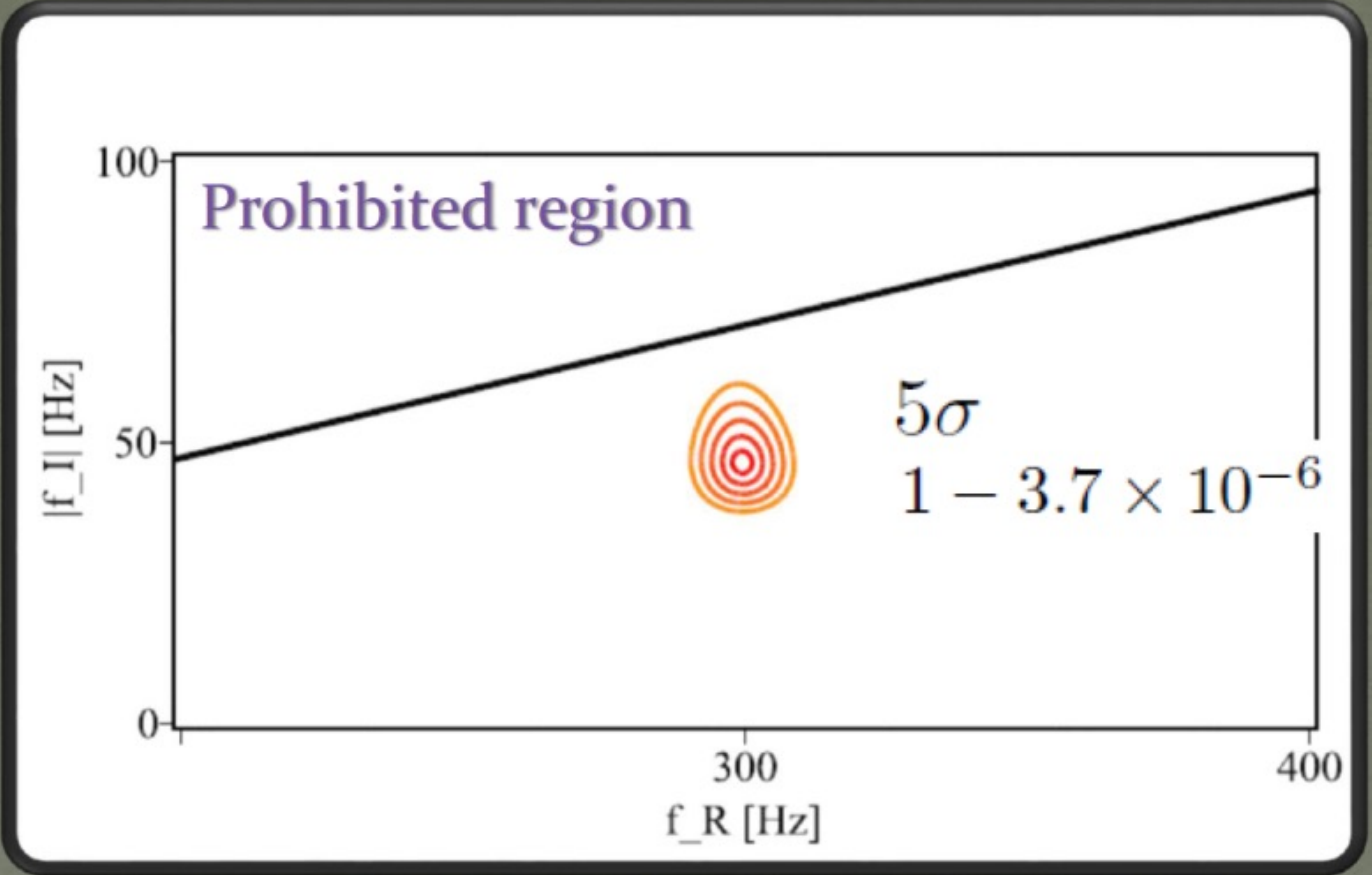
⇒Smaller radius than Pop I stars

-Stellar wind is driven by radiation pressure on resonance lines of heavier ions or dust grains

-However, Pop III stars do not have heavier ion and dust grain

⇒No wind mass loss

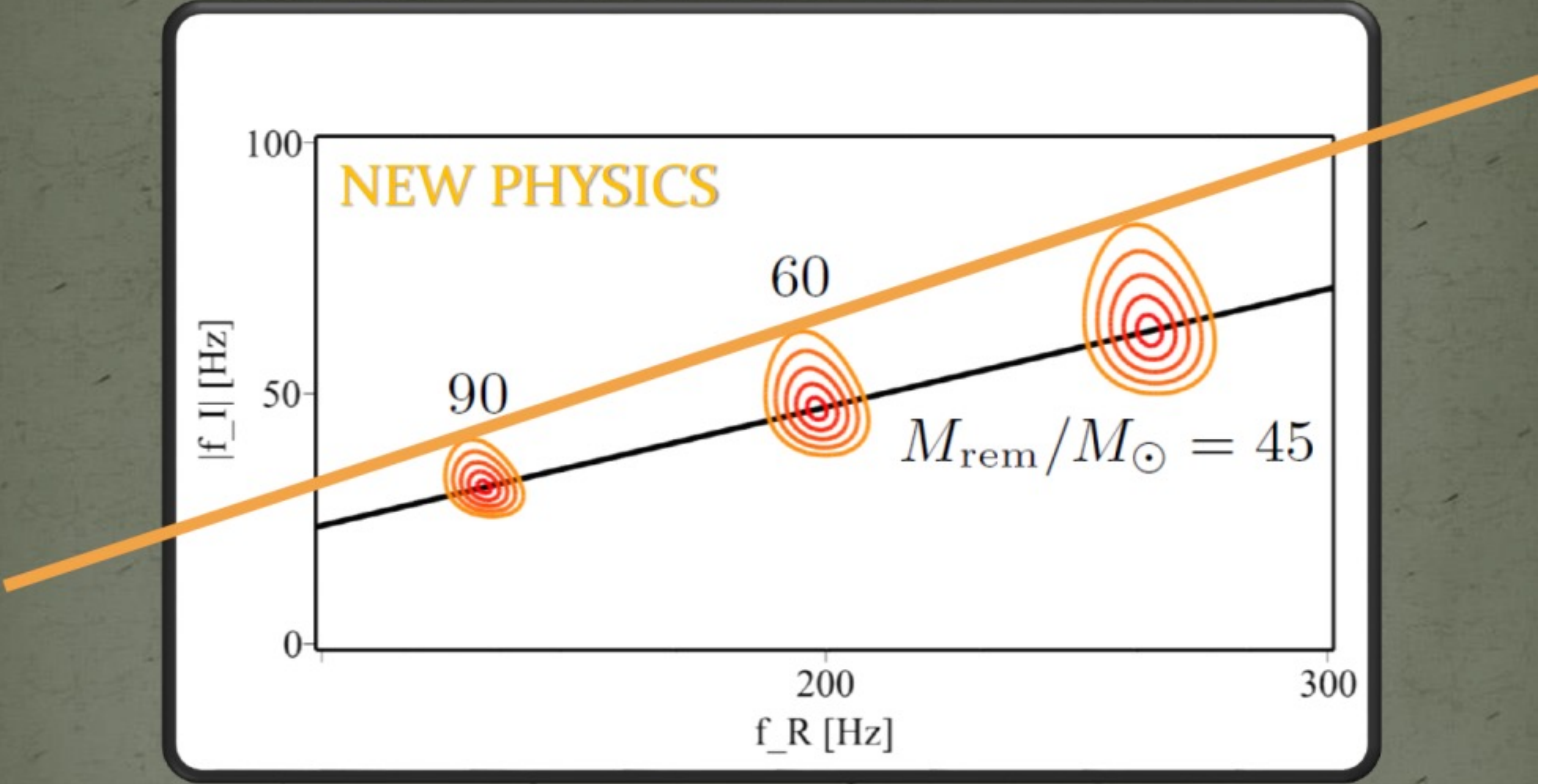
Only ringdown



- SNR = 50 (200Mpc) for the typical Pop III BBH

$$(M = 60M_\odot\,,\ \eta = 1/4) \to (M_{\rm rem} = 57.09M_\odot\,,\ \alpha_{\rm rem} = 0.6867)$$
$$f_R = 299.5\mathrm{Hz}\,,\ f_I = -46.34\mathrm{Hz}\,, (f_c = 299.5\mathrm{Hz}\,,\ Q = 3.232)$$

Only ringdown

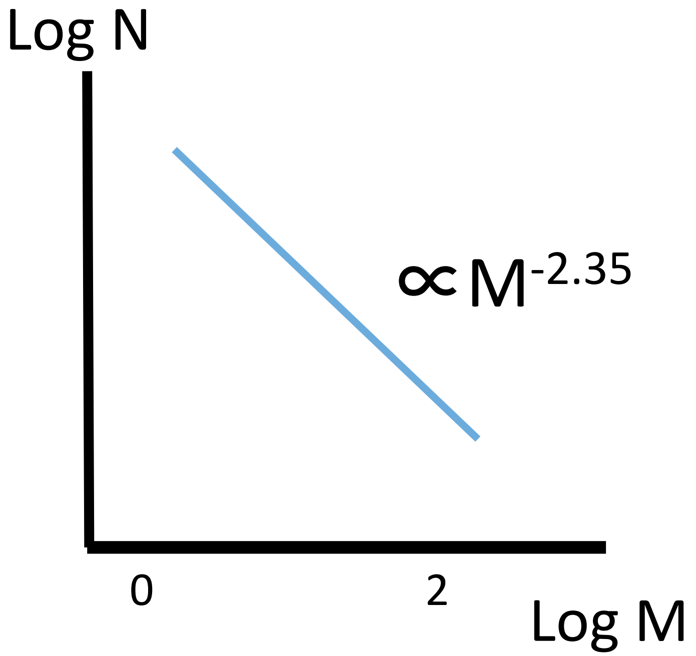


- SNR = 50 (200Mpc), Schwarzschild case

IMF

- Pop I

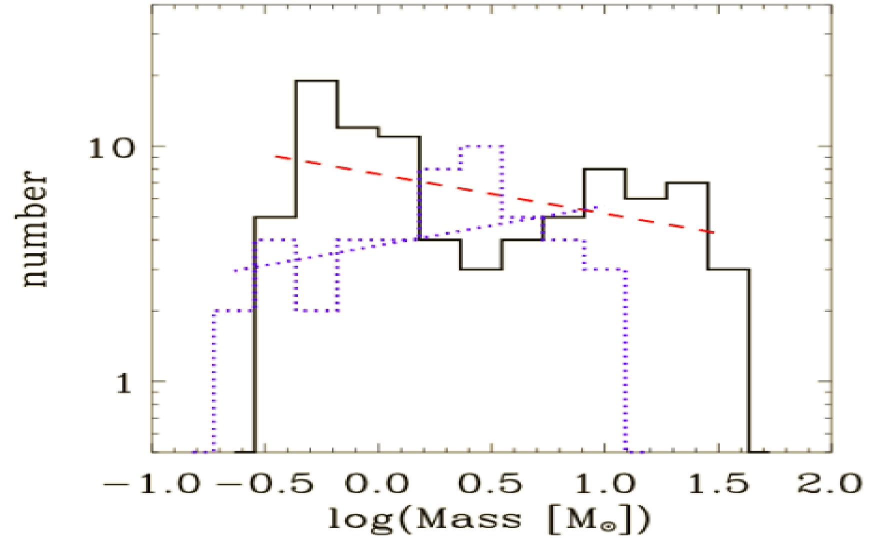
Salpeter



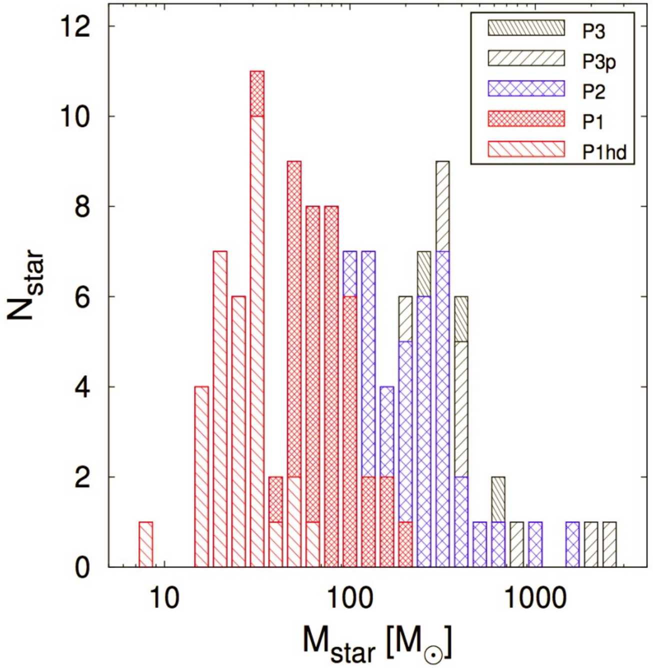
- Pop III

Flat？

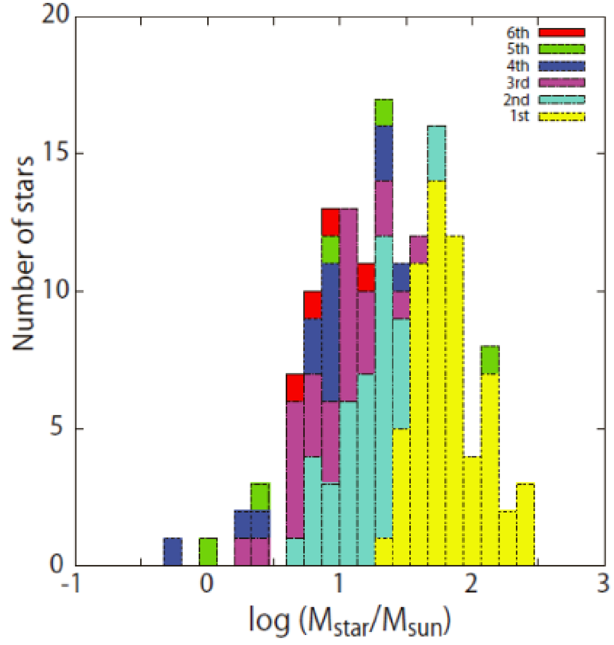
Log Flat？



Stacy & Bromm 2013



Hirano et al.2014



Susa et al. 2014

IMF dependence

	Standard(Flat)	M^{-1}	Salpeter
NSNS	0	2	5
NSBH	185335	168100	93085
BHBH	517067	350169	132534
merged NSNS	0	2	5
merged NSBH	50	68	64
merged BHBH	115056	74745	25536

Uncertainties of Pop III binary population synthesis

- Initial condition
  - IMF
  - mass ratio
  - separation
  - eccentricity
- Binary interactions
  - Common envelope
  - Mass transfer
  - Supernova kick

eccentricity distributions

- General eccentricity distribution (Heggie 1975)
  $P(e) \propto e$ (Standard)
- CygnusOB2 association（Kobulnicky et al. 2014）
  $P(e)=const.$
- Observations of O stars(M>15Msun) (Sana et al.2012)
  $P(e) \propto e^{-0.5}$

eccentricity dependence

	Standard(e)	const	$e^{-0.5}$
NSNS	0	0	0
NSBH	185335	183460	181650
BHBH	517067	522809	523285
merged NSNS	0	0	0
merged NSBH	50	43	38
merged BHBH	115056	111106	107594

Uncertainties of Pop III binary population synthesis

- Initial condition
  - IMF
  - mass ratio
  - separation
  - eccentricity
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Mass transfer

$$\dot{M}_2 = -(1-\beta)\dot{M}_1$$

- β=0 : conservative
- 1>β>0 : non conservative

In Standard model, we use the fitting function

$$\begin{cases} \dot{M}_2 = \min\left(10\frac{\tau_{\dot{M}}}{\tau_{\mathrm{KH},2}}, 1\right)\dot{M}_1 & \text{Secondary is MS or He-burning} \\ M_2 = -M_1 & \text{Secondary is giant} \end{cases}$$

(Hurley et al. 2002)

This is fitted for Pop I stars.

Thus, we check β=0,0.5,1 cases.

Mass transfer dependence

	Standard(func.)	0	0.5	1
NSNS	0	0	5	1359
NSBH	185335	185335	193921	218311
BHBH	517067	517067	549893	531452
merged NSNS	0	0	5	1358
merged NSBH	50	50	199	119
merged BHBH	115056	115056	117094	50119

Supernova kick

- Pulsar kick ~200-500km/s

  Pulsar observation suggest NSs have the natal kick at the SN.

- BHXRBs have large distance from galactic plane.

  Black hole natal kick?（Repetto,Davis&Sigurdsson2012）

⇒We check the kick dependence.

σ=0km/s (Standard)、σ=100km/s、σ=300km/s

$$P(v_k) = \sqrt{\frac{2}{\pi}} \frac{v_k^2}{\sigma_k^2} \exp\left[-\frac{v_k^2}{\sigma_k^2}\right]$$

SN kick dependence

	Standard(0)	100 km/s	300 km/s
NSNS	0	283	8
NSBH	185335	32701	11922
BHBH	517067	191755	70728
merged NSNS	0	17	1
merged NSBH	50	2527	3893
merged BHBH	115056	117415	51928