

# Can nuclear physics solve “the missing gold problem” in the evolution of Galaxy?

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科研費  
KAKENHI

# Talk plan

- Astrophysics part:

a brief summary on the missing problem in GCE, or a bit ironical overview from nucleosynthesis person

- Nucleosynthesis part:

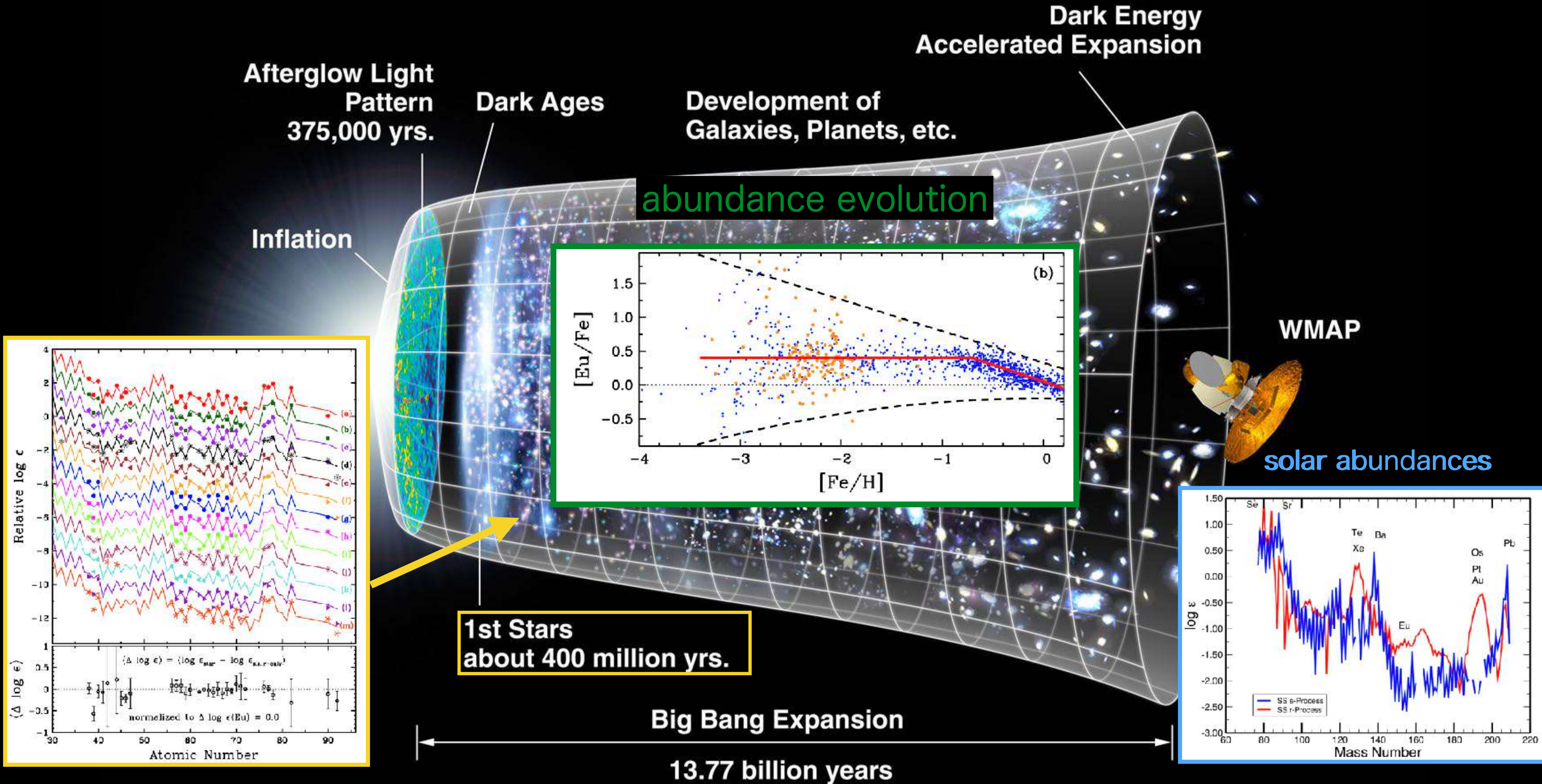
- production of 3rd r-process peaks)
- basics: production of r-process peaks
- previous and on-going studies
- MC-based sensitivity study

- Summary and perspective (personal)

# **1. Astrophysics part:**

**“The missing gold problem” in GCE**

# Evolution of r-process elements

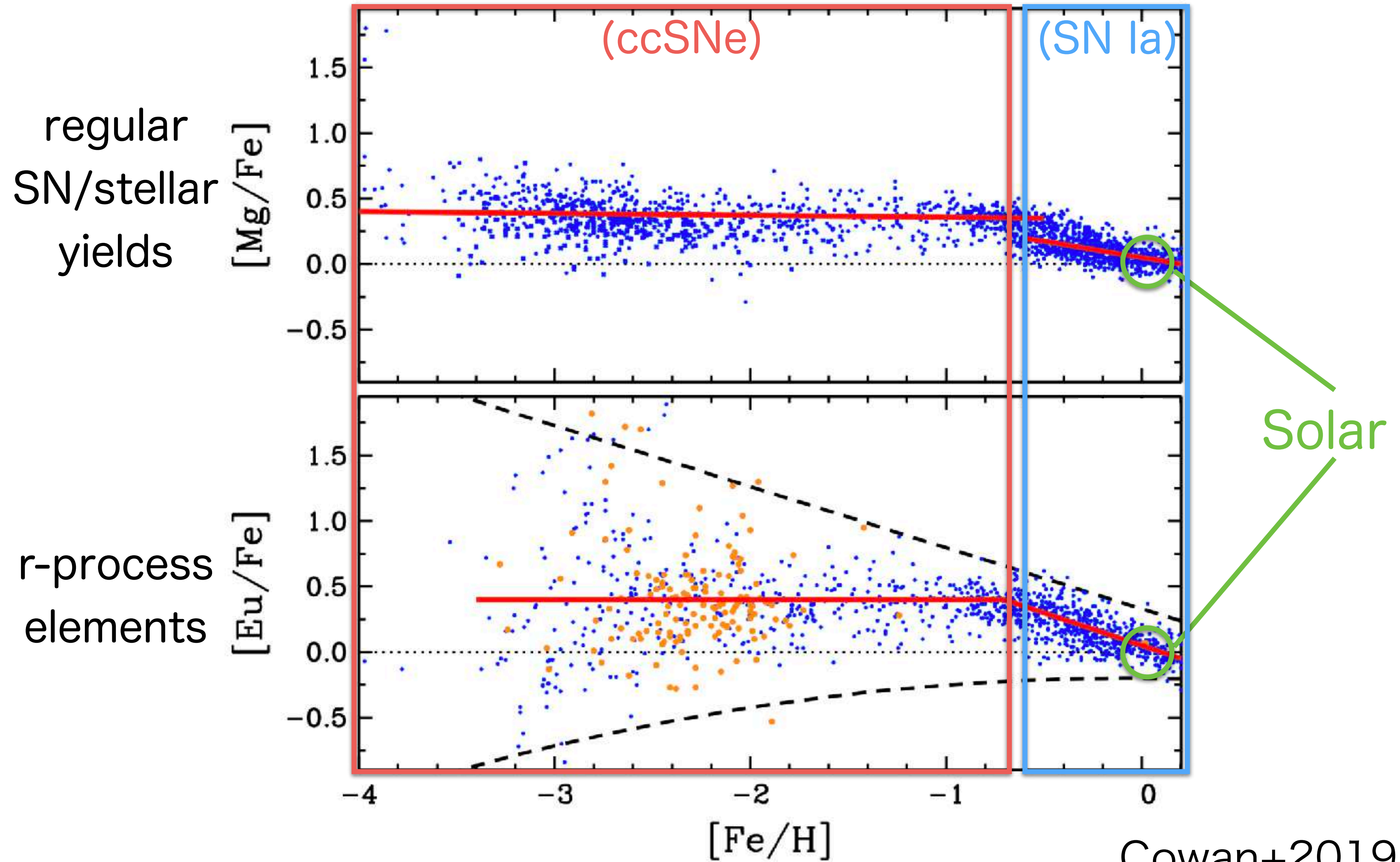


# Galactic Chemical Evolution (GCE)

Galactic halo stars

massive stars

low mass stars



Cowan+2019

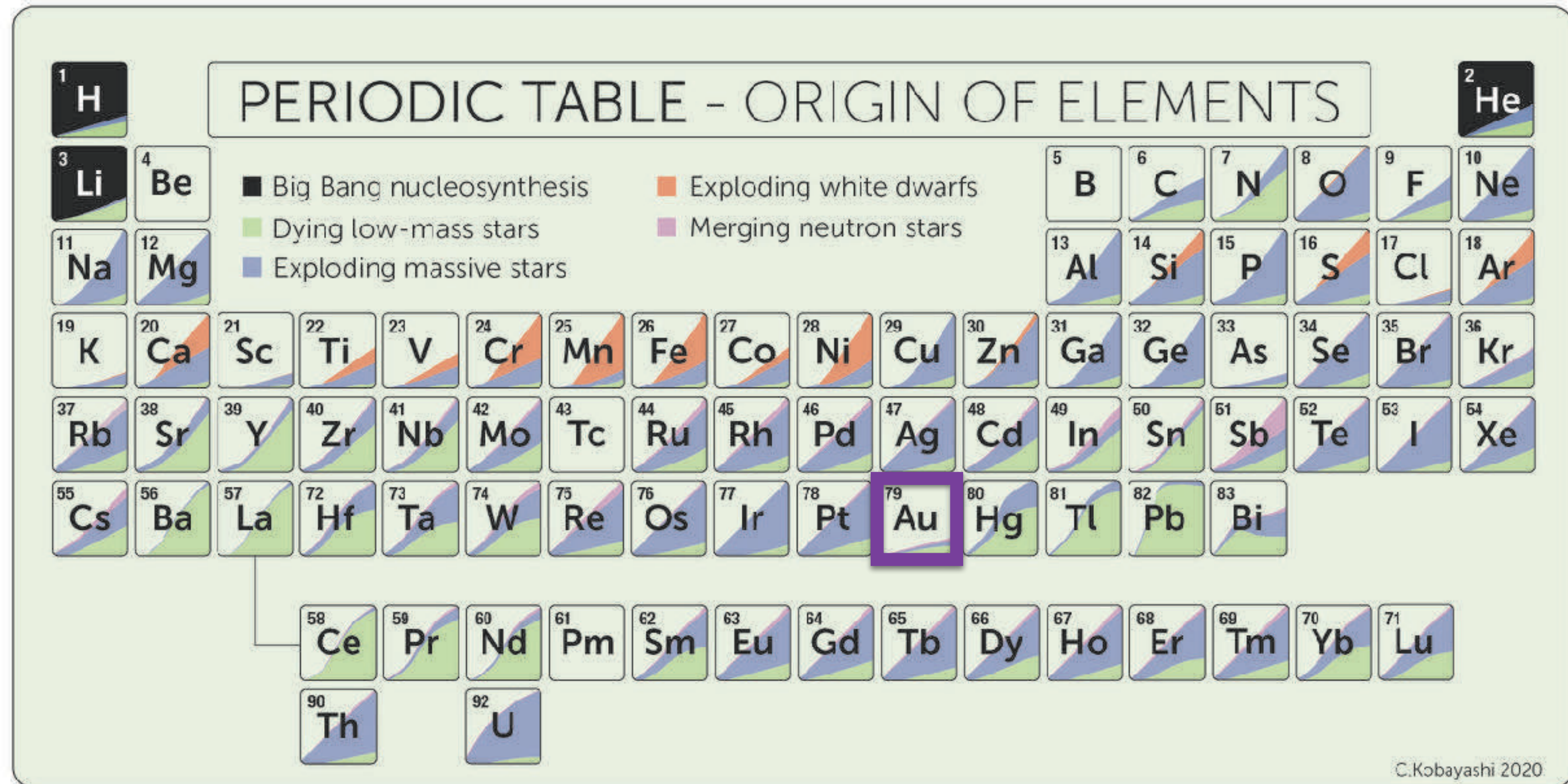
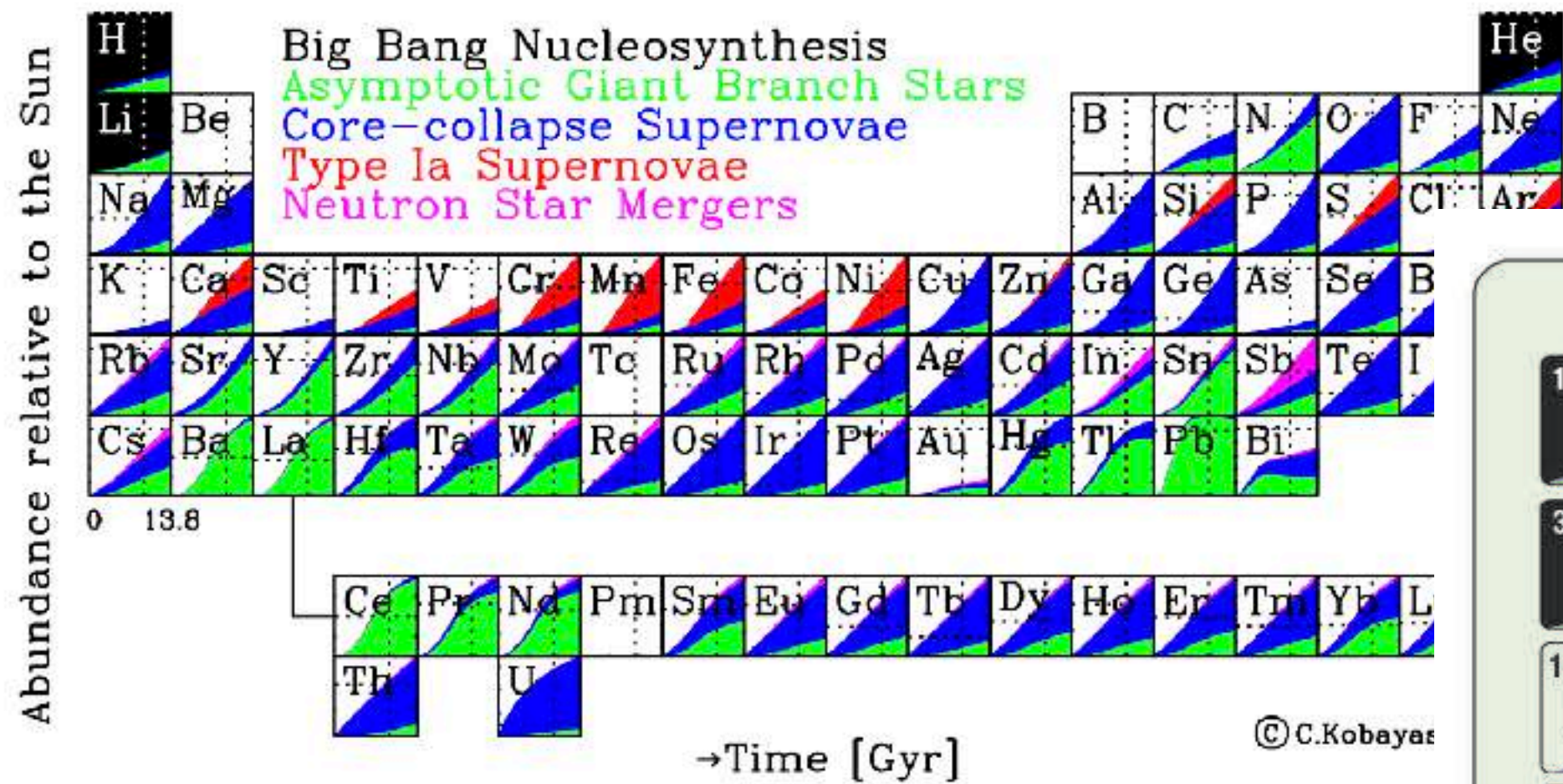
Iron (time evolution)

# The beginning of the story

C. Kobayashi, A. Karakas, M. Lugaro, ApJ 2020

“The origin of Elements from Carbon to Uranium”

- challenge to explain all the isotopes in GCE (Galactic Chemical Evolution)
- they assume supernovae are the main source of r-process (than NS mergers)



# Astrophysical r-process sites

## core-collapse SNe

- NO direct r-process observation
- **Theoretically difficult**
- not very neutron-rich

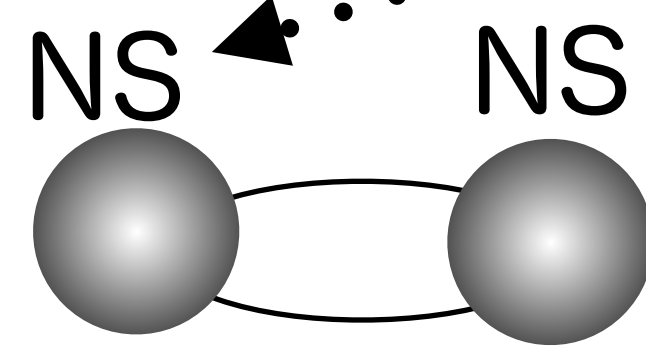
Massive stars

( $10 > M_{\text{sun}}$ )

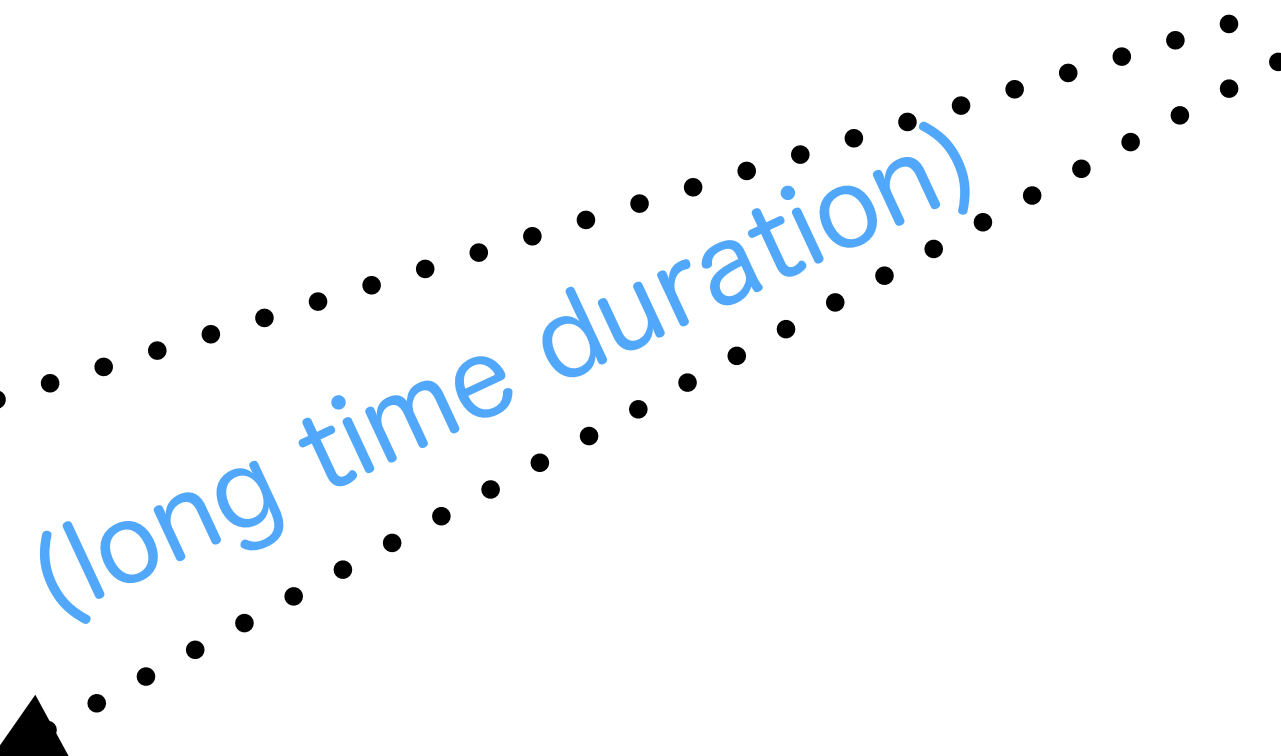
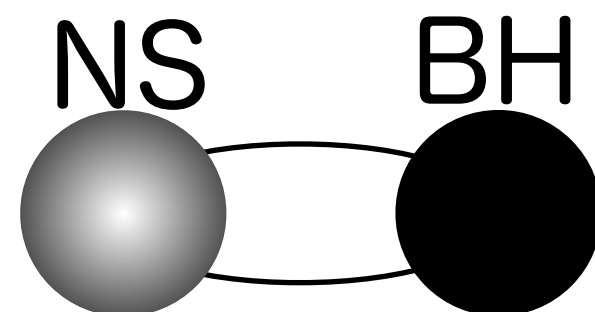


## NS (binary) mergers

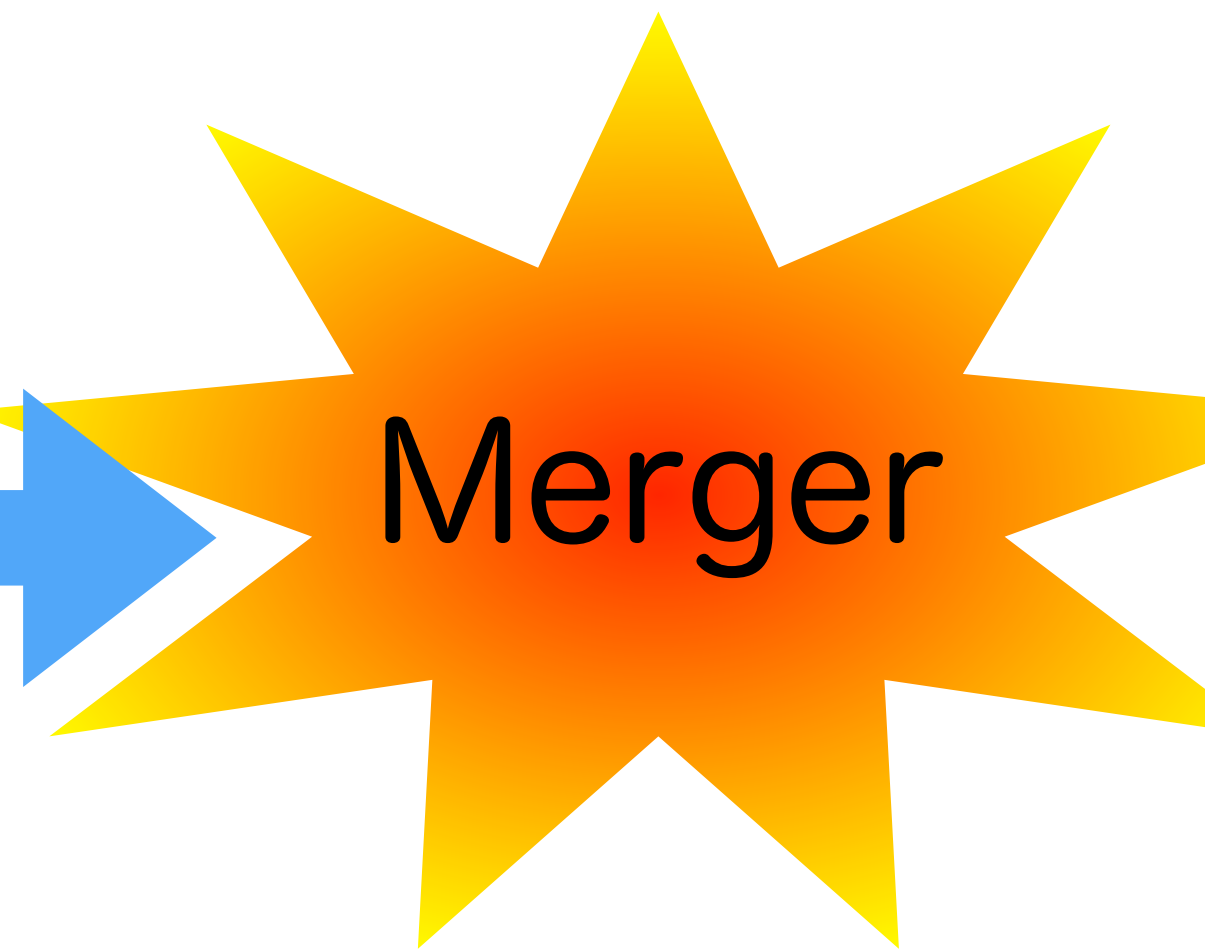
The “observational” evidence with gravitational waves (GW170817)



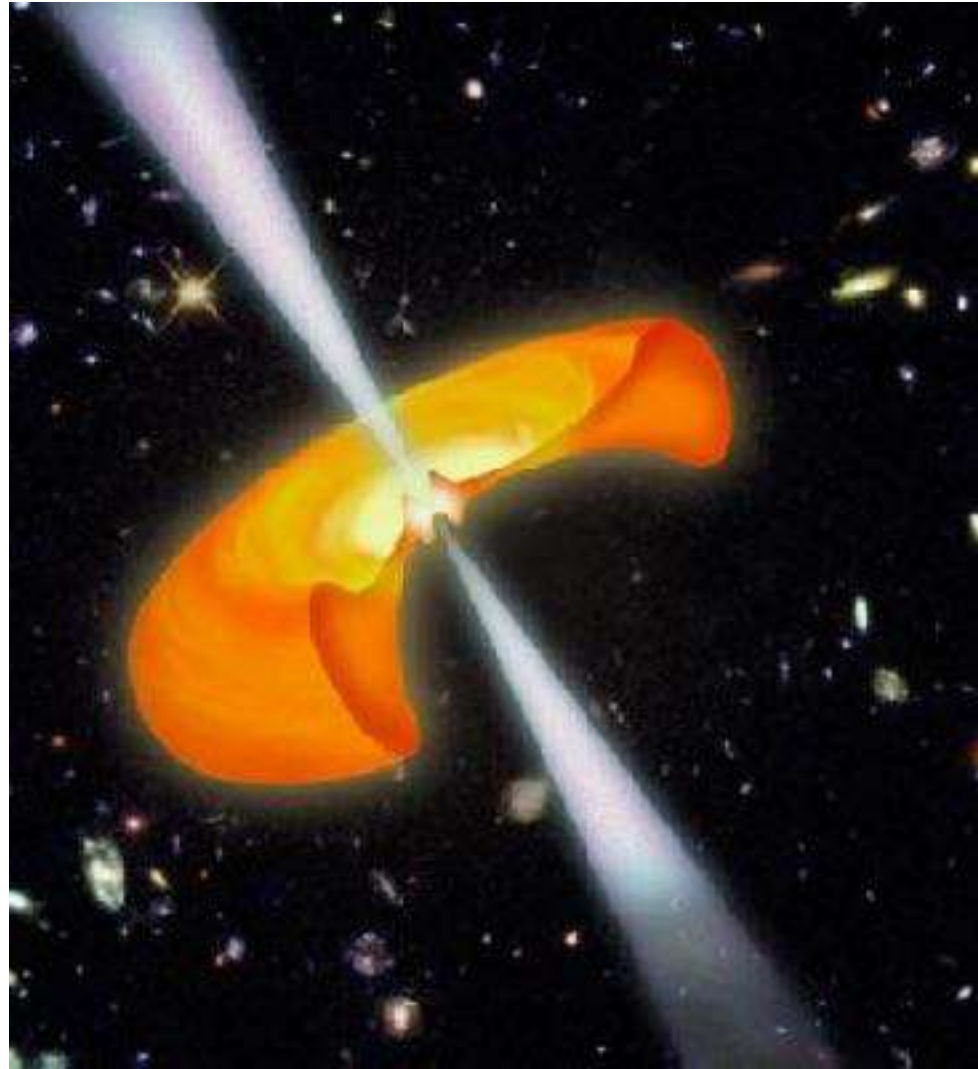
NS binaries



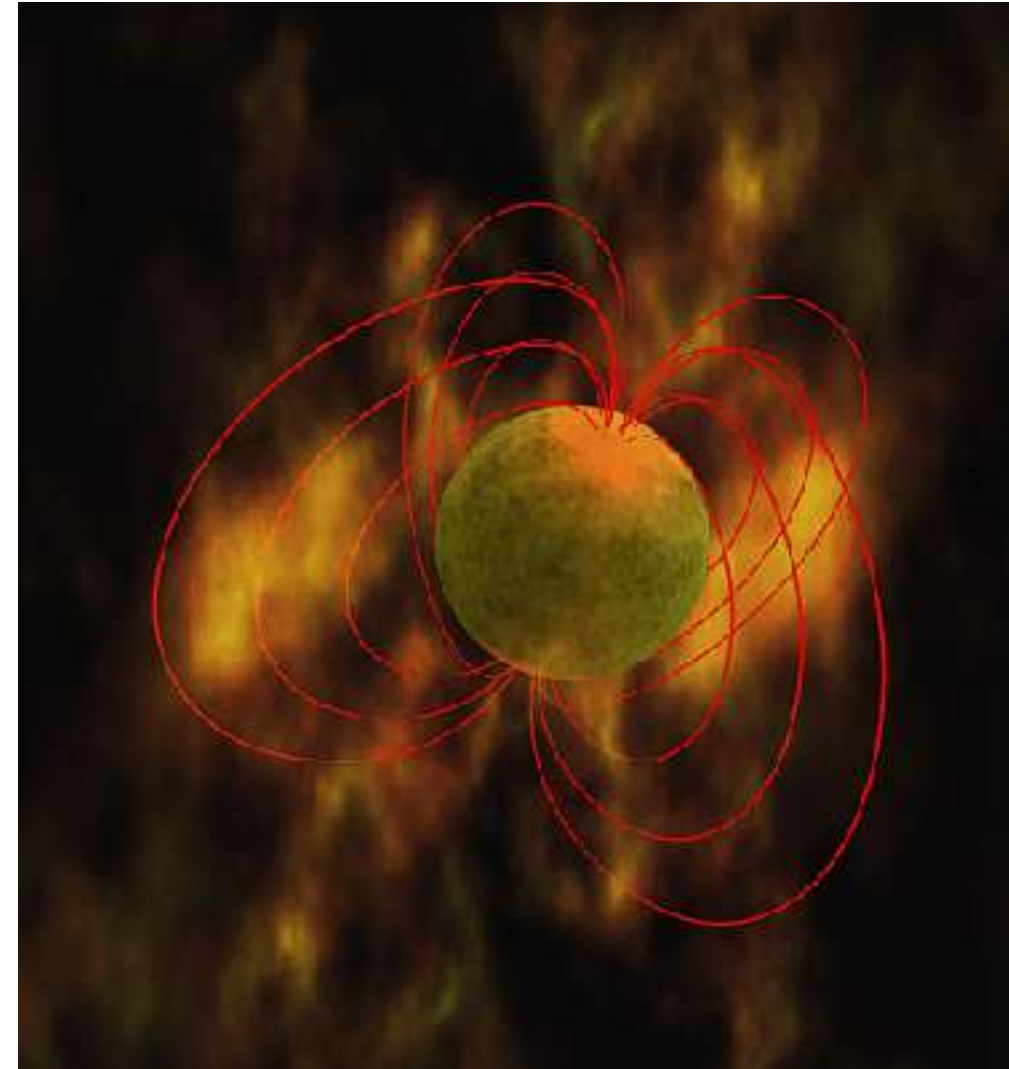
(long time duration)



# Magneto-rotational SN scenario



hypernova/jet-like SN

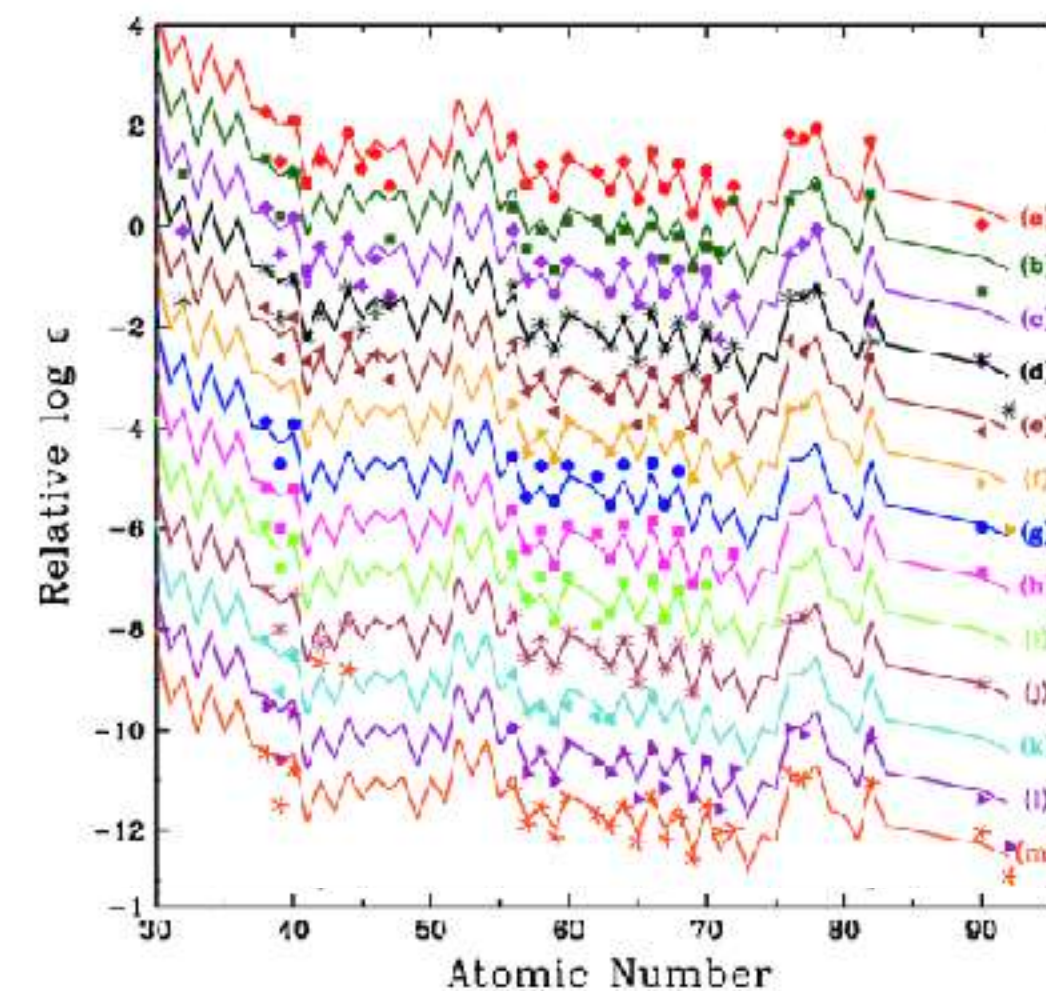


magnetars

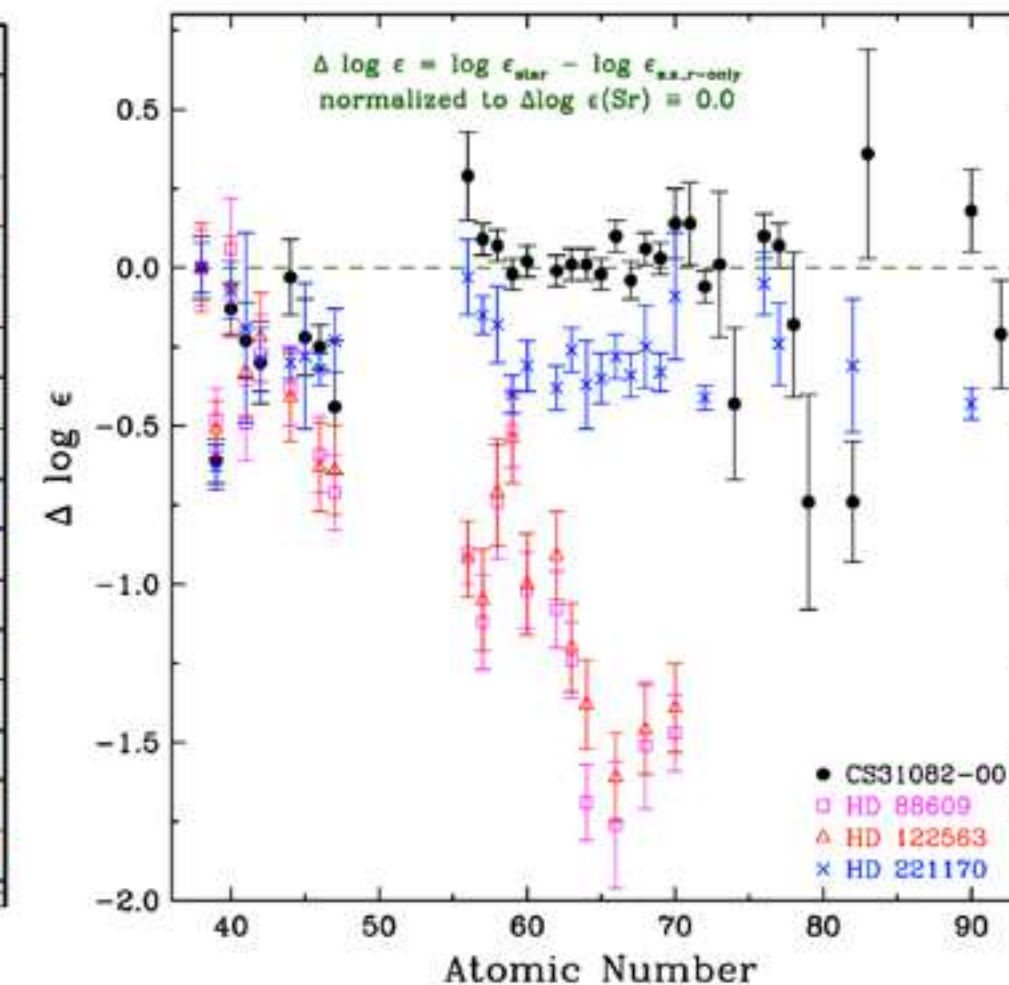
- Magnetars
  - strong magnetic field  $\sim 10^{15}$  G ( $\sim 1\%$  of all neutron stars)
- Magneto-driven Supernovae?
  - GRB central engine
  - Hypernovae?
  - (magnetar driven) Super luminous SNe?

- **variety of r-process patterns** in metal-poor stars
- can be rare  $\sim 1\%$  of ccSN rate
- Galactic chemical evolution
  - needed as **external sources** with NS mergers?
  - MR-SNe, “hypernovae”, collapsars etc.??  
(see, e.g., Wehmeyer+2015, Tsujimoto&NN 2015, Cescutti+2017, Siegel+2019, Kobayashi+2020 etc.)

r-process in MP stars



“weak” r-process?

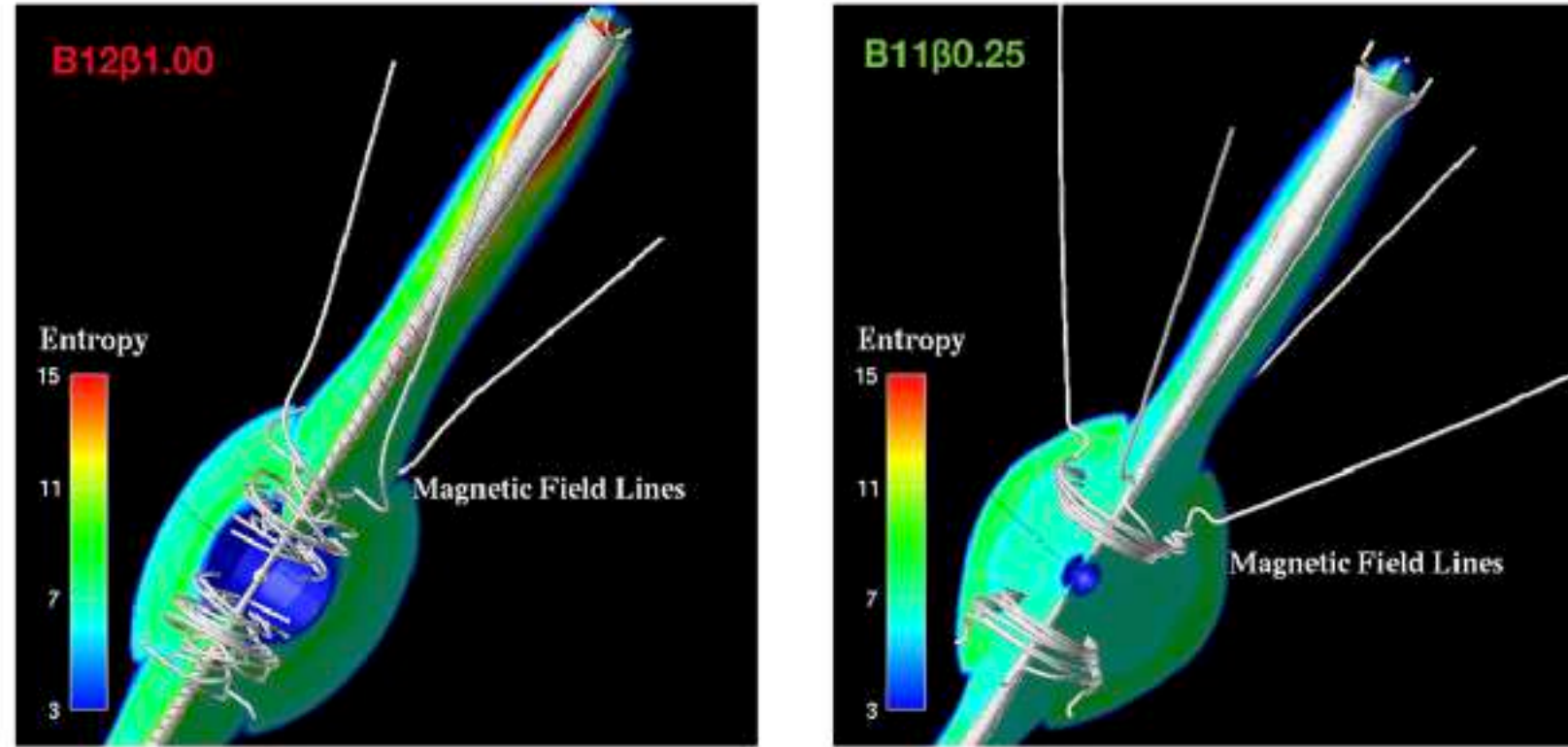


Cowan+2021

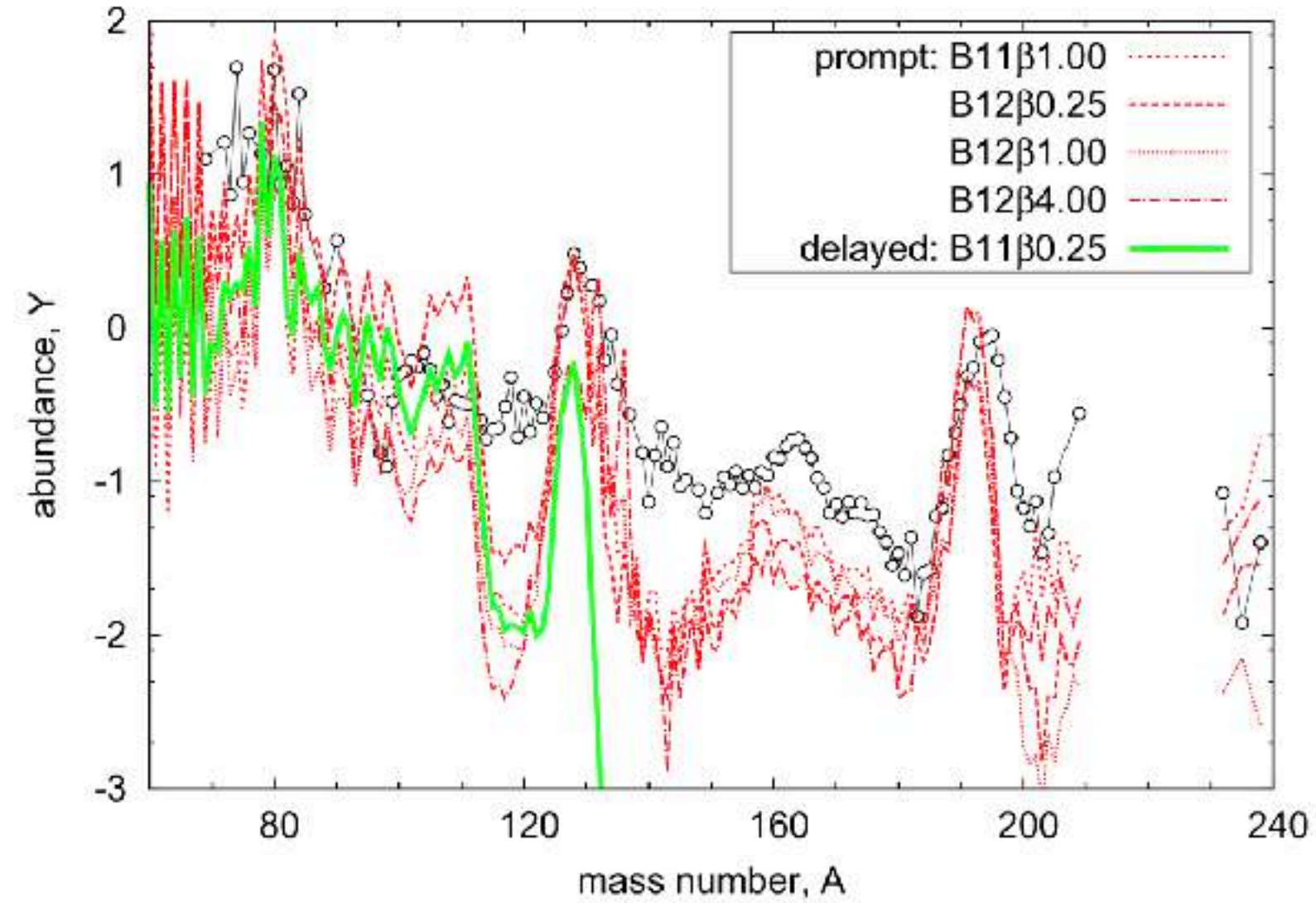


# Prequel story?

Nishimura+(2015) ApJ

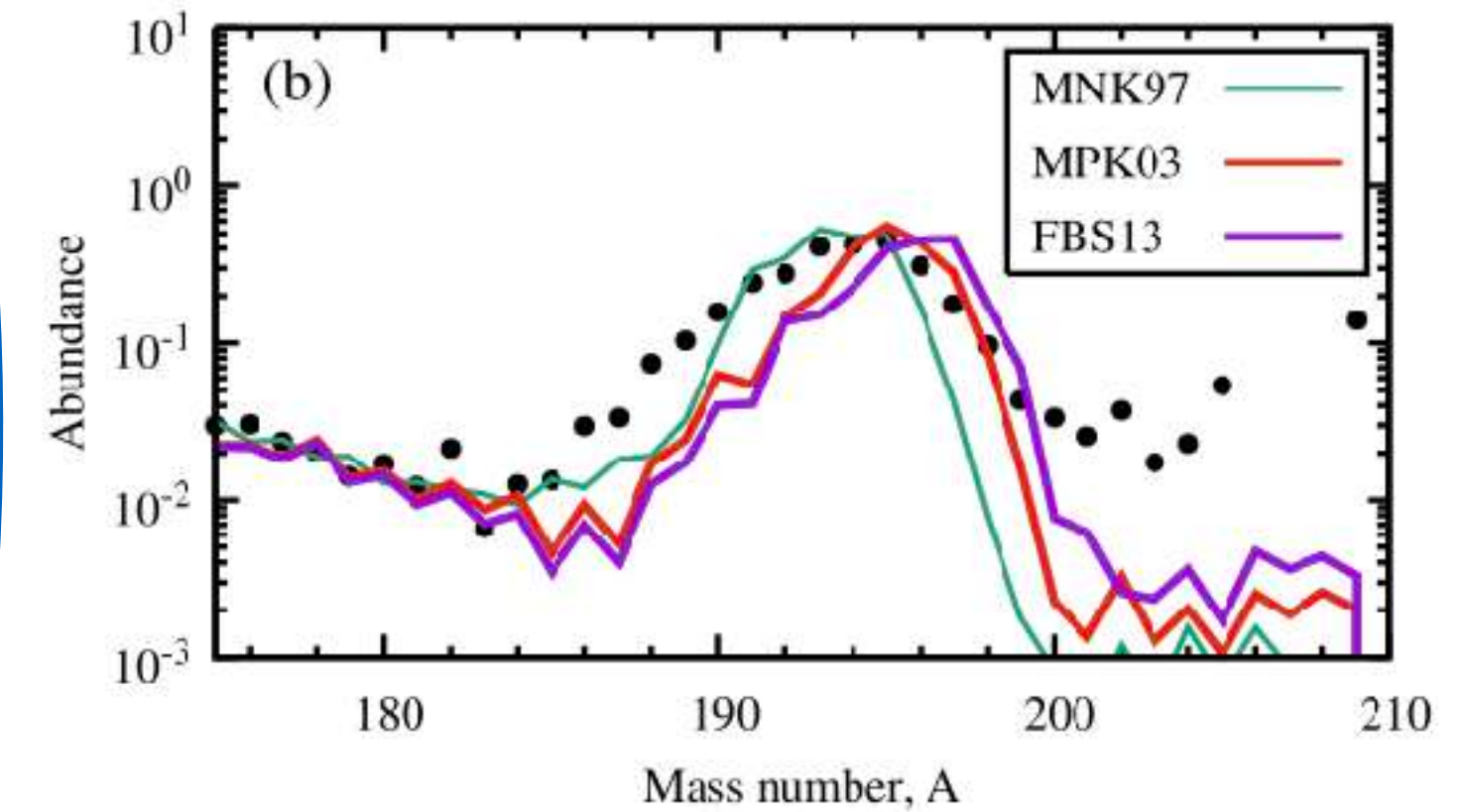


used by Kobayashi+2020



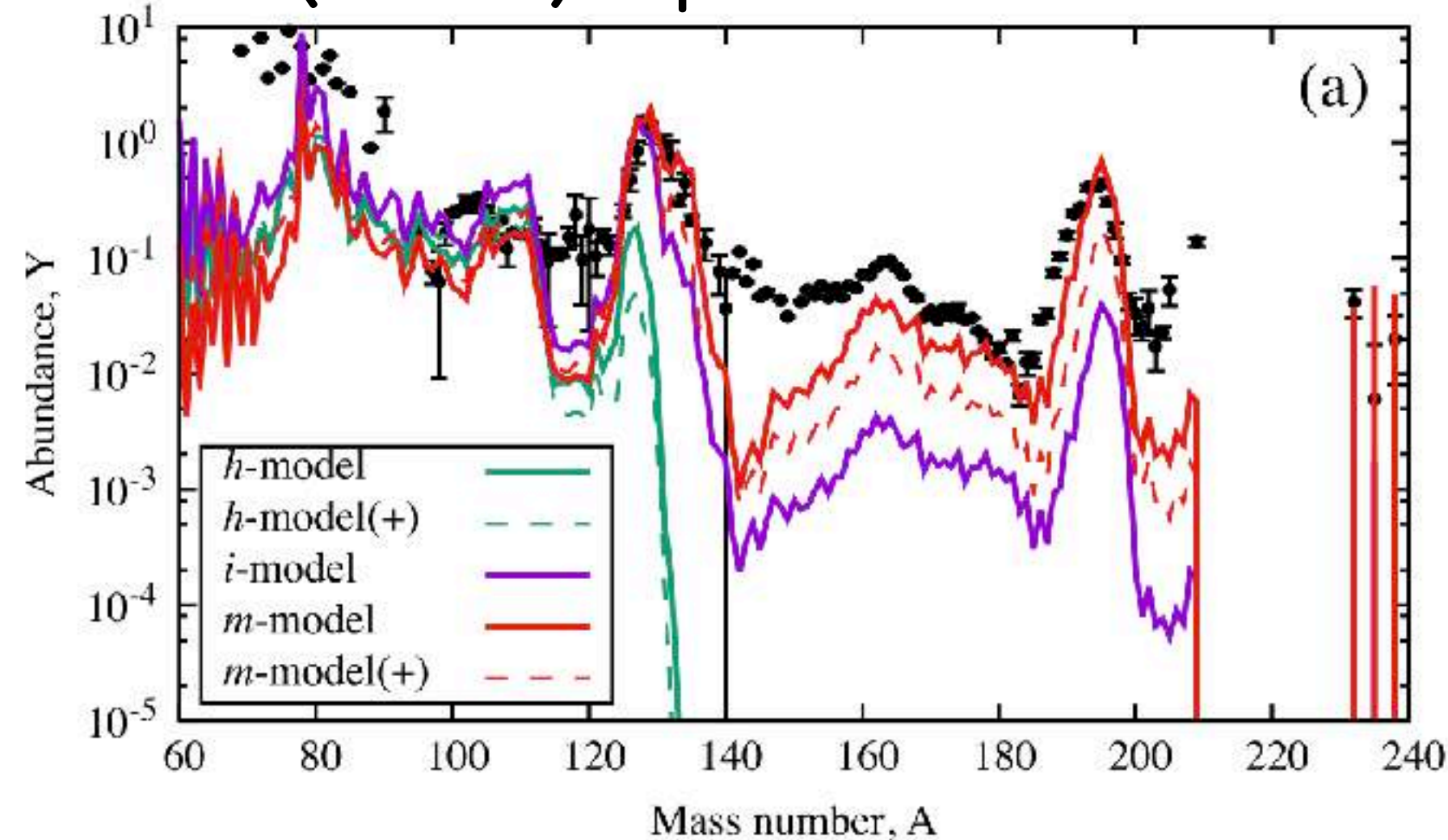
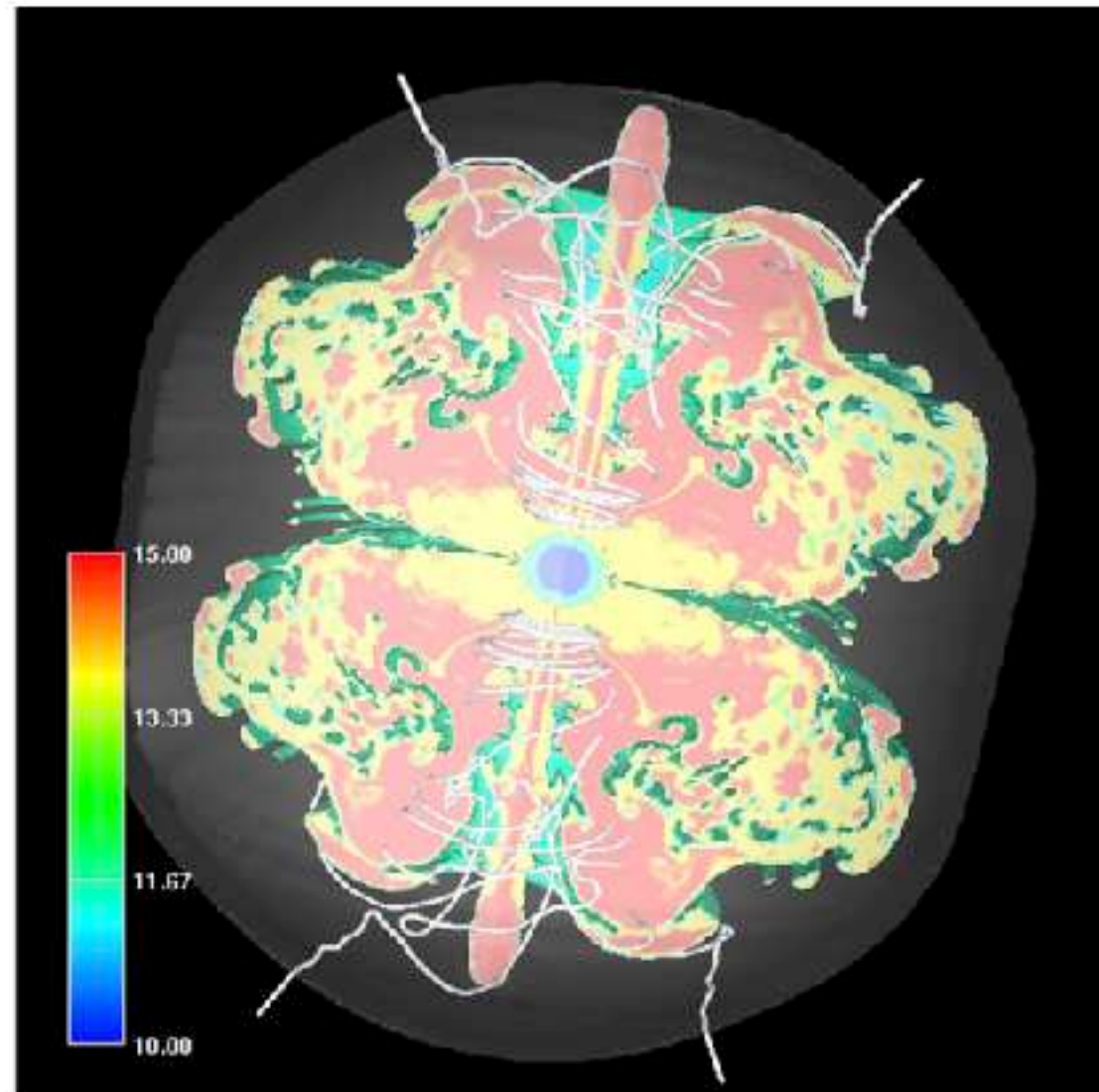
update my r-process calculations ( $\beta$ -decay)

NN+2016, PLB (2016)

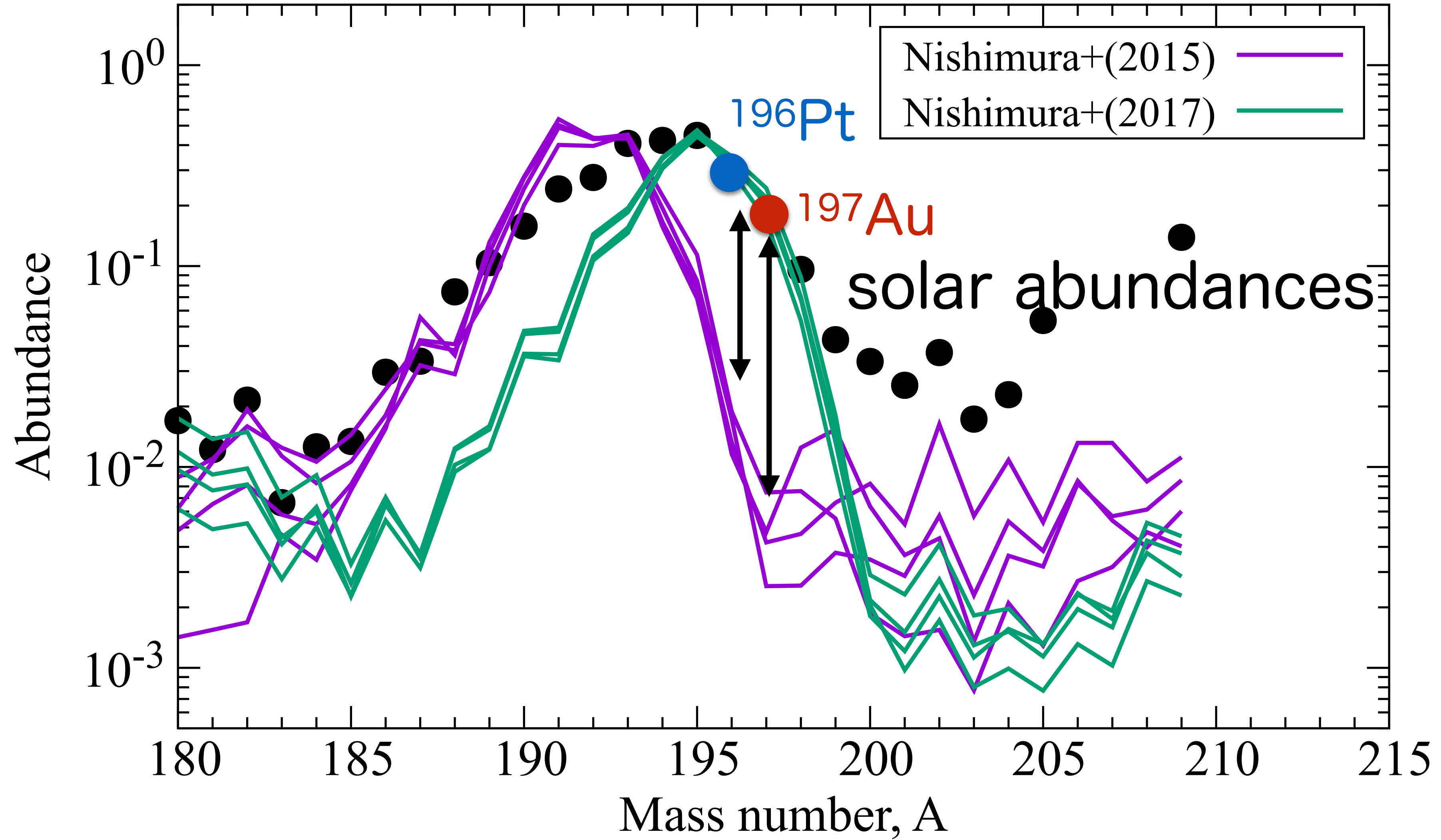


comparison of different beta-decay rates

Nishimura+(2017) ApJL



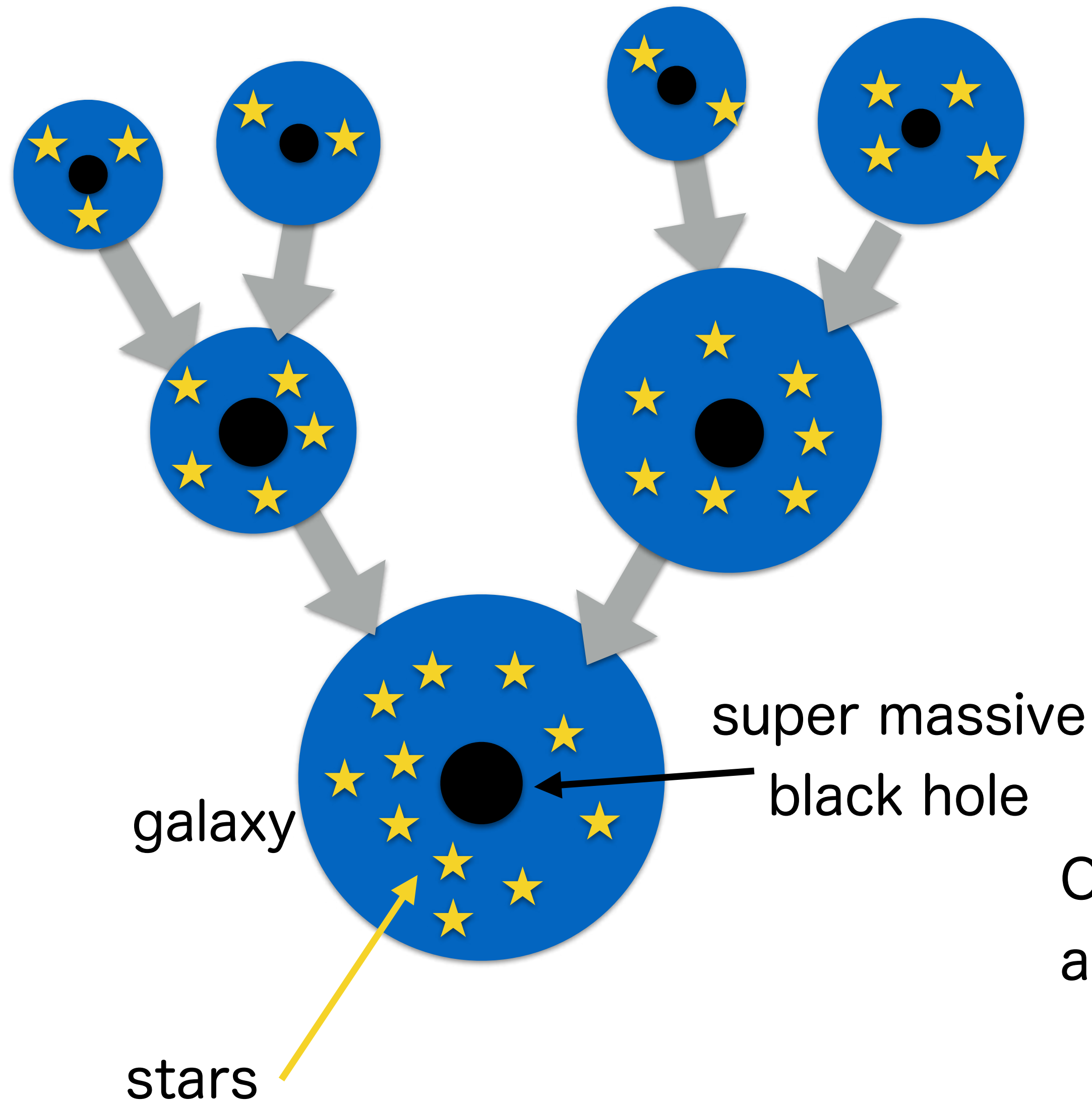
# Prequel story?



# Evolution of galaxy

galaxy formation simulation

Hierarchical structure formation



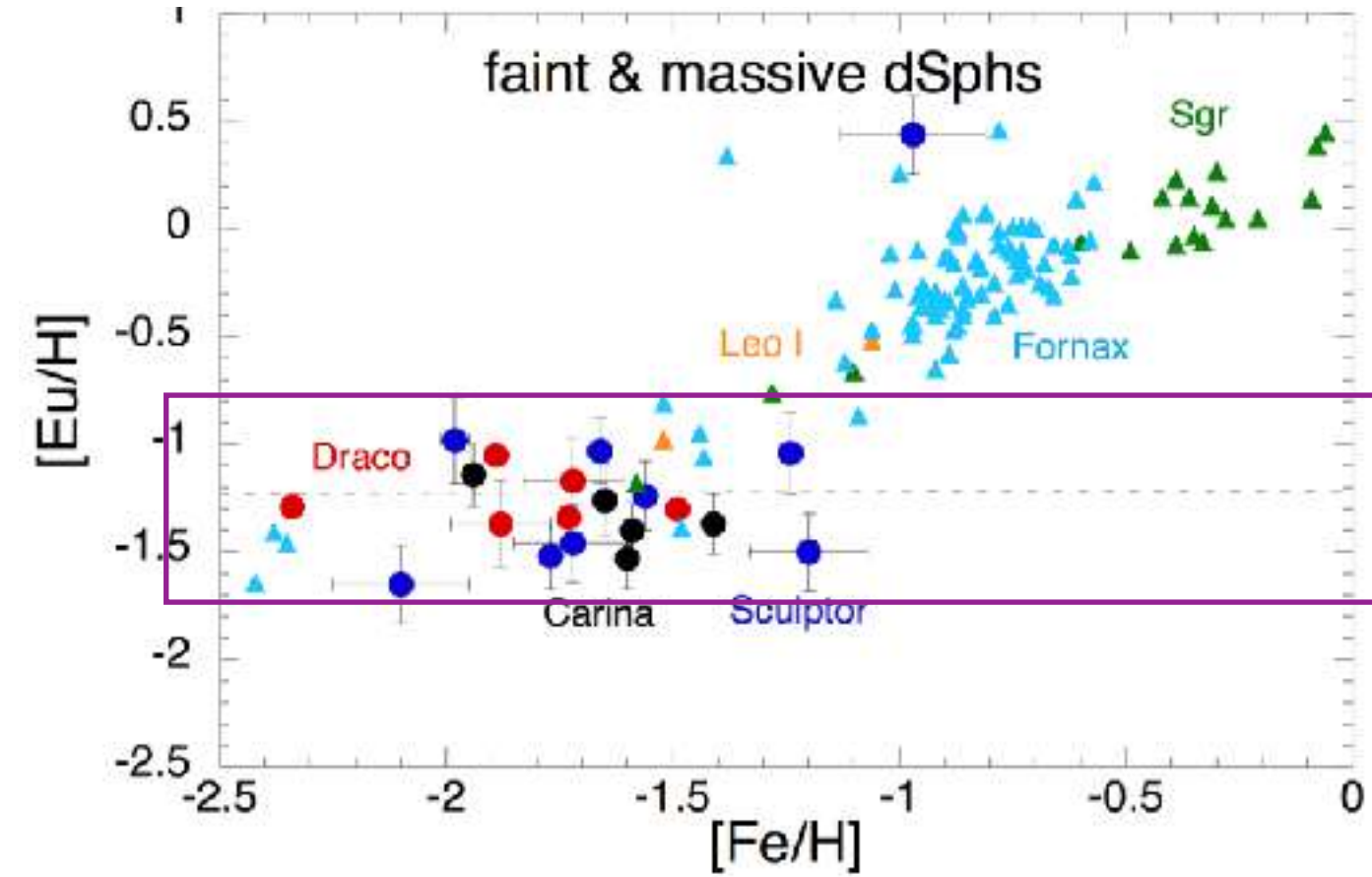
Takeda, Nukatani, Saito (2007)

Our Galaxy (Milky Way) was merged with a dwarf galaxy (Gaia-Enceladus) 10 billion yrs ago

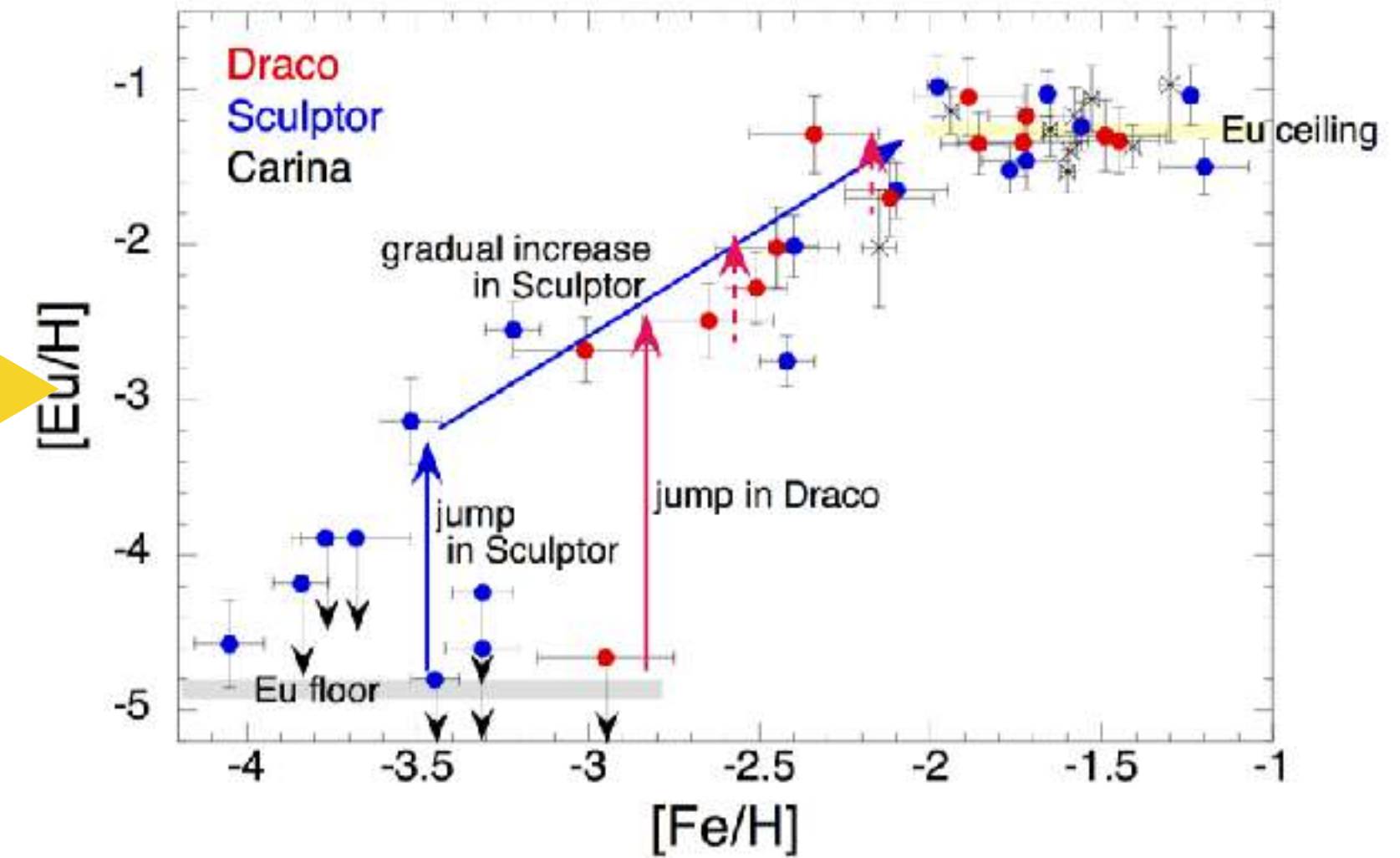
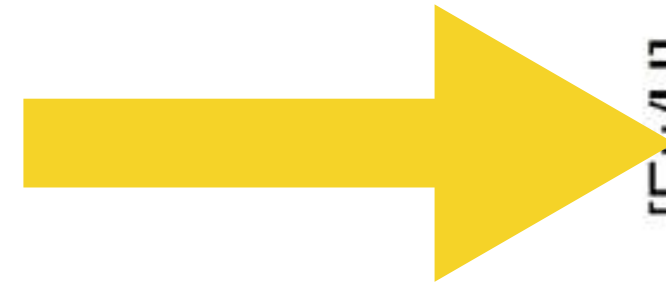
# GCE: early dwarf spheroidal galaxies

Tsujimoto & NN ApJL (2015)  
(see also, Tsujimoto+ PASJ 2015)

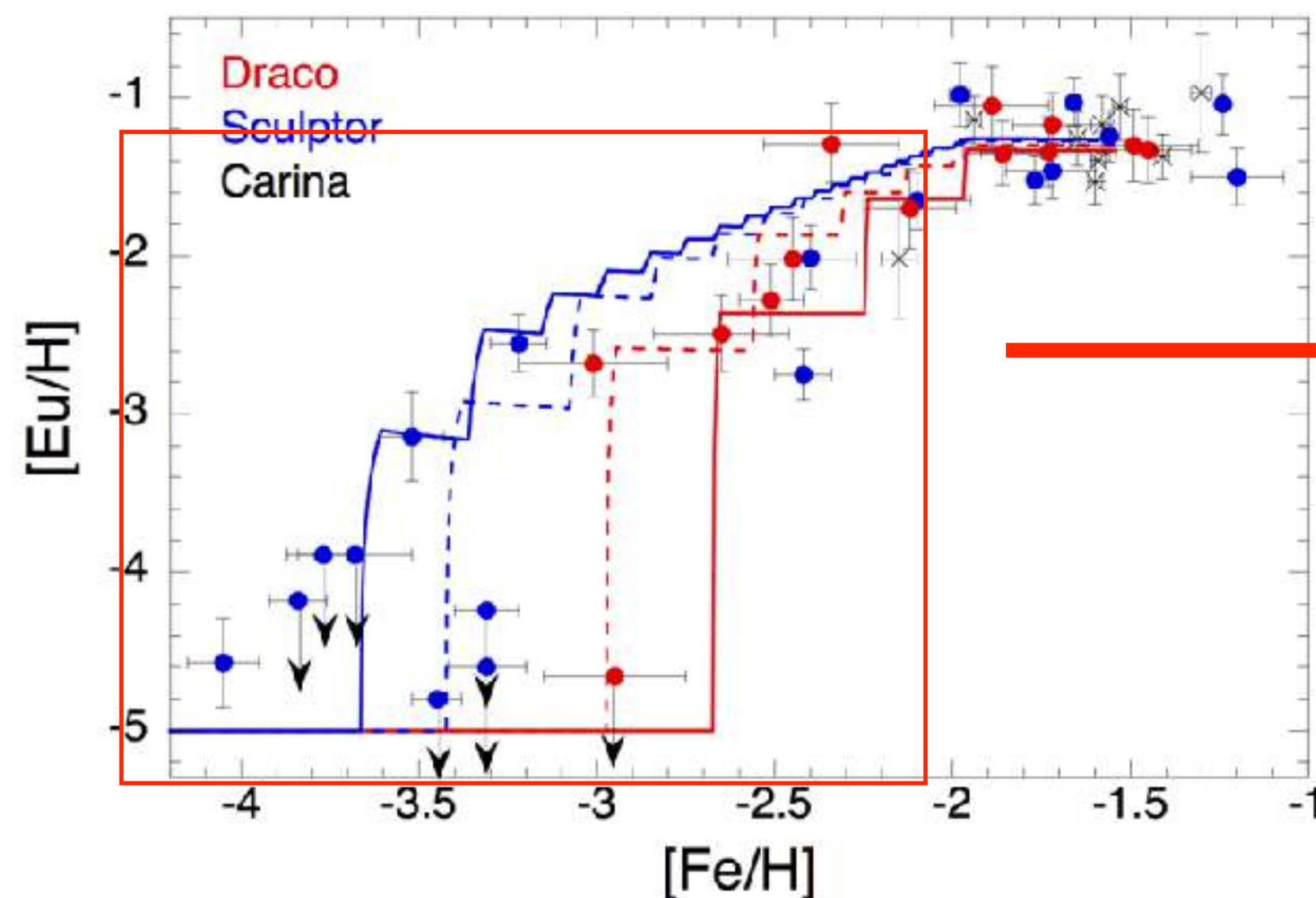
Tsujimoto & Shigeyama A&A (2014)



[Fe/H] < -2.5



Tsujimoto & NN, ApJL (2015)



Chemical evolution models

GCE models suggest:

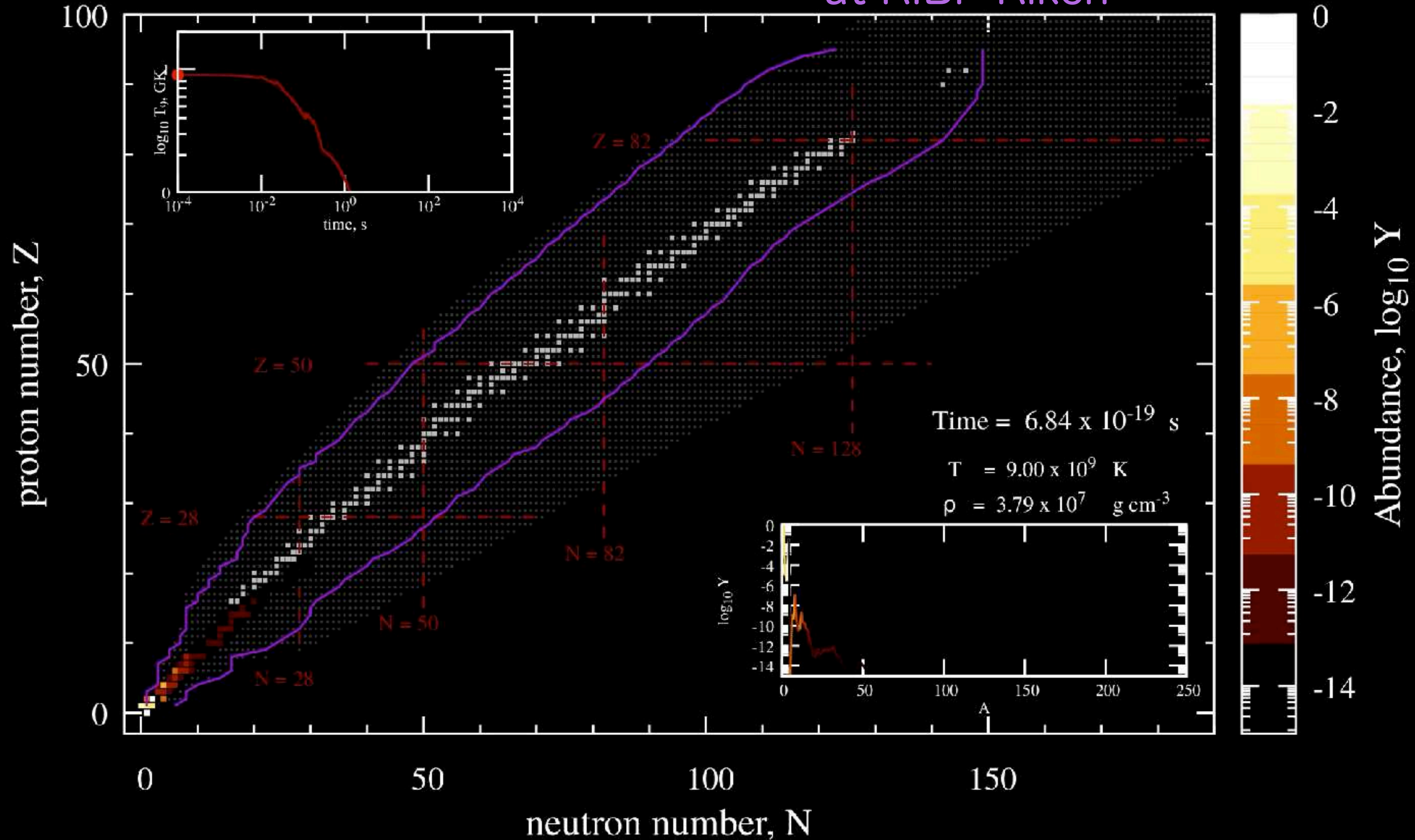
- rate event: 1/200 CC-SNe
  - large Eu ejection:  $\sim 10^{-5}$  Msun
- agree with our MR-SN models  
(e.g. Nishimura+ 2015)

## **2. Nucleosynthesis part:**

**N = 126 isotones and production of  
the r-process 3rd peak**

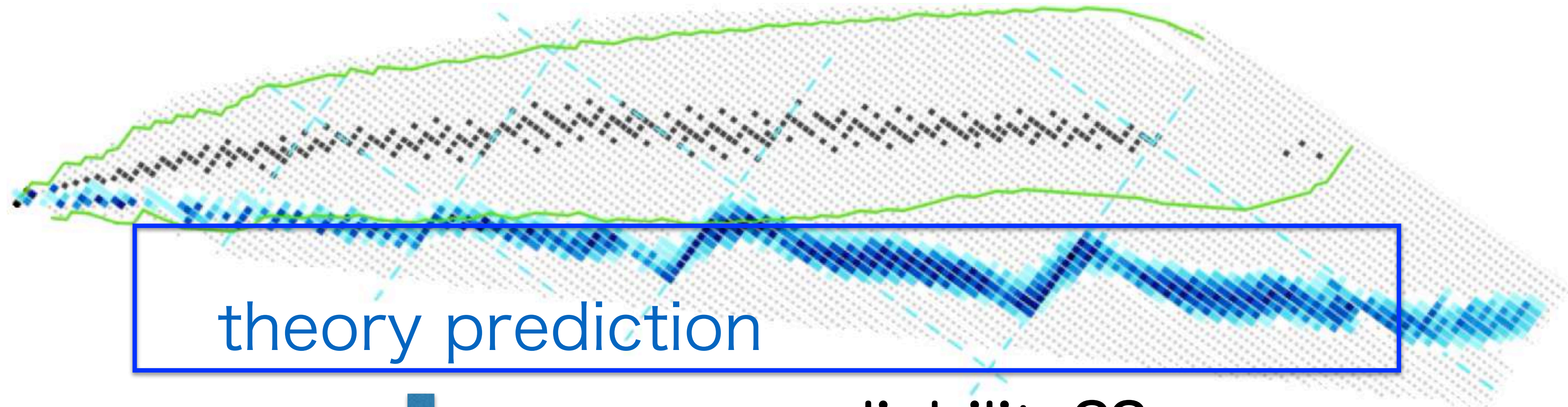
# r-process nucleosynthesis “flow”

1e-4 pps limit  
at RIBF Riken



# Theoretical Prediction

r-process path is beyond experimental accessible region

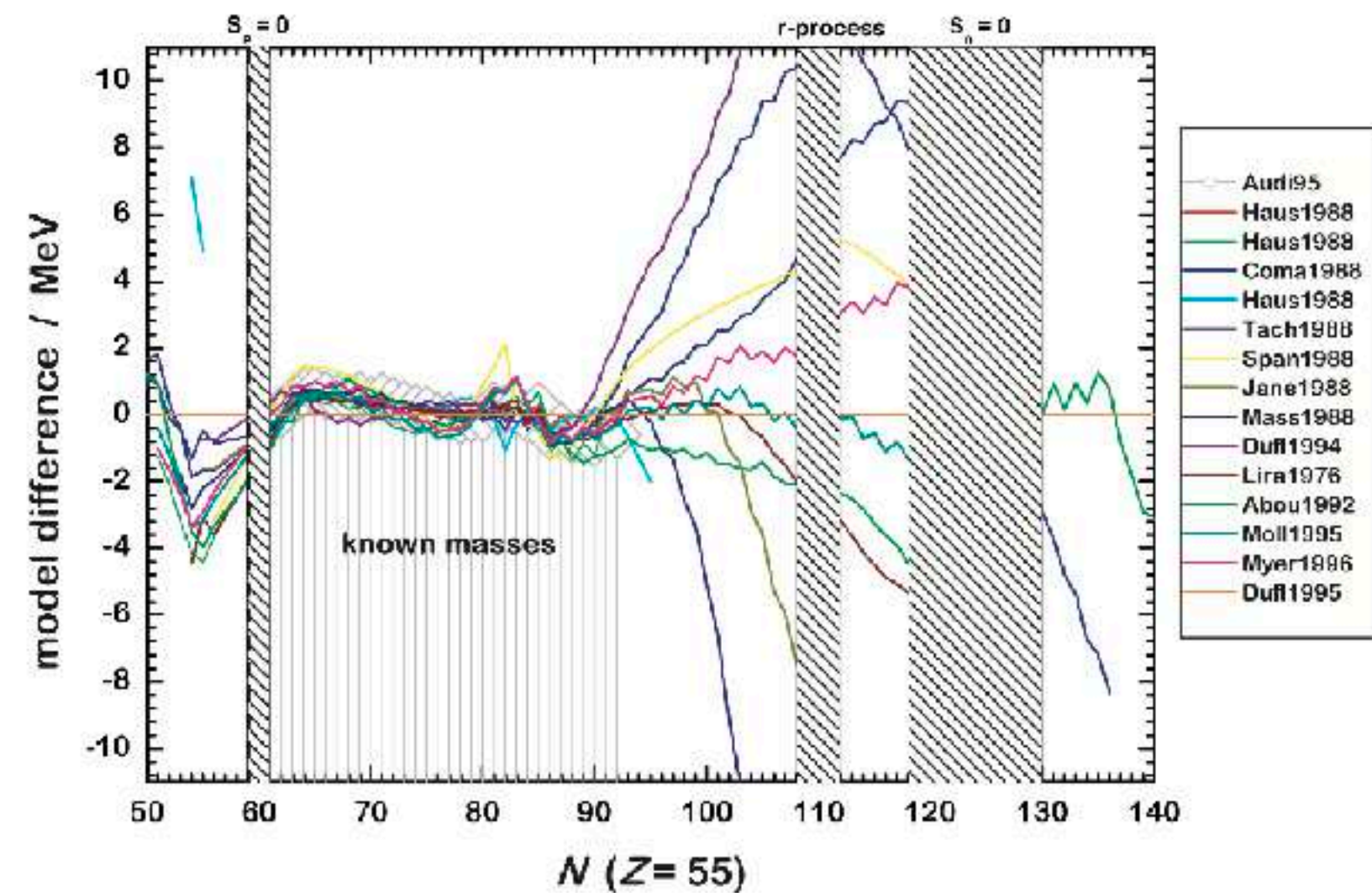
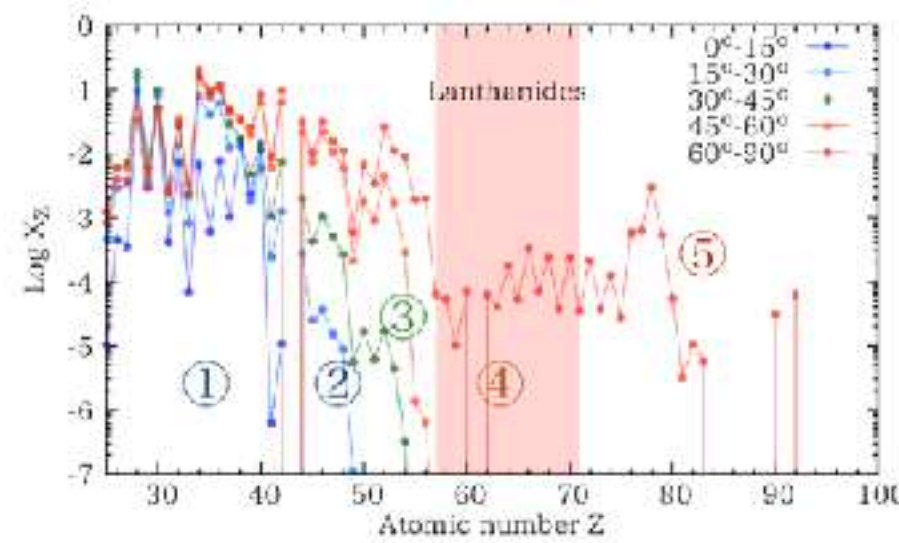
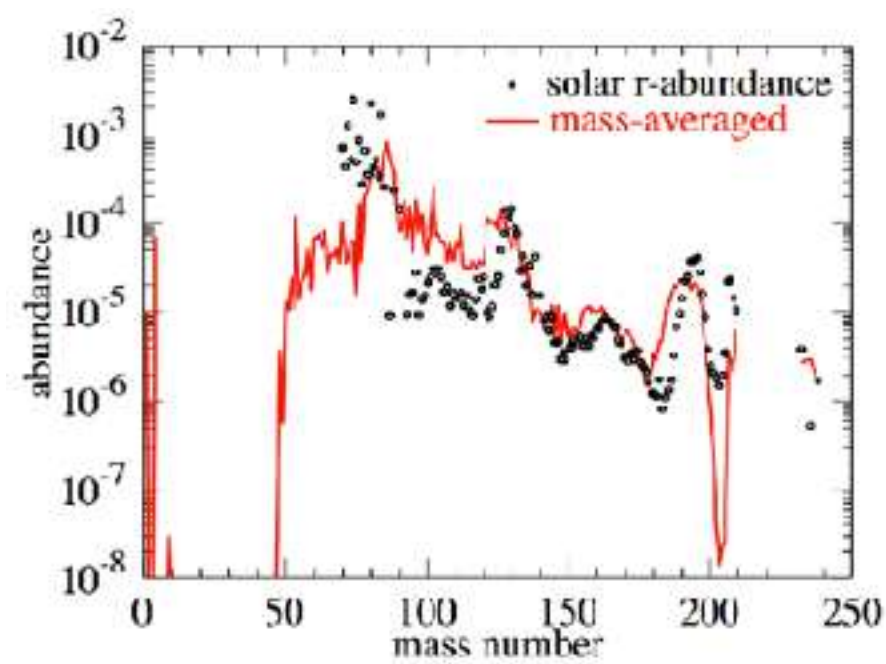


astrophysics



reliability??

large uncertainty in theory



# Basics of the r-process

## • nuclear physics

• “classical” solution :  $(n, \gamma)/(\gamma, n)$  equilibrium, waiting-point approx.

① r-process path :  $(n, \gamma)/(\gamma, n)$  equilibrium  $\rightarrow$  nuclear mass

② r-process abundance : half life (ratio) on the path  $\rightarrow$   $\beta$  decay  
(abundance) / ( $\beta$ -decay half life) = constant

③ decay : smoothing by n-emission  $\rightarrow$   $\beta$ -delayed n emission  
(fission of heavy nuclei : cannot consider?)  $\rightarrow$  fission

① reverse reaction (photodissociation) is determined by the detailed balance :

$$\lambda_{\gamma}(Z, N) \propto T^{3/2} \exp \left[ -\frac{S_n(Z, N)}{kT} \right] \frac{\langle \sigma v \rangle_{(Z, N-1)}}{E}$$

neutron separation E                      neutron capture rate

r-process path =  $(n, \gamma)$  and  $(\gamma, n)$  balance

② abundance

determined by the ratio of  $\beta$ -decay rate (half life) on the path;  
(in contrast to the s-process determined by  $(n, \gamma)$  rate )



# Modern nucleosynthesis simulations

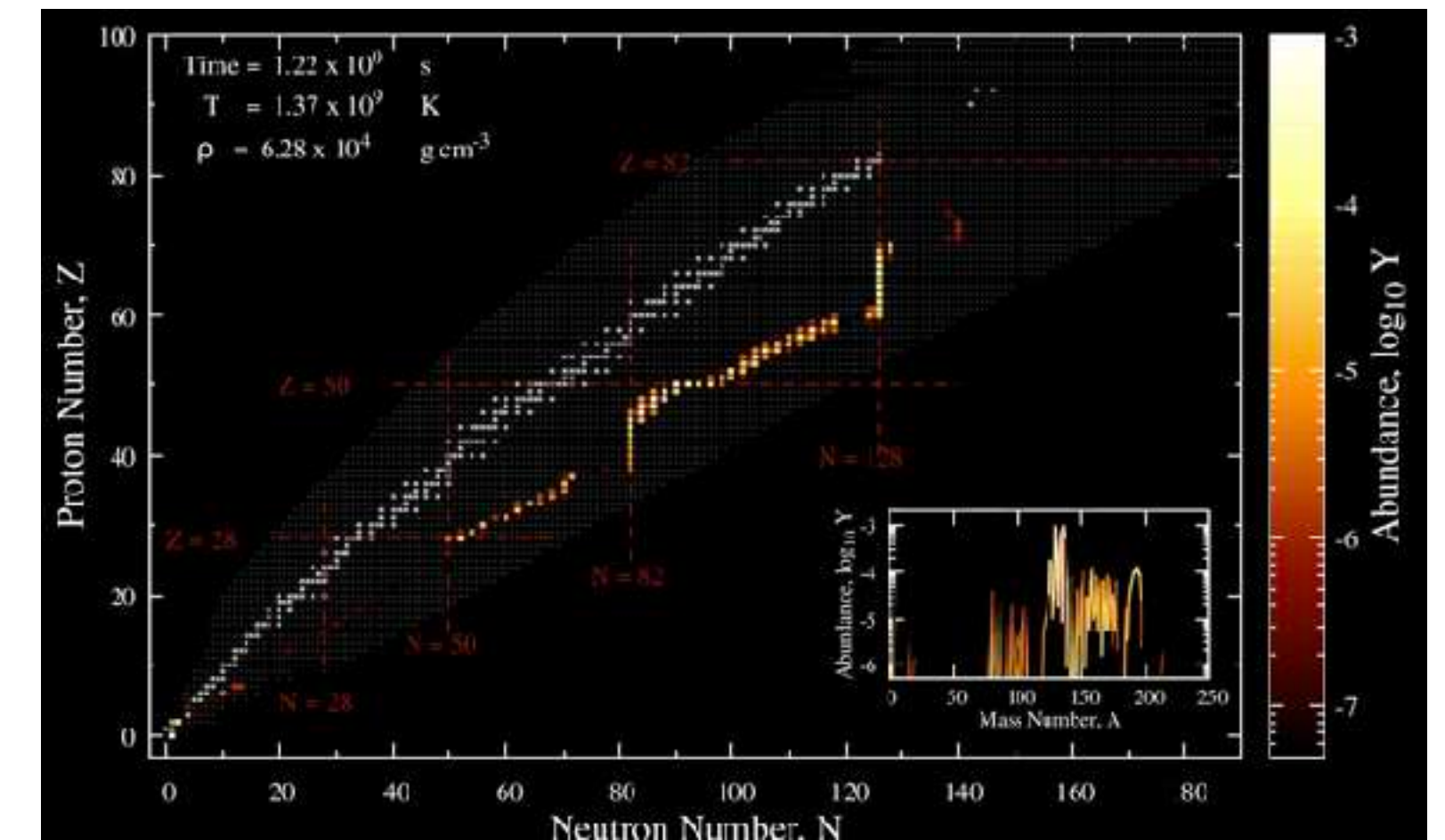
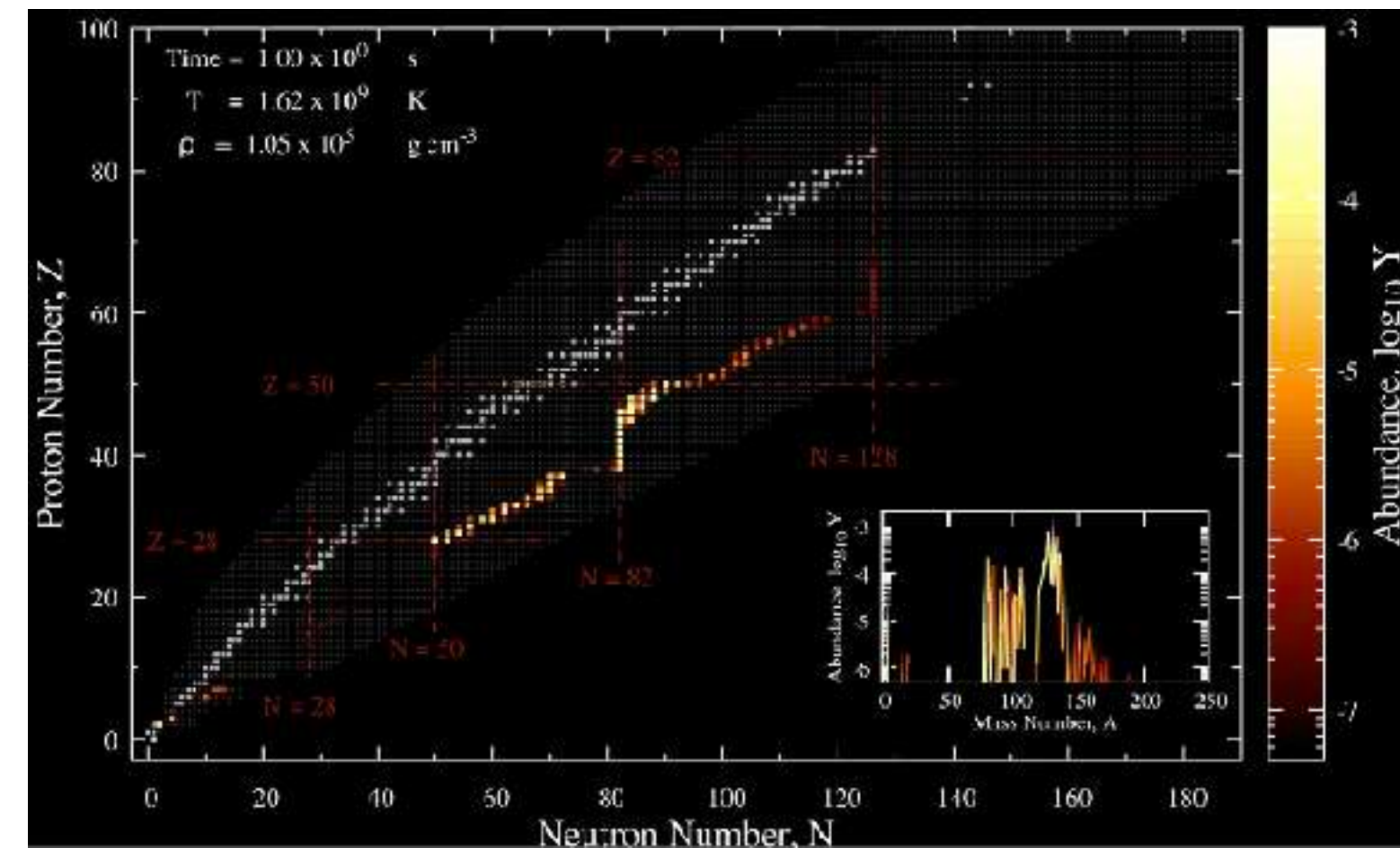
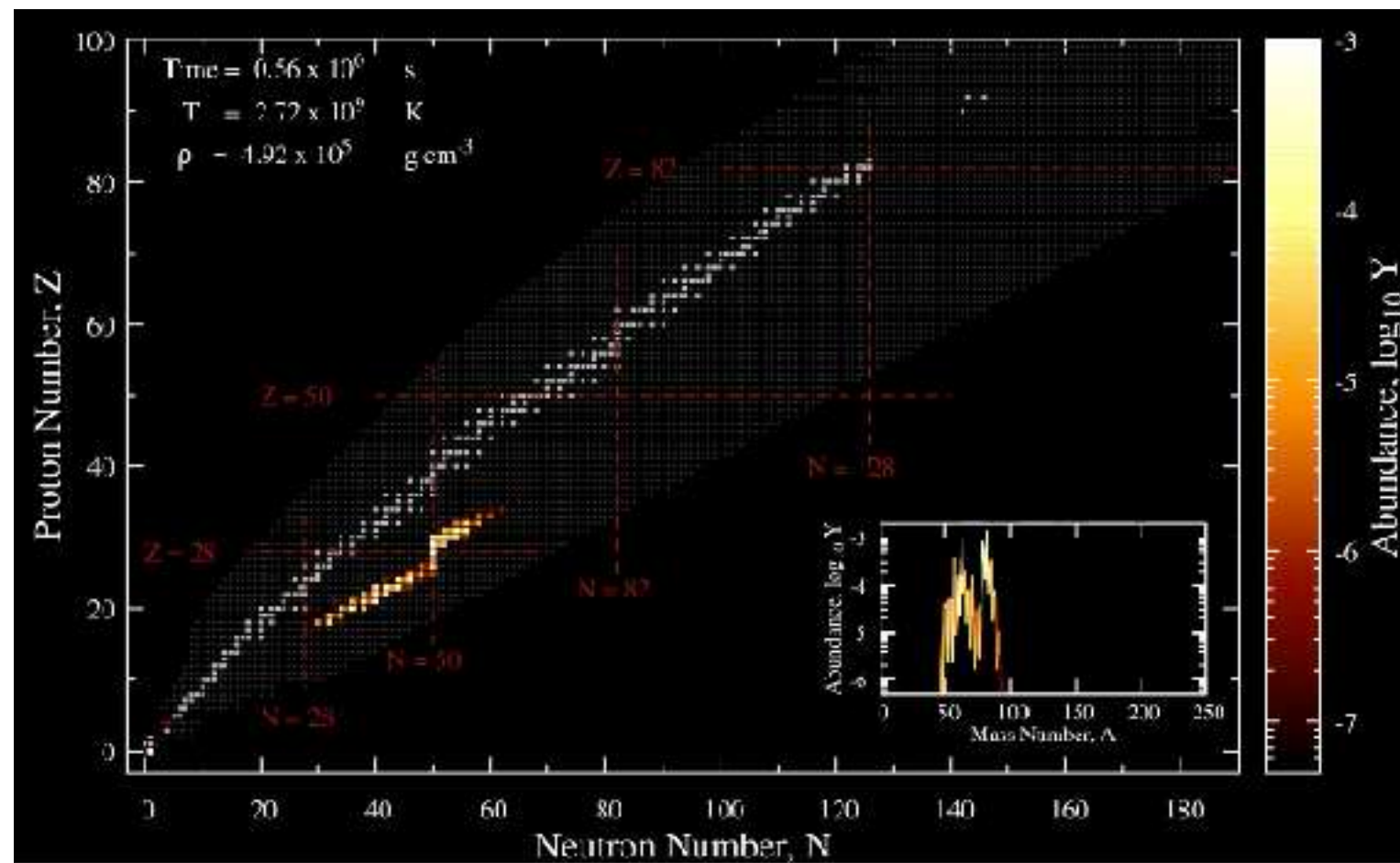
## • Nuclear physics

### • method : calculations by nuclear reaction network

• consider all possible reactions and decays

• (n,g)/(g,n) partially achieved or not realized (cold r-process)

•  $\beta$ -decay, (n,g), fission for a wide range of n-rich nuclei



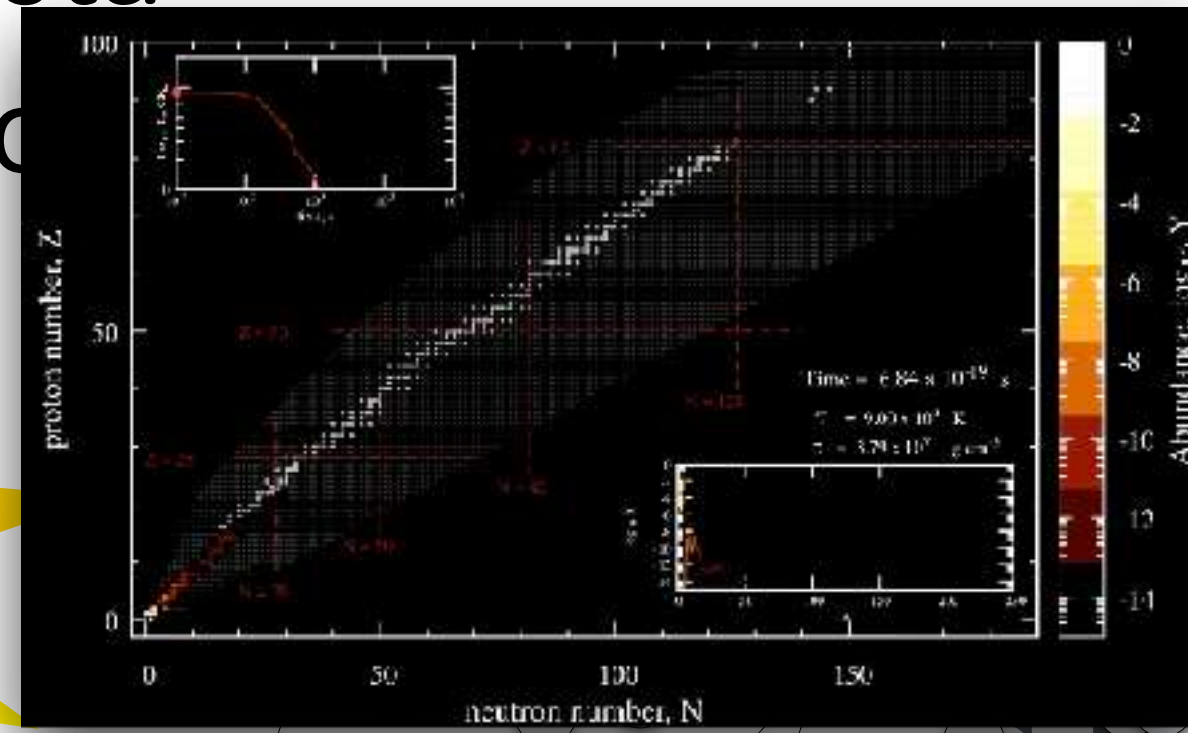
# r-process: nucleosynthesis mechanism

neutron-rich ejecta  
in stellar explosion

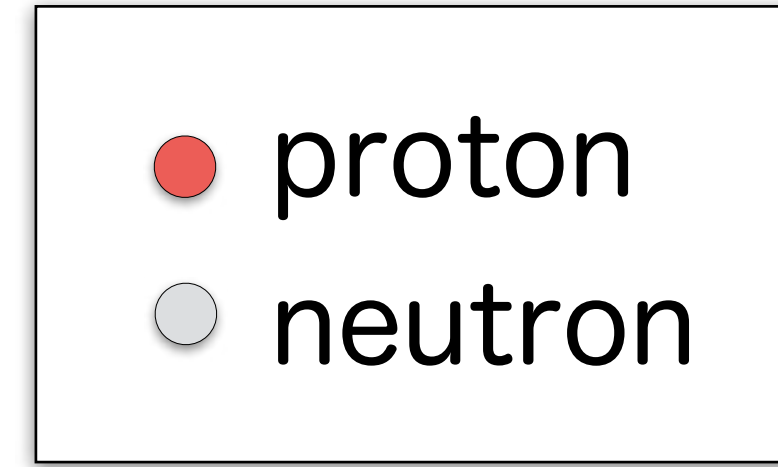


high  $T$  &  $\rho$

① NSE ( $> 1$  MeV)

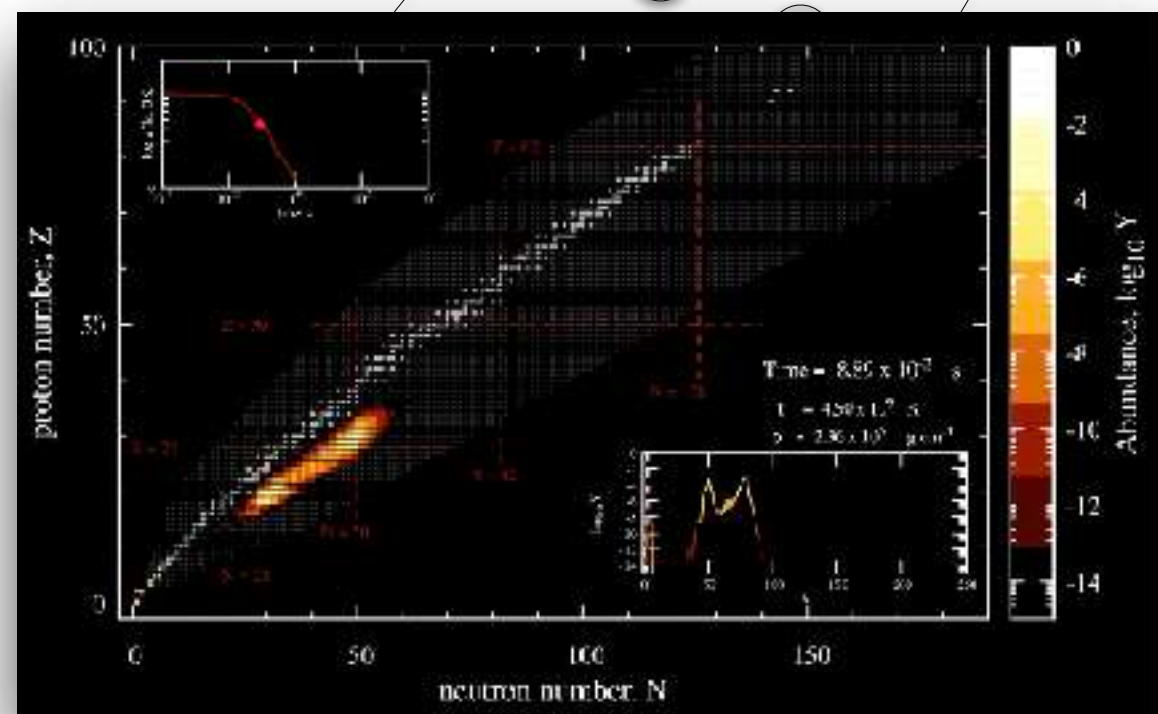


d  $\alpha$ )



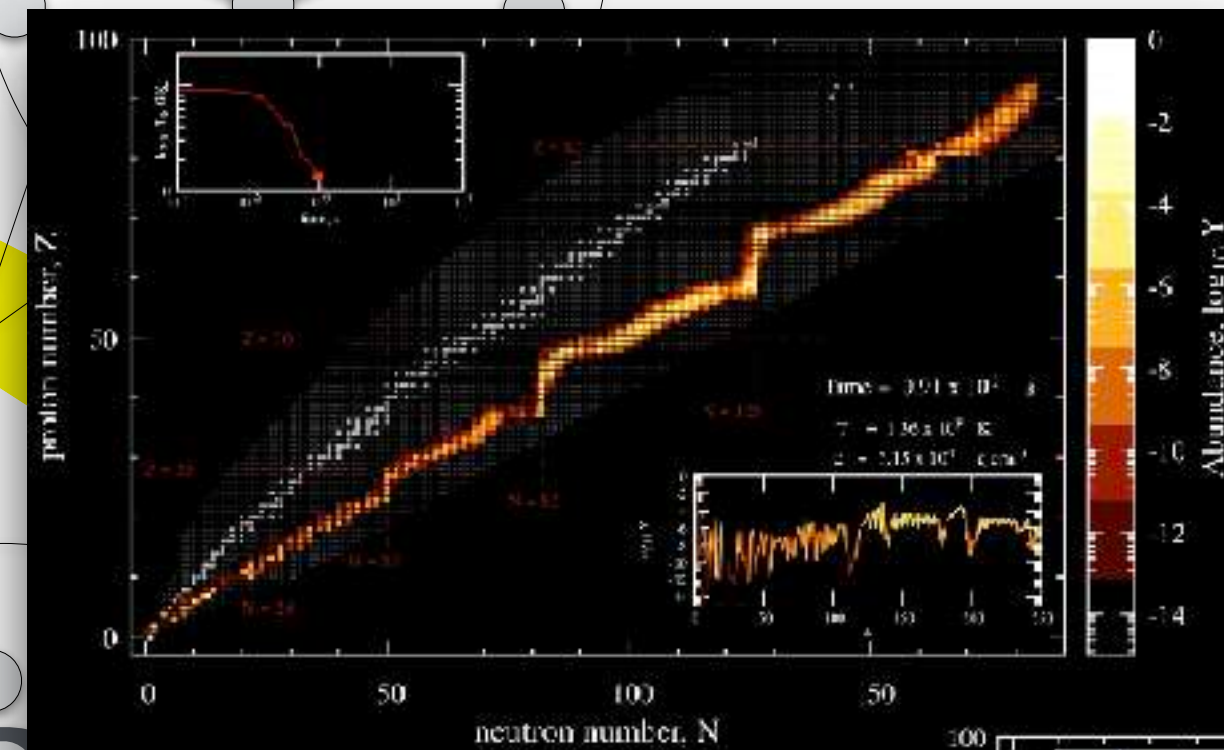
③ n-capture:  $A \uparrow$   
 $(A, Z) \rightarrow (A+1, Z)$

⑤  $\beta$ -decay  
+ fission

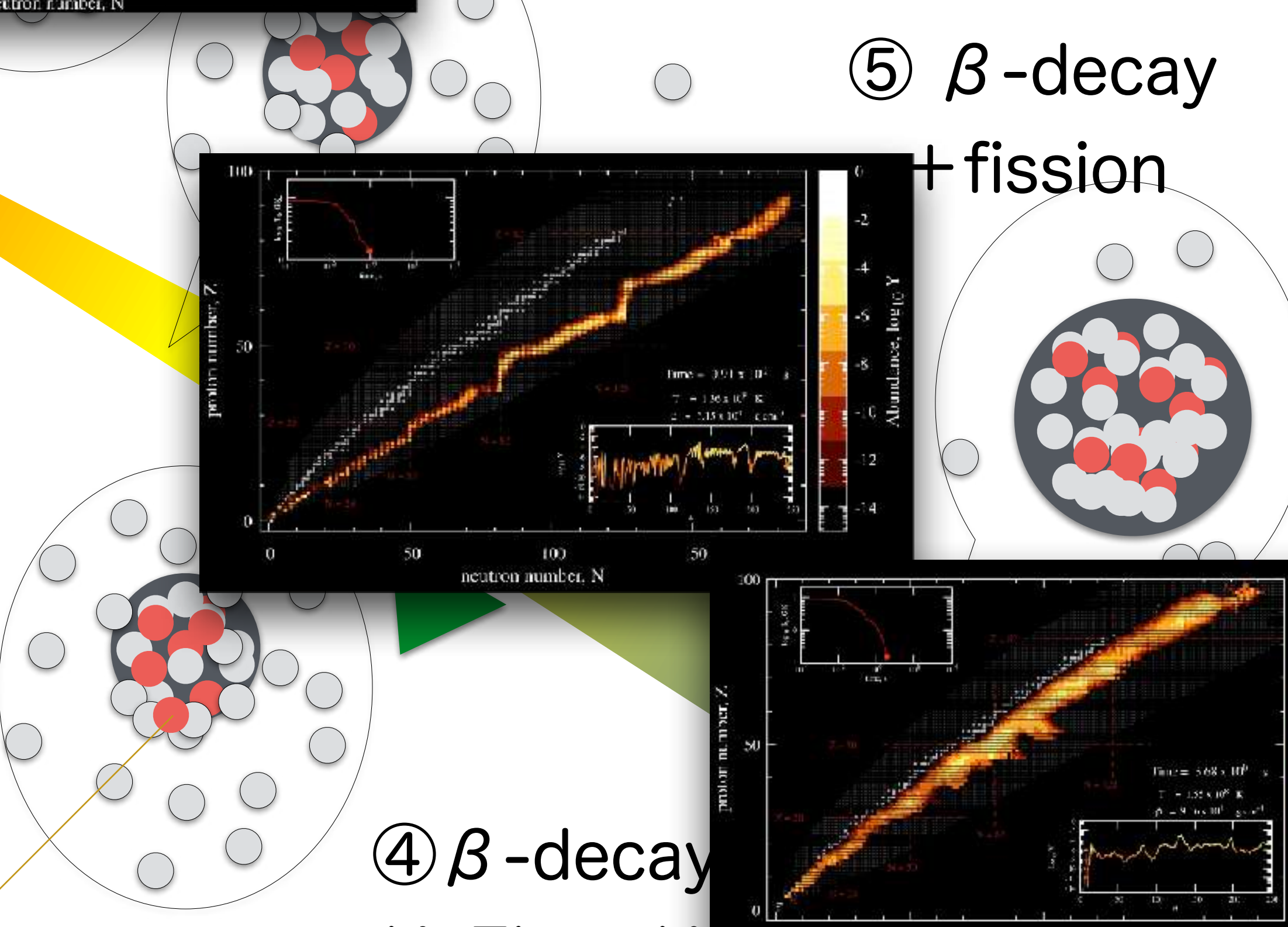
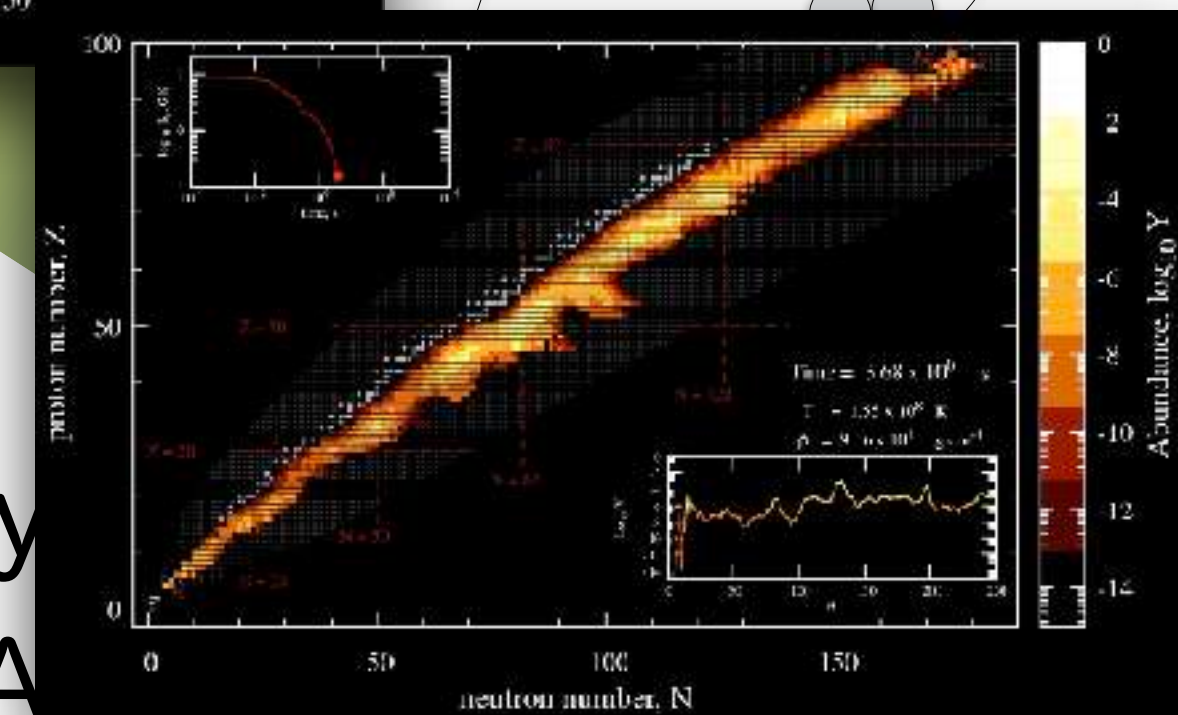


② seed formation

( $\alpha$  + recombination)  
 $A \sim 100$

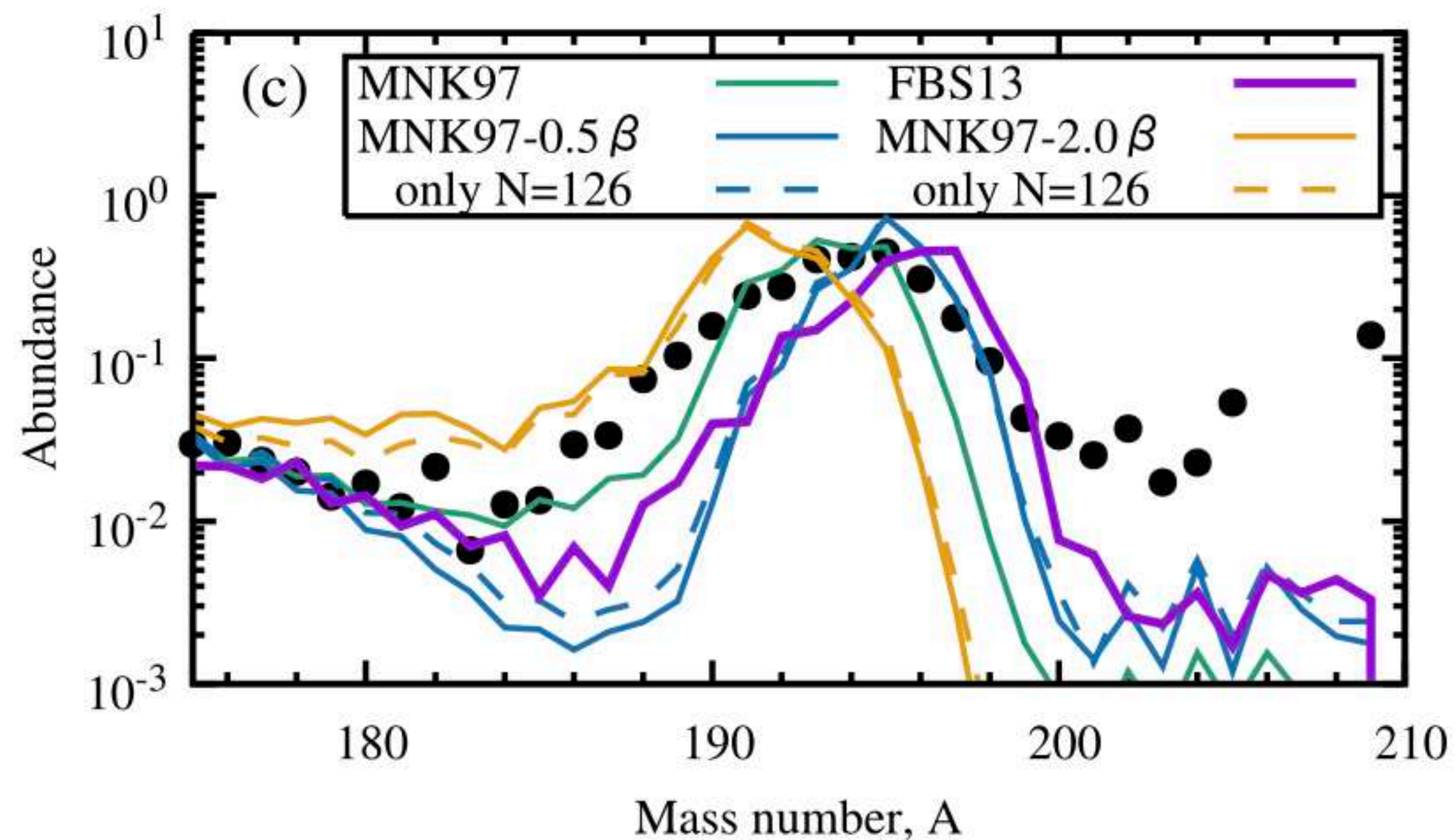


④  $\beta$ -decay  
 $(A, Z) \rightarrow (A, Z-1)$

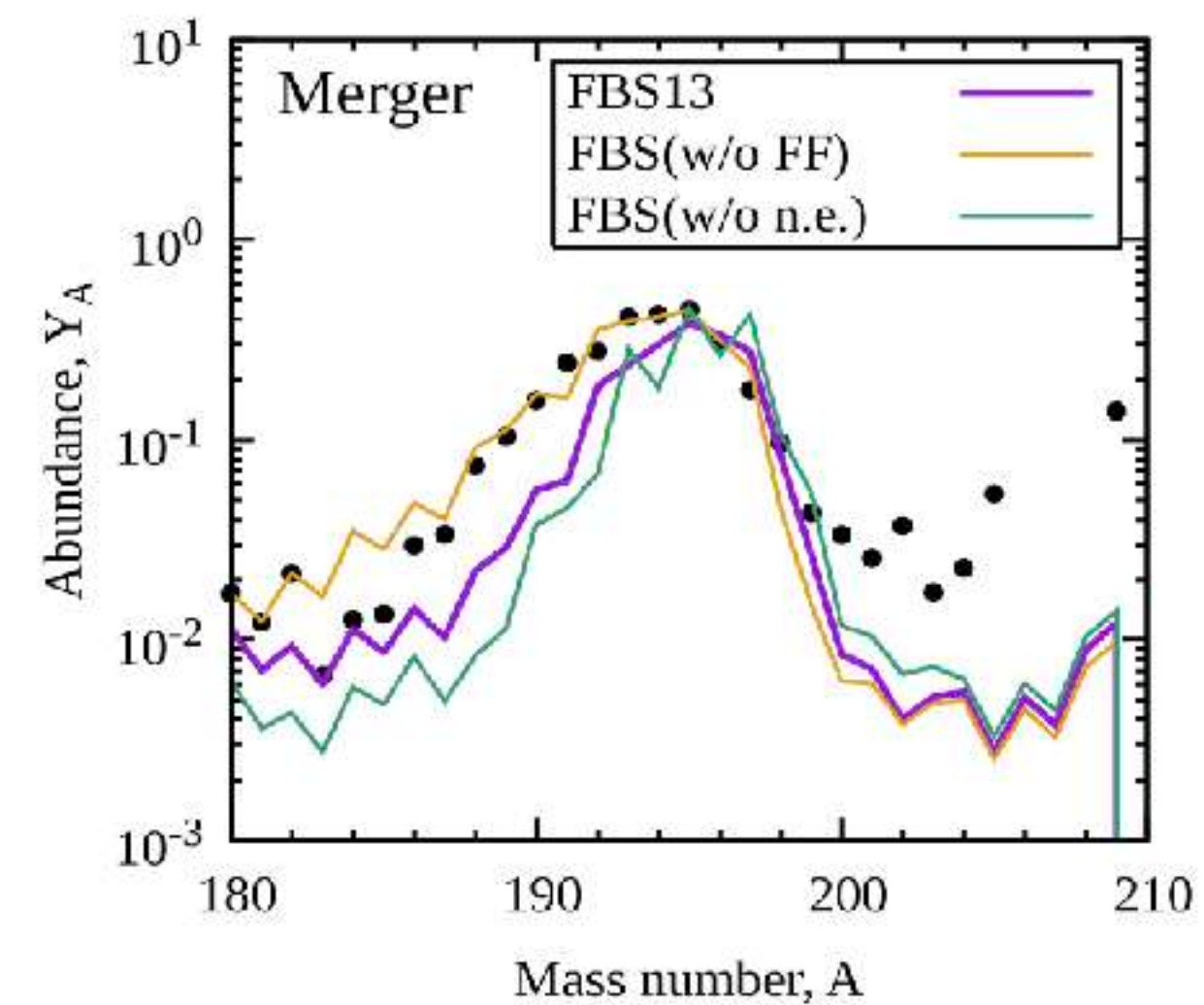
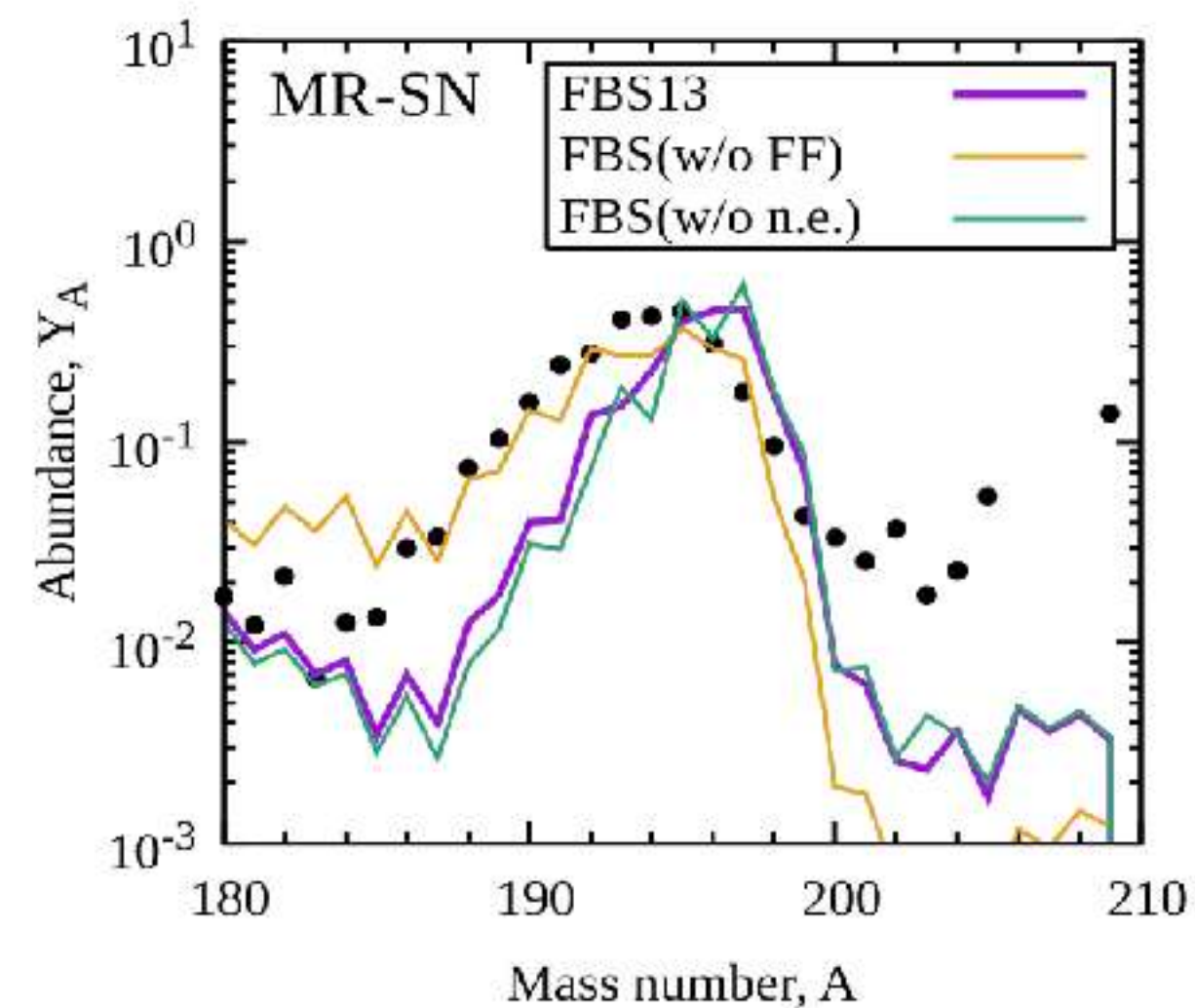


# $\beta$ -decay impacts on the 3rd peak

NN+2016



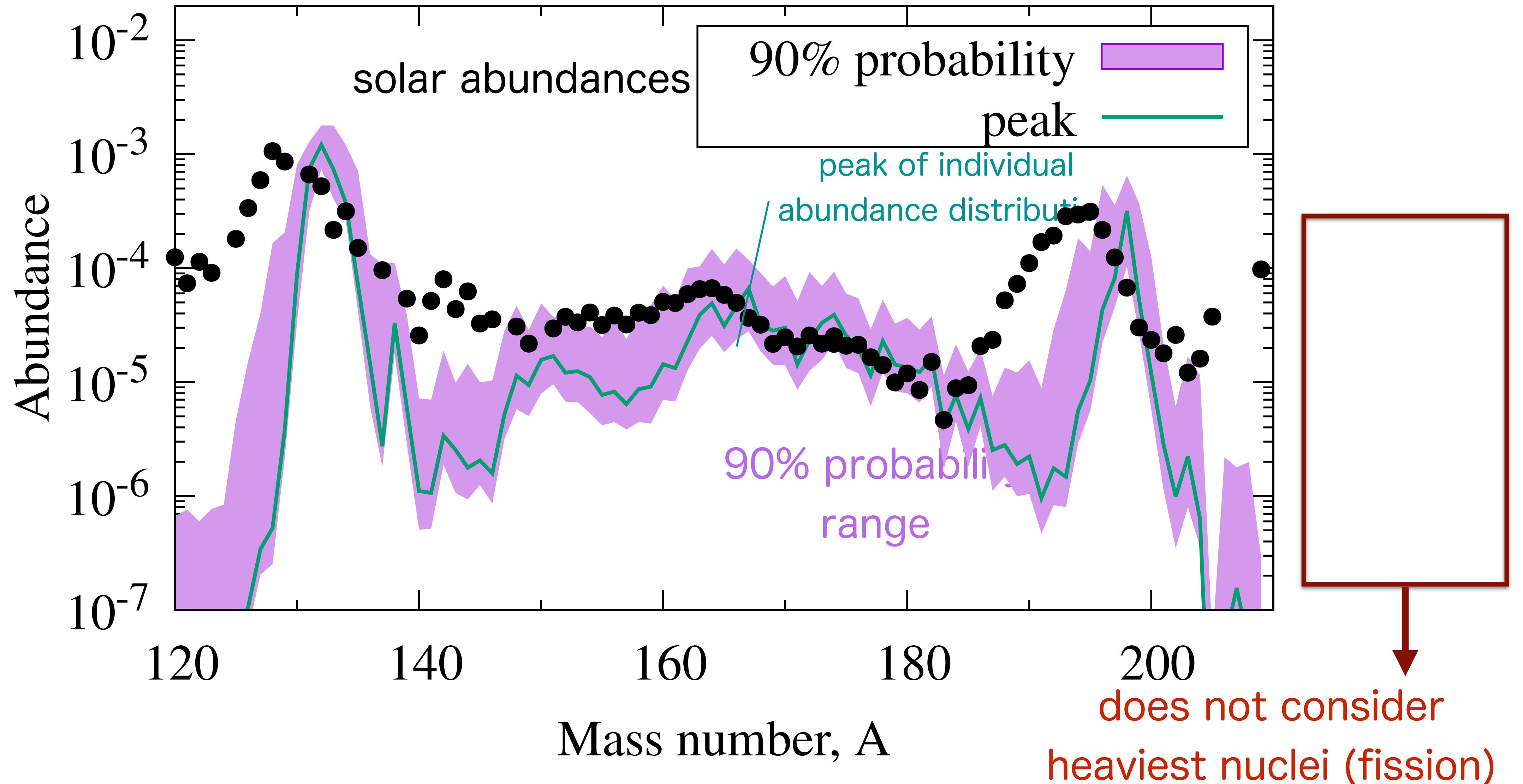
different astrophysics models



# Collective uncertainties on the r-process

impacts on reaction/decay rate variation with MC nucleosynthesis

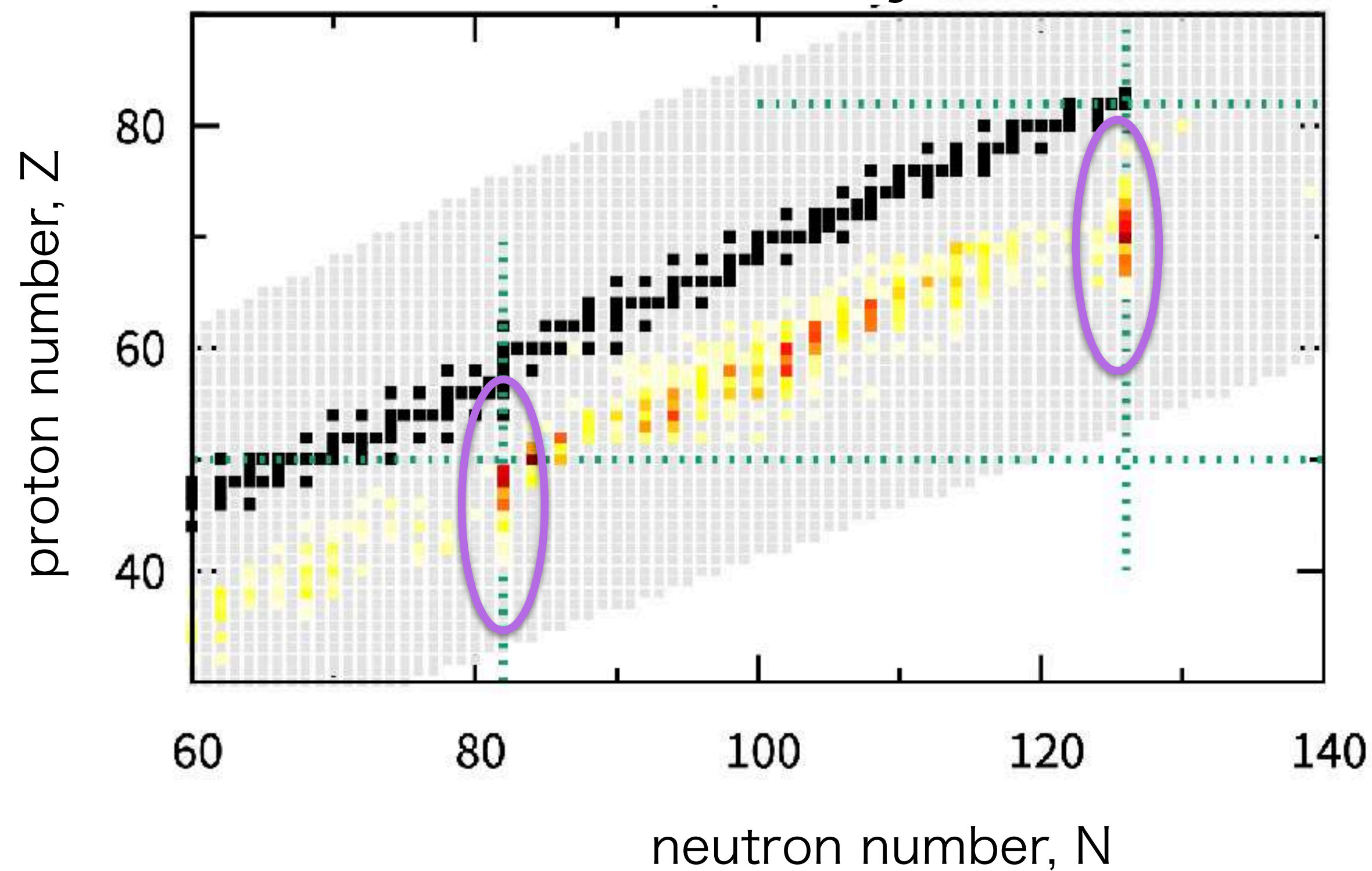
(n,g) x50,  $\beta$  x10



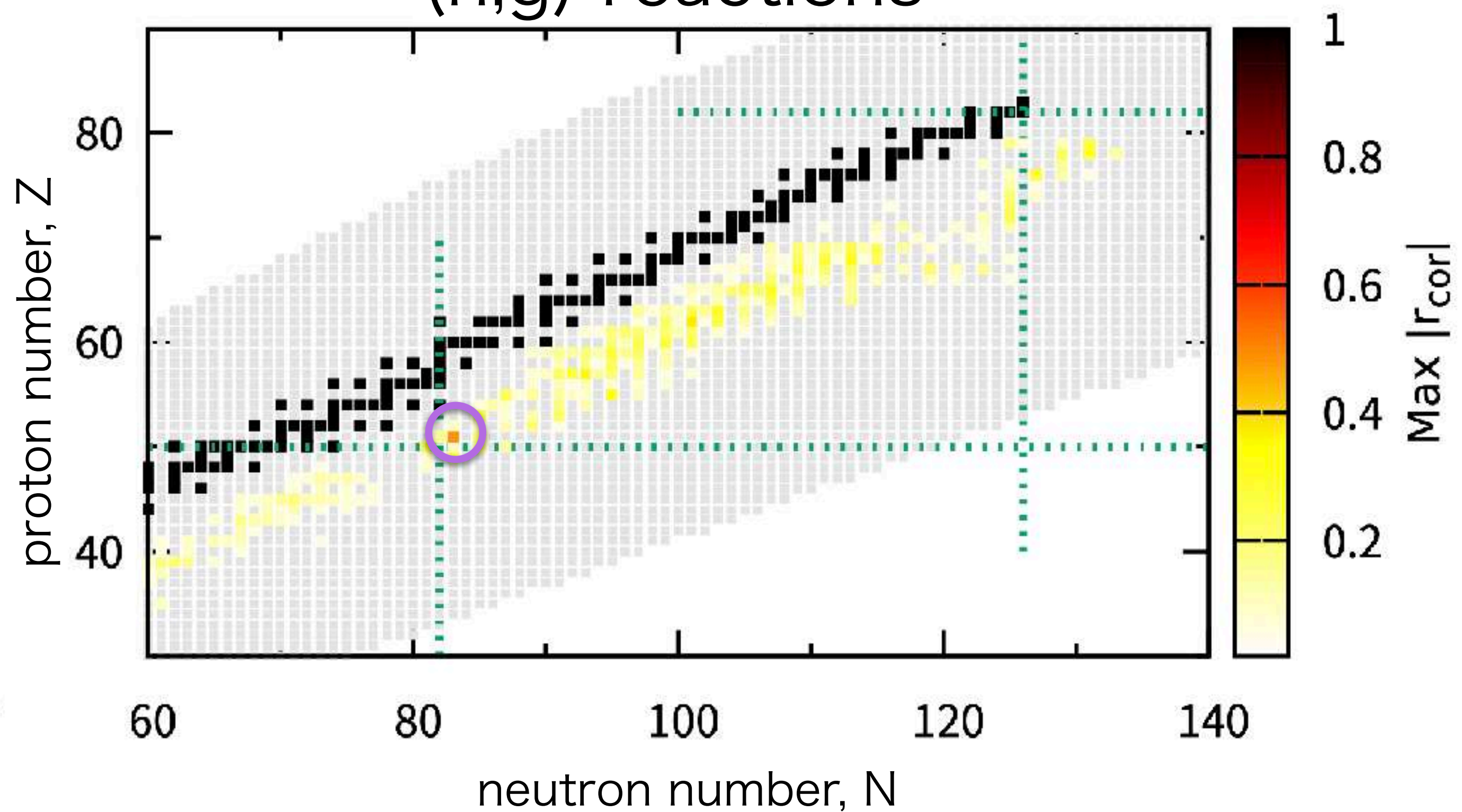
# Individual impacts

Impacts on the r-process of individual rates

$\beta$ -decay



(n,g) reactions



# Summary and perspective(personal)

## 1. “The missing gold problem” in GCE

- It may be caused by “insufficient” theoretical prediction
- Though variation/differences by astrophysical parameters, significant underproduction of Au (gold) in some particular GCE studies would be improved by r-process prediction
- Anyway, Galactic chemical evolutions studies are difficult...

## 2. N = 126 halflives and the 3rd peak formation

- $\beta$  decay of n-rich N = 126 isotones have significant impacts on the 3rd peak formation
- This is also (obviously) confirmed by comprehensive MC sensitivity calculation  
→ investigations by improved “theory” data is also expected

## 3. Future perspective (personal)

- Recent theoretical progresses in Japanese community (DFT studies, Minato-san, Anil etc)
- Personally, although I had already done, it would be worth to revisit it now.
- I need to clear up any confusion (“the missing gold problem”) in the GCE study.